

Euclid

Structural-Thermal-Optical Performance

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ABSTRACT

Context. The *Euclid* system performance is defined in terms of image quality metrics tuned to the weak gravitational lensing cosmological probe. The weak lensing measurement induces stringent requirements on the shape and stability of the VIS instrument system point spread function (PSF). The PSF is affected by error contributions from the telescope, the focal plane and image motion, and is controlled by a global error budget with error allocations to each contributor.

ims. During spacecraft development, we verified through a structural-thermal-optical performance (STOP) analysis that the built and verified telescope with its spacecraft interface meets the in-orbit steady-state and transient image quality requirements, under temperature-induced loads, in all permitted spacecraft attitudes after all permitted attitude transitions. In the first year in orbit, we compared the expected with the actual performance.

ethods. For the purposes of the STOP analysis, a detailed finite-element mathematical model was set up and a standard set of test cases, both steady-state and transient, was defined, comprising combinations of worst-case boundary conditions. Iterations of the analysis were performed in conjunction with the major reviews of the spacecraft verification cycle. After launch, the model is used in sensitivity analyses using realistic boundary conditions.

esults. The STOP analysis addressed the interaction of all spacecraft components in transmitting temperature-induced loads that lead to optical strain deformation. The results of the prelaunch analysis demonstrated that temperature-induced optical perturbations will be well below the allowable limits for all permitted observing conditions. During the first year in orbit, we used the STOP analysis predictions to help interpret the measured performance as a function of environmental variables. Unpredicted disturbances were discovered (heat pulses from instrument operation propagating into the telescope) and unexpected sensitivities were revealed (high dependence of T_{BP} on Solar aspect angle (SAA), nearly absent dependence on α angle (AA) after the attitude domain was redefined for straylight avoidance). In-orbit temperature variations are small (<300 mK) and so are their effects on the telescope structure (displacements $< 1 \mu\text{m}$, rotations $< 1 \mu\text{rad}$), but they are detected in the time histories of the image quality metrics and are a non-negligible factor in the PSF stability budget demanded by the weak lensing science ($\Delta e < 2 \times 10^{-3}$ over 11 000 s). Taking everything into account, our analysis confirms the excellent overall performance of the telescope.

ey words. Space vehicles – Telescopes – Cosmology: observations

1. Introduction

Euclid is a space-based survey mission to improve our understanding of the nature of dark energy and dark matter with an unprecedented level of accuracy and precision (Euclid Collaboration: Mellier et al. 2025). The mission has been optimised for the measurement of two dark energy probes: weak gravitational lensing (WL) and galaxy clustering (GC). These complementary probes involve a statistical analysis of galaxies detected in *Euclid*'s survey area of $14\,000 \text{ deg}^2$, which covers a major fraction of the extra-galactic sky that can be observed in the visible and near-infrared. To meet the high-precision cosmology objectives, a top-down design approach was adopted, where the science requirements were translated into a complete set of technical, operations, calibration, and data processing requirements. To deal with the complexity of budgeting and performance allocation, model-based system engineering was applied, see Lorenzo Alvarez et al. (2016) and Vavrek et al. (2016).

The WL probe is based on the determination of galaxy shear from the measurement of the shape of galaxies derived from their size and ellipticity. The shear determination de-

mands tight control of systematic effects down to the depth and resolution achievable; see Amara & Refregier (2008), Paulin-Henriksson et al. (2008), Paulin-Henriksson et al. (2009), and Cropper et al. (2013) for a detailed assessment of the *Euclid* case. The required properties of the PSF detectable with the visible imager (VIS) resulted in a set of image quality (IQ) requirements that have a major impact on the design of *Euclid*, imposing high stability of mechanical structures, temperatures, and telescope pointing (Racca et al. 2016). The design aims to provide the maximum knowledge of the PSF needed for the correction of the sources of systematic effects, whose aggregated error contribution should be significantly less than the statistical errors in the total error budget (Cropper et al. 2013).

Despite the very stable conditions that can be achieved in space, there can still be transient IQ variations induced by short- and medium-term thermal variations. These are caused by changes in Solar attitude during the execution of the survey plan and variations in heat dissipation by electronic units. At the design level, countermeasures taken include high-efficiency thermal insulation, the use of materials with a low coefficient of thermal expansion, and carefully managed spacecraft oper-

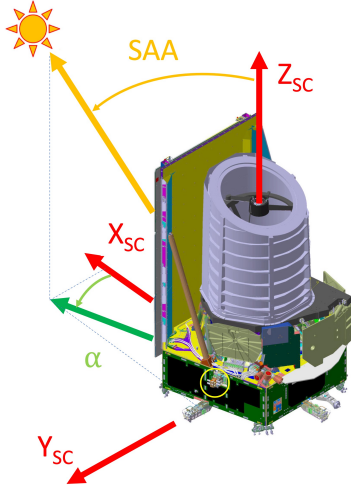


Fig. 1. Definition of Solar aspect angle (SAA) and α angle (AA) with respect to the spacecraft axes. SAA is the angle between the Sun direction vector and the Z axis of the spacecraft, while AA is the angle between the X axis and the projection of the Sun direction on the X-Y plane. The spacecraft is a schematic presentation to illustrate the configuration without the protective multi-layer insulation. Thermal disturbances occur whenever $\text{SAA} > 90^\circ$ where the Sun illuminates the bottom of the spacecraft, or $\text{AA} > -2.5^\circ$ where the Sun illuminates the micro-propulsion system (MPS) boom, encircled in the figure.

ational modes to minimise variations in power dissipation. The designed Solar attitude range of the telescope is limited to angles $-8^\circ < \text{AA} < 8^\circ$ and $87^\circ < \text{SAA} < 121^\circ$ (see Fig. 1).

As part of the design and development process, the adequacy of the *Euclid* design was verified by dedicated component and instrument model analyses complemented by test campaigns after assembly and integration of subsystems, and eventually the entire system; see Gaspar Venancio et al. (2022). To ensure the compliance of the *Euclid* design with the IQ requirements, we have developed the structural-thermal-optical performance (STOP) analysis, predicting the effects of thermal loads on the IQ stability during *Euclid*'s operation in-orbit. An overview of the thermal and mechanical design and the development steps relevant for the analysis described in this document is provided in Appendix A.

In this paper, we describe the image quality performance of *Euclid* for the WL probe as expected from the STOP analysis, and we compare the STOP predictions with the IQ metrics measured during early *Euclid* operations. Section 2 discusses the requirements and the accompanying image quality metrics. Section 3 describes the approach to the analysis, the mathematical models, the model integration into the STOP model suite and the model validation. Section 4 gives details of the thermal, thermoelastic, and optical analyses, and predictions about the satellite-level performance in terms of IQ metrics. Section 5 provides details of the early in-orbit performance and describes some ancillary analyses performed to support satellite commissioning. Finally, in Sect. 6 we compare the performance with the predictions and in Sect. 7 we draw the conclusions and give suggestions for the application of STOP analysis during in-orbit operations.

2. IQ performance requirements

The spacecraft requirements applicable to the STOP analysis were derived from the science requirements for the WL image

Table 1. *Euclid* image quality requirements for WL measurements, the values give the maximum limits. The science requirements refer to the knowledge of the PSF at a fiducial wavelength of 800 nm and a maximum integration time of 700 s. The spacecraft requirements (column 3) give the derived allocations for the spacecraft. The science and spacecraft requirements for the bottom two lines refer to a period 700 s and 11 000 s, respectively.

Variable	Science req.	Spacecraft req.
FWHM ["]	0.18	0.155
Ellipticity	0.15	0.14
R^2 [arcsec ²]	$4R_{\text{ref}}^2$ †	0.055
δe_i	2×10^{-4}	2×10^{-3}
$\delta R^2/R^2$	10^{-3}	2×10^{-3}

Notes.

† R_{ref}^2 is based on a Gaussian profile with a FWHM of $0''.2$.

quality, taking into account the apportionment to spacecraft, calibration, data processing, and survey strategy. The allocation to the spacecraft design was driven by the practical realisation of the design, by the expected performance of the intended in-flight calibrations, and the capability of the PSF model. The spacecraft requirements include contributions from the design residuals, integration alignment error, focus error, thermal stability, and spacecraft-induced image motion.

The size of the PSF is defined as the full width at half maximum (FWHM) averaged azimuthally about the PSF centroid.

$$\text{FWHM} = \frac{1}{\pi} \int_0^\pi \text{FWHM}(\alpha) d\alpha. \quad (1)$$

The ellipticity e is defined in terms of the Gaussian-weighted quadrupole moments, where $e = \sqrt{e_1^2 + e_2^2}$, and e_1 and e_2 are the two projected ellipticity components obtained from

$$e_1 = \frac{Q_{xx} - Q_{yy}}{Q_{xx} + Q_{yy}}, \quad e_2 = \frac{2Q_{xy}}{Q_{xx} + Q_{yy}}, \quad (2)$$

where Q_{ij} are the weighted quadrupole moments, as defined in Schneider et al. (2006) that involve a Gaussian weighting function with a standard deviation $\sigma = 0''.75$, based on the typical size of the smallest galaxies useful for weak lensing (Meneghetti et al. 2008). The IQ metric $R^2 = Q_{xx} + Q_{yy}$ bounds the PSF wings to a reference value R_{ref}^2 obtained from a Gaussian profile with FWHM of $0''.2$. The spatio-temporal variation of the PSF shall allow a model of the PSF as a function of time and position in the image plane with a residual model uncertainty smaller than δe_i and $\delta R^2/R^2$.

The science requirements and the apportioned spacecraft requirements are listed in Table 1. Besides the static requirements for a fiducial exposure time of 700 s, it also provides the transient requirements for the spacecraft imposing a temporal stability for $\delta R^2/R^2$ and the ellipticity components. The temporal stability of the spacecraft is required to improve the model of the PSF by aggregating stars from more than one field (Cropper et al. 2013). The spacecraft requirements are defined between any two of eight consecutive dither exposures of two adjacent survey fields (Euclid Collaboration: Scaramella et al. 2022), with a maximum duration of 11 000 s for 8 consecutive dither exposures. The corresponding science requirements refer to the knowledge over an

exposure time of maximum 700 s and therefore have smaller values. In the remainder of this paper, we will refer only to the spacecraft requirements.

3. STOP analysis approach

The STOP analysis considers all thermal loads that affect the stability of the telescope in-orbit. The thermal loads acting on the telescope come from its environment, which include the Sun, payload module (PLM) containing the telescope and scientific instruments, service module (SVM) including electronic units, and the sunshield. The STOP analysis shall predict the impact on the IQ parameters, and the resulting distortion-driven error terms are considered in the system performance budget. The results of the analysis during the design phase were used to improve the system design as part of an iterative process toward the final design.

The design of the spacecraft must meet its requirements in all possible conditions during the lifetime of the mission. To define design cases, extreme conditions are considered for the permissible attitudes with respect to the Sun, for changes of thermo-optical properties of the exposed surfaces because the Solar absorptivity of the paint is expected to increase with age, and for electronic unit power dissipation taking ageing into account. The situations analysed are therefore conservative and relative to compounded worst cases.

Each iteration of the analysis proceeds through three coordinated sets of mathematical models according to the following sequence: temperature sets at appropriate “thermal” locations → interpolation to corresponding “structural” locations → displacement and distortion of the optical path and elements → application to the optical model → resulting optical aberrations (IQ metrics).

3.1. Mathematical models

The system thermal mathematical model (TMM) was assembled from models representing each major component: SVM, PLM, and sunshield. The component models were prepared and tested separately. For the PLM, two configurations were provided, operational worst-hot and operational worst-cold. In the system analysis, the worst-cold configuration was associated with beginning-of-life conditions and the worst-hot configuration with end-of-life conditions, due to the degradation of the thermo-optical properties of the surfaces.

The STOP structural analysis finite element model (FEM) considers the model assembly of the major components using MSC NASTRAN software (2008 and 2016 versions). The PLM and sunshield are attached to the SVM structure by means of rigid elements and bars. There are no structural links between PLM and sunshield. The total mass of the model is 1974 kg. Customization of the FEM for thermo-elastic analysis included the application of thermal expansion coefficients. The coefficient of thermal expansion (CTE) can vary with temperature and for some materials the difference between room temperature and the operating temperature of *Euclid* can be large, e.g. the CTE of Silicon Carbide (SiC) used in the PLM structures changes by a factor of about 4. For these sensitive materials, CTE values measured at a nominal operating temperature of 135 K were applied.

The optical model consists of the telescope aperture and the VIS image plane, via the optical train composed of the primary mirror (M1), secondary mirror (M2), folding mirror 1 (FOM1), folding mirror 2 (FOM2), tertiary mirror (M3), dichroic filter, and folding mirror 3 (FOM3), see

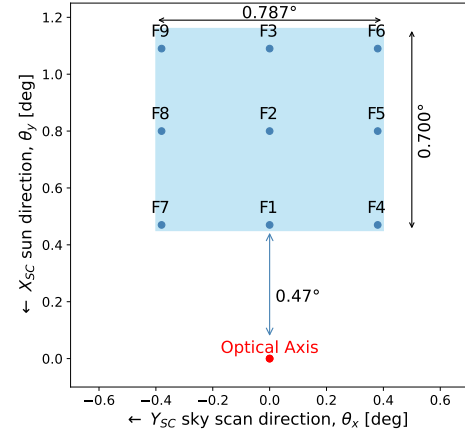


Fig. 2. VIS image plane field points. The Korsch telescope design provides an off-axis exit pupil. See Fig. 1 for the definition of the spacecraft X, Y, and Z coordinates.

Euclid Collaboration: Mellier et al. (2025) for a schematic. The PSF was computed for 9 sample points in the VIS focal plane by ray tracing, see Fig. 2. In the ray-tracing, the mirrors are represented by the location of their vertices, translated and rotated according to the thermo-elastic results. The effect of temperature on the curvature of the mirrors, whose main effect is defocus, was added using the analytical procedure outlined in Appendix B.2. Ray tracing optical design software Code V¹ was customized to process the displacements and rotations provided by the structural analysis. All coordinates are referred to a global reference frame centred on the M1 vertex. Numerical noise in the computation of the ellipticity, the most sensitive figure, due to small differences in pupil sampling, is on the order of a few parts in 10^4 . Other contributors, such as small differences between the actual coded algorithms, are one order of magnitude below.

The STOP analysis used the Multidisciplinary REentry Analysis (MaREA) software package, developed by Thales Alenia Space, to efficiently transfer data between mathematical models of different types. Using MaREA functions, the three mathematical models (thermal TMM, structural FEM and optical CODE V) were integrated into one workflow. The process was then implemented in the SIMULIA Isight® execution engine² for sequential runs of multiple analysis cases and statistical post-processing of the results. See Appendix B.1 for a detailed description.

3.2. Definition of analysis cases

IQ parameters were determined for a number of possible in-orbit situations in which the telescope is in a thermally stable state. The variables and their values for the steady-state cases are provided in Table 2. We obtained the IQ parameters for each thermal combination involving four Solar aspect angles (SAAs) and three α angles, for two operational modes: nominal (NOM) and higher dissipation communications (COM) mode, two values of the conductive interface, and two Solar irradiances, resulting in 96 cases.

¹ see <https://www.synopsys.com/optical-solutions/codev.html>

² SIMULIA Isight® is a commercial process automation and simulation tool developed by Dassault Systèmes.

Table 2. Definition of variables for steady-state analysis cases.

Variable	Value(s)
Solar irradiance:	
BOL, EOL*	1293 Wm ⁻² , 1388 Wm ⁻²
SAA	87°, 90°, 105°, 121°
AA	8°, 0°, -8°
Operational Mode	NOM, COM**
SVM top floor interfaces:	
conductive	20 °C, 24 °C
radiative	worst-case MLI***
Sunshield to PLM interface:	
radiative	worst-case MLI

Notes.

- * begin of life (BOL) and end of life (EOL) conditions
 ** Nominal (NOM) and Communications (COM) modes
 *** Multi Layer Insulation

Table 3. STOP Transient sizing case (hottest to coldest).

Variable	Value
Solar irradiance	1388 Wm ⁻² (EOL)
PLM Operational Mode	OpHotWorst
Start SAA	121°
Start AA	8°
Start SVM Operational Mode	COM
End SAA	87°
End AA	-8°
End SVM Operational Mode	NOM

The transient case is intended for the verification of the image quality stability requirements (ellipticity and R^2). Table 3 provides the definition of the transient case, covering the transition from the hottest to the coldest condition (EOL, at constant Solar irradiance), as allowed by the Solar illumination requirements. A mode transition of the SVM is included to envelope the extreme worst cases of environment variation, but is not representative of any real-life case, where the changes in attitudes and dissipation are always minimised.

4. STOP analysis in the satellite design phase

4.1. Thermal analysis

The thermal response of the telescope structure to changes in the environment provides already a strong indication of the possible IQ changes which are quantified in the subsequent STOP analysis steps. A general property consistently observed is the linear correlation between the temperatures of the main optical elements and the average temperature of the telescope baseplate. This property of the baseplate is due to its large mass and surface area. The baseplate accounts for more than half of the mass of the SiC parts in the telescope dominating the telescope thermal capacity. In addition, the baseplate has the largest surface area facing the SVM, making it the main component for the radiative exchange between the PLM and the SVM.

The PLM settings (dissipation, heater power) are important in determining the telescope temperatures and, to a lesser extent, their variations. All thermal states were calculated from an initial state with a uniform 20 °C temperature in all thermal

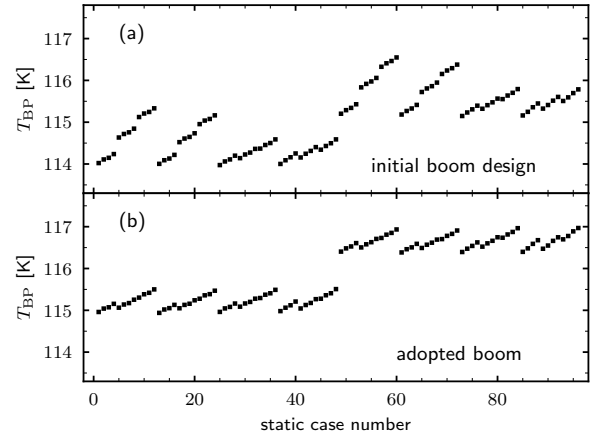


Fig. 3. Baseplate average temperature versus STOP case number, see Sect. 4.1.1 for the sequencing description. **Upper panel (a):** full-length micro-propulsion system (MPS) boom, the topside can be illuminated by the Sun. **Lower panel (b):** half-length MPS boom, the topside cannot be illuminated by the Sun.

nodes. Consequently, the calculated deformations included the large cool-down component caused by the temperature transition from 20 °C to the operating temperature of each thermal node. The common cool-down component was removed in the structural analysis by taking the difference of each state vector (translation and rotation) with a reference state vector associated with STOP case #19 (SAA = 90°, AA = 0°, interface temperature = 20 °C, nominal mode, see also Sect. 4.1.1).

4.1.1. Steady-state thermal analysis

The steady-state analysis addresses the 96 STOP cases defined in Table 2. The results are presented in plots of temperature versus STOP case number using the following convention.

- The first set of 48 points refers to beginning of life (BOL) conditions and the second set of 48 points (#49 to #96) refers to end of life (EOL) conditions. BOL is obtained from the minimum yearly irradiance at summer solstice at L2 distance from the Sun, and EOL from the irradiance at winter solstice at L2.
- In both the BOL set and the EOL set, the first set of 12 points is calculated at SAA = 87°, the second set at SAA = 90°, the third set at SAA = 105° and the fourth set at SAA = 121°.
- Within each set of 12 points defined above, the first four are characterised by AA = -8° (Sun on the spacecraft -Y side), the second set of four by AA = 0°, and the third set by AA = +8° (Sun on +Y side).
- Within each subset of four consecutive points, the first two share the NOM operational mode and differ by the conductive interface temperature, 20 °C or 24 °C, and the second two share the COM operational mode and again differ by the conductive interface temperature, 20 °C or 24 °C.

From the resulting Fig. 3 one can trace the dependence on SAA (increasing left to right) and on azimuth (increasing bottom to top).

Figure 3a shows the average temperature of the baseplate T_{BP} for the 96 steady-state thermal cases using the initial design configuration of *Euclid*. For SAA = 87°, an increase of T_{BP} as large as 1.4 K was found when rotating the spacecraft from AA = -8° to AA = +8°. Such a large temperature change was not expected. Initial estimates indicated that the ellipticity would

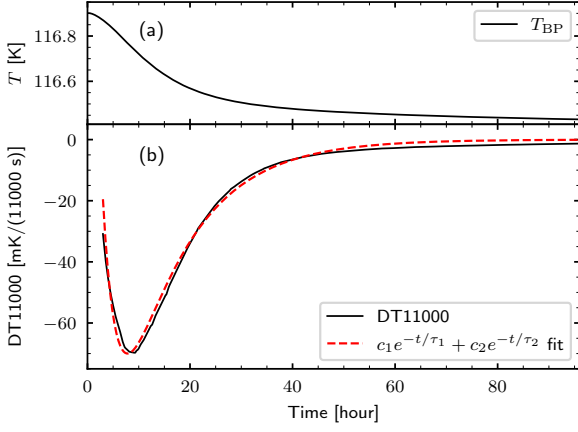


Fig. 4. Temperature transient analysis. **Upper panel (a):** baseplate temperature evolution over 96 hours. **Lower panel (b):** the temperature difference DT11000 derived from the temperature evolution in (a); the dashed curve shows a match of DT11000 with two exponential terms with $c_1 = 168$ mK and $c_2 = 315$ mK.

change by 5% per K temperature variation in T_{BP} , the derived variations would violate the IQ requirement for ellipticity.

The dependence of T_{BP} on Sun azimuth was attributed to the presence of the MPS boom on the spacecraft +Y_{SC} side, in front of and below the VIS radiator. Radiation from Sun-facing side of the boom reflected off the back of the sunshield and reached the telescope baffle, which is mounted on the baseplate, see Fig. C.1 in Appendix C. This finding led to a redesign of the MPS boom. In a first approach, the multi-layer insulation (MLI) on the boom topside was reconfigured such that the topside remained in the shadow. For the final configuration, the boom length was reduced to half its initial length (Fig. 3b). This design solution would still give sufficient torque authority without causing a significant increase of T_{BP} compared to a solution without boom. The boom redesign is described in Appendix C.

4.1.2. Transient thermal analysis

We simulated a transient case that should provide maximum stress to the PLM by imposing extreme changes; see Table 3. Assuming EOL, we define an attitude transition simultaneously comprising the maximum allowed ranges of SAA and AA. The transition brings the Sun from the upper limit of the allowed SAA (121°), a warm condition for the SVM, to the lowest SAA (87°), a colder condition. The SVM warming is due to the illumination of the bottom platform by the Sun whenever SAA > 90° (cf. Fig. 1). A mode change from COM to NOM is also included, causing a high to low variation in power dissipation in the SVM. The transition causes a significant cooling of the SVM and a slight warming of the PLM.

In the calculation, the attitude transition is instantaneous, applied at $t = 0$, and the subsequent evolution is tracked for four days. It should be noted that four days are not enough to reach a steady condition, we estimate a drop below 0.1% after seven days assuming exponential decay and a time constant of one day; therefore, the end state differs somewhat from that found in the corresponding case of the steady-state analysis. However, four days are long enough to capture all the essential features of the transient.

The temperature inputs in the IQ analysis are the temperature differences over a 11 000 s sliding time window: $DT11000 = T(t) - T(t - 11\,000\text{ s})$. Figure 4a shows the telescope baseplate

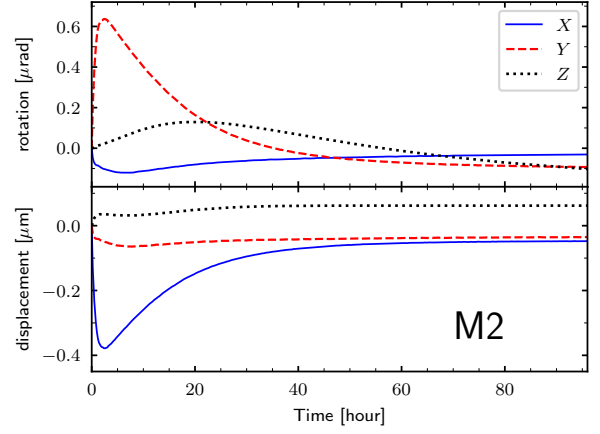


Fig. 5. M2 rotations (upper panel) and displacements (lower panel) around the spacecraft axes after a transition according to the parameters in Table 3.

temperature evolution over 96 h after the attitude transition and Fig. 4b shows the corresponding DT11000 for the case stipulated in Table 3. The dip in the DT11000 curve, lowest after about 8 h, is explained as due to the interaction of two temperature components that act in a SAA slew. The displacement of the Sun from SAA = 121° to 87° cools the SVM and heats the sunshield. The two effects propagate to the PLM baseplate with different time constants. The superposition of the two effects produces the dip. The DT11000 curve is well matched by the combination of two exponential functions with time constants of $\tau_1 = 12.3$ h, associated with the PLM, and $\tau_2 = 2.8$ h, associated with the SVM, see Fig. 4b.

A sensitivity analysis was performed to study the dependence of the temperature perturbation on the environmental variables and the spacecraft properties. The results are summarised as follows. The smaller the attitude variation, the lower the temperature fluctuation and the minimum fluctuation is achieved when limiting SAA to less than 105°, and the azimuth range to $-5^\circ < AA < 5^\circ$. Minimising the change in power dissipation in the SVM, too, reduces the magnitude of the temperature fluctuation. Reversing the start and end points of the transition described in Table 3 does not produce a symmetric temperature change, and the magnitude of DT11000 is lower in the “warming” transition. The properties of the surface finishes have a small but noticeable effect on the position of the peak and its duration, which may be relevant in the case of ageing.

4.2. Structural analysis

The temperature variations from the thermal analysis are applied to the corresponding elements of the structure to obtain the displacements and rotations induced in the telescope optical train. The structural analysis addresses both the 96 steady-state temperature perturbations defined in Sect. 4.1.1 and the transient perturbations in Sect. 4.1.2.

4.2.1. Steady-state structural analysis

The elements of the optical train comprise 8 points representing the centre positions of the main mirrors M1, M2, and M3, the folding mirrors FOM1, FOM2, and FOM3, the dichroic, and the image plane.

The solid-body motion is removed by associating each element to the reference frame defined by M1. Since the temper-

Table 4. Metrics of the “as built” telescope optical model.

Field Point	Ellipticity	R^2 (arcsec ²)	FWHM (")
F1	0.005	0.046	0.1344
F2	0.007	0.045	0.1338
F3	0.014	0.047	0.1340
F4	0.027	0.051	0.1327
F5	0.033	0.046	0.1342
F6	0.020	0.051	0.1362
F7	0.044	0.049	0.1355
F8	0.050	0.047	0.1347
F9	0.025	0.049	0.1353

ature perturbations are computed for each case from a uniform temperature of 20 °C, they contain the ground-to-orbit temperature transition, which leads to a large deformation. This is removed in-orbit by focussing the telescope using the M2 mechanism. This operation is simulated by taking the difference of each state with that associated to a reference case, as described in Sect. 4.1. Taking the differences is justified in a linear approximation, given the smallness of the effects once the common-mode motions have been removed.

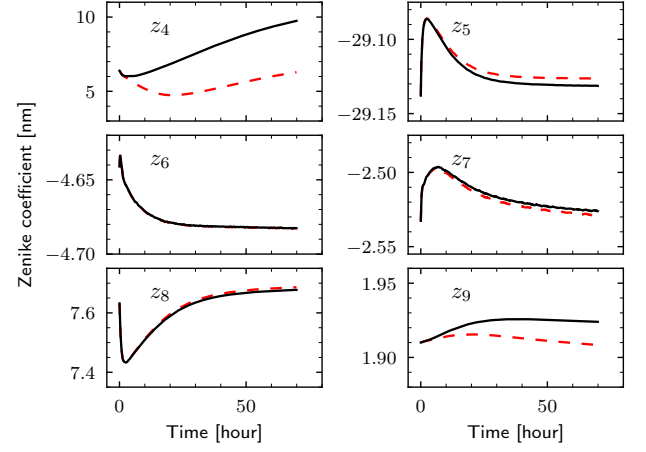
In all nodes, the maximum magnitude of each component of the displacement vector is found to be less than 1 μm , and the maximum rotation of the reference vector is less than 1 μrad . These displacements and rotations become significant in the optical analysis in Sect. 4.3.

4.2.2. Structural stability analysis

We have computed the displacements and rotations of the optical elements with respect to their reference positions as a function of time for 4 days following the attitude transition defined in Table 3. All optical elements exhibit a rapid change in the first few hours after the transition, concurrent with the temperature perturbation of the baseplate described in Sect. 4.1.2. The change appears as a damped oscillation, with a quick rise and then reversing direction with a slow exponential return. See, for example, the y-component of M2, displayed in Fig. 5, with peak magnitude 0.65 μrad .

4.3. Optical analysis

For optical performance calculations of the *Euclid* telescope, the PSF at the desired points were generated by the CODE V ‘as built’ optical model based on the nominal optical layout constructed after correlation with the results of the PLM thermal vacuum test; see Appendix A. The PSF is then sampled with a pitch of 4 μm and processed using a 2048 $\times$ 2048 fast Fourier transform (FFT). The map is then truncated to 256 $\times$ 256 points centred on the maximum value. The truncation is justified by the fact that none of the weighted IQ metrics (ellipticity, R^2 , and FWHM) is significantly dependent on the part of the PSF energy that is beyond $128 \times 4 \mu\text{m} = 512 \mu\text{m}$ or 4'':26 from its centre. The image quality metrics are calculated using a discrete version of the equations given in Sect. 2. Table 4 shows the metrics of the nominal telescope realization. This process is time-consuming, and not suited for the STOP analysis where many different static cases, as well as temporal transient cases, have to be analysed. To ease and quicken the process, simplified models have been used, as described in Appendix B.2.

**Fig. 6.** 11 000 s variation of Zernike coefficients z_4 to z_9 . Dashed lines: telescope only; solid lines: telescope and mirrors.

4.3.1. System steady-state IQ performance

The IQ metrics were calculated for the 96 cases defined in Table 2, both without and with the inclusion of the predicted deformation of the mirrors. We find that the influence of the telescope optical train deformation is very small and that the mirrors contribute to almost all the variance among the cases.

The different cases indicate notable variations in ellipticity due to changes in azimuth and operating mode. Little variation is caused by the different SAA. All variations are well within 0.5% for the contribution by the optical train deformation alone. The deformation by the mirrors amplifies the ellipticity variations to about 1.0% for the most perturbed field points F2, F6, F7, and F9 (cf. Fig. 2). The most perturbed field points are well within the maximum allowed ellipticity (0.14). A similar behaviour is found for R^2 . The FWHM is hardly affected by thermo-elastic perturbations. We point out that R^2 is more sensitive to the PSF wings than the FWHM.

4.3.2. System in-orbit IQ stability performance

The stability of the telescope metrics is evaluated starting from the time series of displacements/rotations of the optical train described in Sect. 4.2.2, originating in the extreme thermal transient described in Sect. 4.1.2. The transients were interrupted after 4 days, in a condition close to the final steady state (ΔT with respect to the corresponding static STOP case < 0.1 °C in all relevant nodes). The optical metrics were calculated up to 70 h, for reasons of containing the time-consuming computations, but the essentials of the transients are clear anyway. See Appendix D for the computed timelines.

To understand the underlying variations in the shape of the PSF causing the IQ metrics, we analyse Zernike components z_4 to z_9 (adopting Fringe’s ordering and notation) from which the PSF can be constructed. The variations of the components are presented in Fig. 6. It shows that the largest variations are found in the defocus term (z_4) followed by astigmatism (z_5) and coma (z_8), both of which are already an order of magnitude smaller than z_4 . The additional defocus contribution due to mirror deformation (cf. Appendix B.2) increases monotonically towards an asymptotic value and dominates the value of z_4 near the end of the transient. The z_4 telescope term has a minimum after about 20 hours and then goes back to its original value. The mirror term correlates with the mean M1 temperature, whereas the telescope term correlates with the vertical displacement of M2 relative to

Table 5. *Euclid* Image Quality performance (VIS channel). The STOP performance values were calculated before launch; they are worst cases and do not belong to the same realisation. The in-orbit ellipticity and FWHM values were derived from a VIS frame taken from the Early Release Observations programme; the in-orbit R^2 were measured from VIS frames during telescope alignment.

Variable	STOP	In-orbit	Required
FWHM ["]	0.146	$0.141 \pm 0.006^*$	0.155
Ellipticity $ e $	0.074	$0.024 (0.05)^{**}$	0.14
R^2 [arcsec ²]	0.053	$0.048 (0.050)^{**}$	0.055
$\delta\epsilon$	1.5×10^{-3}	–	2×10^{-3}
$\delta R^2/R^2$	1.5×10^{-3}	–	2×10^{-3}

Notes.

* uncertainty range is based on the minimum and maximum measured values in the field.

** average value, the maximum value in the field is given between parentheses.

M1. Changes in the curvature of the mirrors, in particular M1, affect the focus, which is the dominant term. This justifies the telescope re-alignment using only M2 adjustments. Sensitivity analysis shows that virtually all ellipticity variation is cancelled when the M2 ΔZ variable is constrained to zero.

The higher order terms z_5 and z_8 both exhibit a fast transient with almost the same magnitude but opposite signs. The transient peaks after about 2 h and then returns to its original value, with a small offset, after 24 h. This transient can be modelled as the sum of two opposing terms with time constants of about 10 h and about 1 h (cf. Sect. 4.1.2). The 10 h constant is associated with the baseplate. The fast component can be attributed to the telescope truss, as suggested by the strong correlation with the lateral M2 displacement ΔY . The short time constant can be explained by a strip of thermal nodes associated with the external baffle and VIS radiator with a view factor to the back of the sunshield (see Appendix C) and changing its temperature on a timescale of less than one hour. This heat is dissipated via the baseplate.

4.4. Performance synthesis before the launch

Table 5 shows the performance compared to the requirements. Each performance value is the worst case found among all cases studied, among all field points and all time instants.

The results of the prelaunch STOP analysis showed that the design is compliant with the applicable requirements under the most extreme conditions, both steady-state and transient. The margins are generally large, even more so taking into account that we have exercised instances where the performance comes close to the specified upper boundary, and the transients are the largest allowed by the design limits.

The goal of the STOP analysis was to study in the design stage the impact of the environment on the PLM to identify and eliminate any significant harmful effects. The analysis was not intended to predict the detailed behaviour of the satellite and payload in-orbit, a purpose for which other tools are available; see Appendix A for further details on the thermal design. Nevertheless, the findings of the STOP analysis are useful for the interpretation of the telemetry and the assessment of the in-orbit performance, the subject of the following sections.

5. STOP in-orbit

5.1. In-orbit situation

After launch on 1 July 2023, the STOP model was used in support of the commissioning and performance verification phases, to help the interpretation of the in-orbit performance and to verify the STOP analysis predictions.

We examined the spacecraft attitude relative to the Sun, AA and SAA, and data from VIS and Near-Infrared Spectrometer and Photometer (NISP) units whose thermal dissipation affects the telescope. These data were extracted from satellite housekeeping telemetry. Temperature sensors are mounted at thermally important positions in the PLM providing means to monitor thermal responses to environmental changes. In particular, 20 high-resolution (< 10 mK) thermistors, complemented by a number of lower-resolution (< 60 mK) sensors, monitor the thermal state of the baseplate. The record shows that, while the mean temperature is different at each position, the temperature variations are nearly the same, so that any one of the sensors may be used as an index of the thermal state of the baseplate, which correlates with the IQ. Variations in telescope structure and image quality cannot be monitored from housekeeping telemetry but must be inferred from stellar images obtained with VIS.

Instrument activities generating heat dissipation were deduced from the readout shutter unit (RSU) movements for VIS, and the filter and grism wheel movements of NISP. These movements serve as a close proxy for the dissipation cycle of the instruments. The effect on temperature can be detected only when there is a long sequence (> 12 h) of identical cycles that add up. VIS bias or dark frames, which are read out without moving the RSU, cannot be detected in the temperature curve unless they occur in a long sequence during dedicated calibration operations. For the VIS effect on the baseplate, we cannot distinguish between heating by the shutter motor or heating by the readout electronics (ROE) via the flex wires connecting the ROE to the FPA.

The spacecraft operating attitude domain was redefined after a parasitic straylight component was found in the VIS images at positive AA, see Euclid Collaboration: Mellier et al. (2025). Subsequent analysis showed that the MPS boom is illuminated at $AA > -2.5^\circ$. The limiting angles used for the construction of the survey had to be set to $87.1^\circ < SAA < 120.9^\circ$ and $-8.4^\circ < AA < -3.0^\circ$ to meet the survey requirements while keeping the straylight at an acceptable low level. The negative azimuth limit is 0.4° beyond the limit that was considered for the STOP analysis.

During spacecraft commissioning it was found that switching on/off by the telemetry radio frequency transmitter in the SVM induced an unwanted rotation around the spacecraft Z axis, causing a disturbance in the pointing. It was decided to permanently leave the transmitter on to minimise disturbance (Gottero et al. 2024). This was done on 2023-10-06, raising the average T_{BP} with a few tens of mK, but below 50 mK.

5.2. In-orbit thermal performance

During the *Euclid* performance verification phase from October to early December 2023, numerous calibration activities were carried out, requiring different operating modes from those used during the survey. These activities caused exceptional environmental stresses on the telescope and provided extreme situations suitable for studying thermo-optical effects.

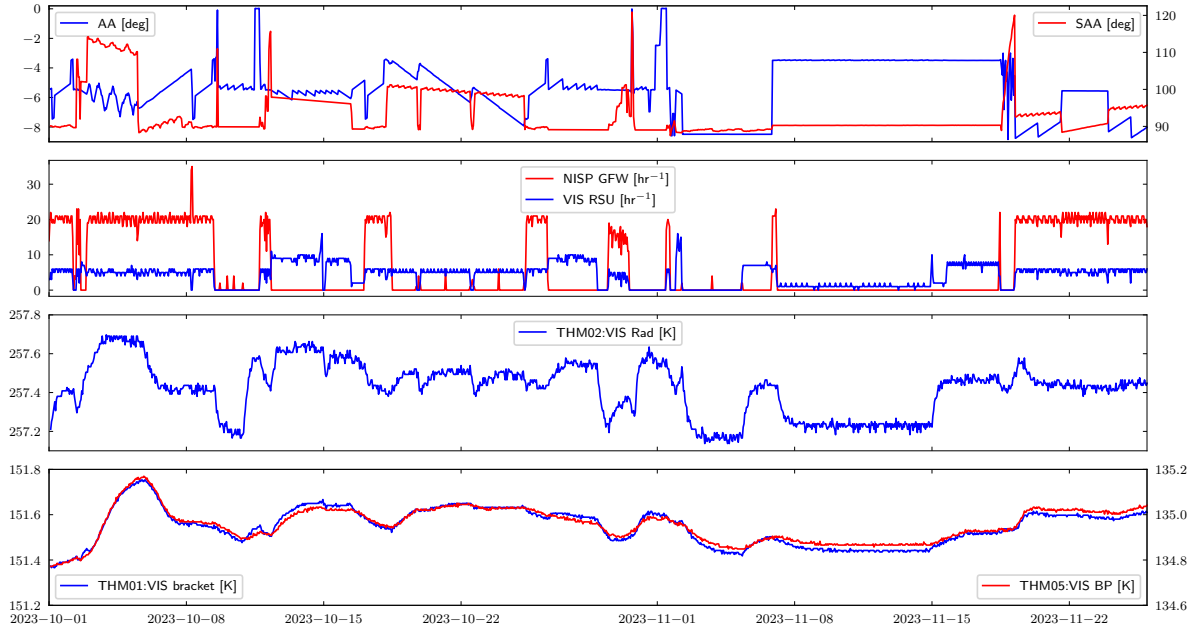


Fig. 7. Time histories of spacecraft attitude and key thermal performance parameters obtained from spacecraft telemetry during the *Euclid* performance verification phase in the period from 2023-10-01 00:00:00 UTC until 2023-11-26 00:00:00 UTC. **Upper panel:** the variations in SAA and AA. Dedicated rotations of the SAA (red line) and AA (blue line) were performed to verify the thermal response of the system. **Upper middle panel:** the history of the cadence (in cycles per hour) of the NISP grism and filter wheel (GFW, red line), and the VIS shutter (RSU, blue line). **Lower middle panel:** the history of the VIS radiator temperature registered by sensor THM02 (blue line), which is an indicator of the VIS instrument activity by means of its power dissipation. **Lower panel:** the temperature responses of a baseplate sensor placed near the VIS bracket (THM05, red line) and a sensor mounted directly on the VIS focal-plane array (FPA) bracket (THM01, blue line).

Before launch, only the Sun incidence angles SAA and AA were considered to be the dominant sources of temperature variation. The in-orbit data indicated that non-standard (i.e. non-survey type) instrument operations unexpectedly dominated the baseplate temperature history. Taking the time histories of various observables, we shall first assess the impact of attitude changes (Sect. 5.2.1) and subsequently investigate the additional impact of other causes of temperature changes (Sect. 5.2.2).

The time histories of AA and SAA, and corresponding temperatures of the baseplate, VIS FPA, and the VIS radiator are shown in Fig. 7, covering the period 2023-10-01 until 2023-11-26, and Fig. 8, covering the period 2024-09-01 until 2024-12-31. The baseplate temperature T_{BP} was obtained from sensor THM05 located close to the VIS instrument. The VIS FPA temperature was obtained from sensor THM01 located on the VIS FPA bracket, while sensor THM02 provided the VIS radiator temperature.

5.2.1. Thermal impact of SAA and AA variations

Fig. 7 includes four episodes in which the Solar aspect angle was moved high above 90° and remained there for several days. In the first such episode, the “hot case”, a step up of SAA from 90° to 113° was executed on 2023-10-02, followed by a 3 day stay on an average of 110° , and then a step down back to 90° on 2023-10-05. The step up was accompanied by an exponential increase of T_{BP} (THM05), headed for a new equilibrium about 400 mK above the initial state, to be reached after about 4 days, but interrupted after 3 days by an SAA step down back to 90° , initiating an exponential drop levelling ≈ 200 mK below the peak.

Three more SAA steps, of smaller magnitude, were then executed: from 90° to 97° starting on 13 October and lasting about 4 days; from 90° to 100° average starting on 2023-10-18 and

lasting about 7 days; and from 90° to 99° average around 2023-10-30 lasting less than 1 day. The longer steps with around 10° SAA variation cause the baseplate temperature to rise by about 100 mK, always followed by an exponential drop when the SAA is brought back to 90° .

All episodes of SAA rise exhibited an exponential time constant of approximately 1 day and a temperature rise rate, from start to plateau, of 5 to 15 mK per degree of SAA. These numbers are rough estimates deduced after accounting for other phenomena acting on the temperature, described below.

From 2023-10-30 on, the SAA remained around 90° , while the baseplate temperature underwent a slow descent, the tail of the exponential drop after the last step down, modulated by other effects that cannot be explained by the SAA.

A thermal response to changes in AA would be expected in Fig. 7, from 2023-10-18 to 2023-10-25 (sloping in steps from -3.8° to -8° over 7 days) and from 2023-11-07 (large step from -8.5° to -3.5°). However, the effect of AA on T_{BP} is hardly discernible. This is most clearly seen at the time of the large AA step on 2023-11-07. The temperature began to rise some time earlier, therefore the AA step cannot be the cause of the rise. The temperature started to drop soon after the step, and the positive AA step could be the cause of the temperature drop. However, the earlier AA-sloping segment contradicts this hypothesis: a negative AA trend should have caused a temperature rise, which did not appear.

Neither SAA nor AA changes can account for numerous small features of the T_{BP} curve of Fig. 7, addressed in the next section.

5.2.2. Dependence on instrument activities

The temperature effects of the VIS operation are evident in the timeline of the VIS radiator (THM02) shown in Fig. 7. The

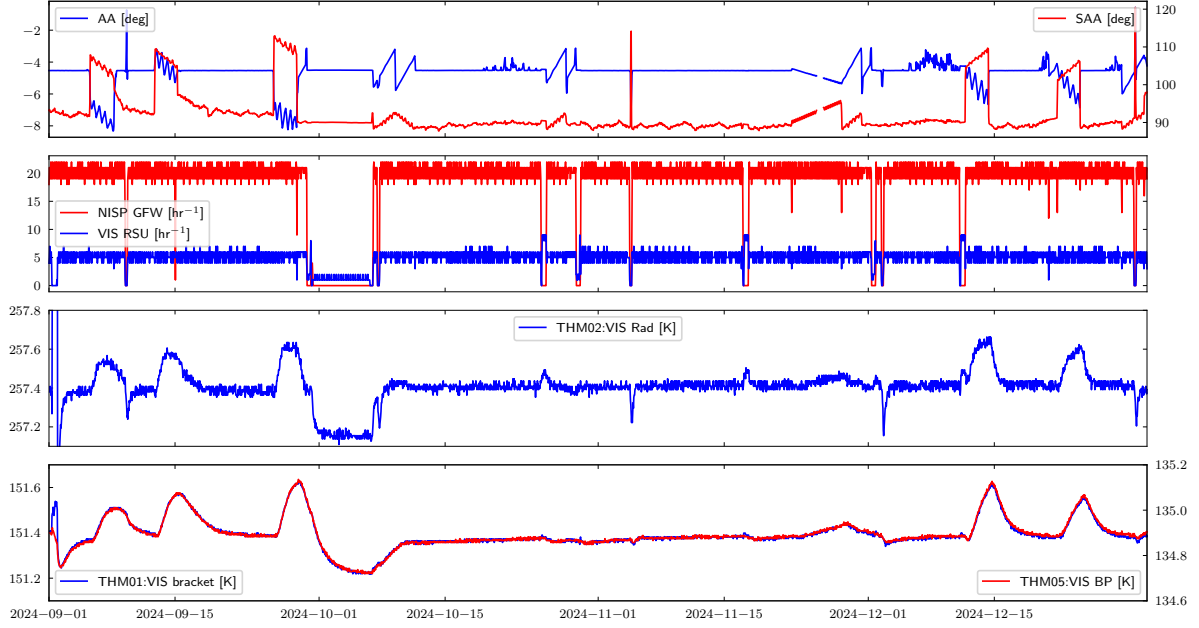


Fig. 8. Time histories of spacecraft attitude and thermal performance parameters obtained from spacecraft telemetry during the period from 2024-09-01 00:00:00 UTC until 2025-01-01 00:00:00 UTC. **Upper panel:** the variations in SAA and AA. **Upper middle panel:** the history of the cadence (in cycles per hour) of the NISP GFW (red line), and the VIS RSU (blue line). **Lower middle panel:** the history of the VIS radiator temperature from sensor THM02. **Lower panel:** temperature responses of baseplate sensor placed near the VIS bracket THM05 (red line) and a sensor THM01 directly on the VIS FPA bracket (blue line).

radiator enables the cooling of the VIS proximity electronics mounted behind the FPA (Euclid Collaboration: Cropper et al. 2025). The radiator temperature, which depends on SAA like all other elements, rises quickly to a new equilibrium when the VIS exposure cadence changes to a higher value and drops with changes to a lower cadence. The temperature of the VIS FPA bracket (THM01) follows the VIS radiator, with a damped magnitude and a longer time constant close to that of the baseplate temperature (THM05). At times when the radiator temperature changes abruptly, indicating a change in the cadence of the readouts, the temperature fluctuation of THM01 gets slightly larger or smaller than that of THM05, causing the two curves in Fig. 7 to separate.

Figure 7 covers parts of the performance-verification (PV) phase and, in particular, the phase-diversity calibration (PDC) demonstration between 2023-11-01 and 2023-11-27. During this period, numerous calibration activities were performed, subjecting the PLM to a rapid succession of enhanced thermal perturbations unlike anything experienced during the normal survey. It is exceptionally suitable for assessing the impact of payload operations on the thermal stability of the PLM. Here, the temperature curves show, in addition to the global impact of the SAA, a distinct signature of each calibration block.

The temperature effect of the VIS operation can be appreciated on 2023-11-05. At that epoch, the SAA had been stable around 90° for 5.5 days, and the VIS and NISP instruments had been inactive for 4 days; as a consequence, both the VIS radiator and the baseplate had dropped to their lowest temperature. On 2023-11-05 08:19:25 UTC, a series of 256 short (92 s) VIS exposures was initiated, lasting 1.5 days. The high cadence was associated with a temperature increase of the VIS radiator by 300 mK, and of the baseplate by 50 mK. The phenomenon can be explained by each readout event injecting heat pulses by the RSU motor and the readout electronics, via the dissipation directly to the baseplate and via the thermal straps connecting FPA bracket and baseplate, respectively.

Late on 2023-11-06, AA was moved from $-8^\circ 5'$ to $-3^\circ 5'$, where it stayed for about 6 days, during which a series of long exposures of 560 s was taken. In this period, we observed a drop of 200 mK in the VIS radiator temperature, and an exponential drop in T_{BP} of 35 mK, with a time constant of 29 hours (the time constant of the baseplate was measured during the telescope cool-down of 5 days from July 15 and was found to be 20 hours, while that of M1 in the same period was about 33 hours). As explained above, this temperature drop can be completely attributed to the lower dissipation due to the longer time intervals between successive readouts, and not to the change of AA.

A small impact of the NISP filter and grism wheel mechanism motors on T_{BP} can be inferred when comparing temperatures at epochs when the wheels are on and off. We deduce a contribution of the NISP wheels of less than 30 mK.

5.2.3. Thermal stability during nominal survey operations

The thermal history of the system for a four-month period during the nominal mission is presented in Fig. 8 displaying the same parameters as in Fig. 7. The baseplate temperature T_{BP} registered by THM05 appears to be dominated by five large excursions in SAA, which for each case was rotated from around 90° , to more than 105° , and back, causing T_{BP} to increase by more than 100 mK. The large SAA rotations were triggered by observations of several days in the *Euclid* Deep Field South for dedicated calibrations.

The period between 2024-09-30 to 2024-10-07 was allocated to a biannual calibration block which contains a NISP non-linearity calibration of 6.7 days, which did not require filter wheel movements and was accompanied by low VIS activity. These operations caused a prominent dip in T_{BP} of more than 100 mK below the baseline. Interruptions involving large SAA changes confined during the 12 h monthly satellite maintenance operations had a much smaller effect on T_{BP} . Other,

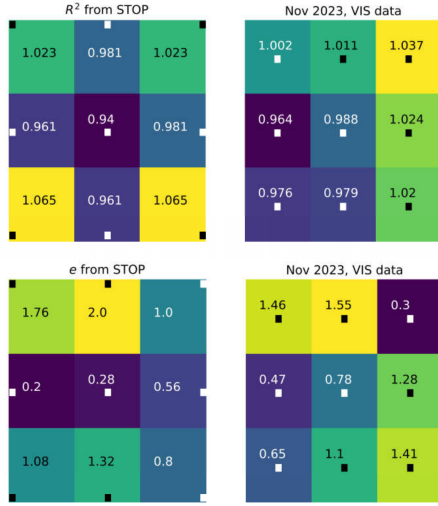


Fig. 9. Visualisation of the relative variation of R^2 and e over the VIS focal plane, +X to the right and +Y going up. The values are normalised by the average of the 9 values, with average R^2 of 0.0479 and 0.0199 arcsec² and average e of 0.025 and 0.021, for the STOP and in-orbit data, respectively. The difference between the values are due to the usage of a different Gaussian weighting in the calculation of the moments (see Sect. 5.3.2). The squares mark the reference positions on the FPA where the values have been obtained. **Left diagrams:** the relative R^2 (top) and e (bottom) predicted by the STOP analysis. **Right diagrams:** relative R^2 and e measured in-orbit in November 2023.

less prominent temperature excursions can all be associated with calibration observations causing changes in SAA, or a reduced VIS cadence in case of recurrent PSF calibration observations. The time history shows that the baseplate temperature remained nearly half of the time in transient periods, during the months September 2024 and December 2024, and the first 10 days of October 2024. This is predominantly caused by a few large SAA rotations, which were less than two weeks apart in time.

5.3. Image quality

5.3.1. Image quality after first alignment of the telescope

Soon after decontamination and cool-down of the PLM in July 2023, the telescope was aligned by adjusting the M2 mechanism which was previously set in launch condition. The whole alignment process was based on images obtained by the VIS instrument. Its objective was to optimise the R^2 image quality metric over the whole VIS FPA, with SAA = 90° and AA = 0°. A preliminary corrective motion of the M2 was carried out on the basis of the PSF shape of stars visible on previously acquired VIS images. This initial motion brought the M2 closer to its optimum position. Then, the three degrees of freedom (translation along optical axis, tip, and tilt) of the M2 mechanism have been scanned successively. Along each scan, acquisitions were performed, and the R^2 of the acquired stars spread over the whole field were computed. After each scan, the M2 position minimising the worst R^2 was computed and applied. After alignment, the CODE V optical model was updated to match the in-orbit measurements. The numerical simulations after updating the model showed the excellent optical quality of the telescope, well within the required performance: a worst case $R^2 = 0.051$ arcsec² and an average $R^2 = 0.047$ arcsec², over the nine reference field points, for a monochromatic source at 800 nm, to be compared with the required value of 0.055 arcsec² at spacecraft level (Ta-

ble 5), apportioned to 0.053 arcsec² at telescope level. A direct extraction of the telescope R^2 from the in-orbit acquisitions during the alignment process was also performed. The extraction includes unaccounted contributions from the colours of the stars and the detector PSF, which increase the value of R^2 . However, the result of this analysis was in line with the numerical simulations: a maximum $R^2 = 0.050$ arcsec² and an average $R^2 = 0.048$ arcsec², both within the requirements.

We determined the FWHM and ellipticity from a nominal science observation obtained in early September 2023 as part of the *Euclid* Early Release Observations (ERO) programme (Cuillandre et al. 2025). The observation of the dwarf irregular galaxy IC10 was selected because it is situated in a field at a low Galactic latitude ($b = 3^\circ$) with high stellar density. The average FWHM and ellipticity $|e|$ measured from a single exposure are listed in Table 5. Despite the additional uncertainties due to the colours of the stars and the detector effects, the in-orbit IQ performance appears to be well within the limits indicated by the STOP predictions and requirements.

5.3.2. Static IQ performance

We obtained the in-orbit IQ parameters R^2 , e_1 , and e_2 by measuring the quadrupole moments of selected stars detected in VIS frames, and by excluding the VIS short exposures of 95 s. To overcome an unacceptable high noise caused by cosmic rays, we reduced the width of the Gaussian weighting of $\sigma = 0''.75$ (see Sect. 2) to $\sigma = 0''.25$. This suppression of the wings of the PSF produced a systematically smaller value of R^2 . The resulting values for R^2 and ellipticity e cannot be directly compared to the required values but can be used as diagnostics for IQ variations. Details of the IQ determination are given in Appendix E.

We investigated the static IQ performance in the period from 2023-11-01 to 2023-11-26 by averaging the R^2 and e over time intervals of 12 h for the stars detected in 9 areas of the VIS FPA. The spatial variation of the mean IQ parameters across the FPA is significantly larger than the variation in each area during the measurement period. The relative IQ variations in the VIS FPA are visualised in Fig. 9 where the observed results are compared with the STOP predictions. The predicted IQ values were calculated in the reference field points defined in Fig. 2 where 8 of them are located at the edges of the VIS FPA to determine extreme design results. This is different from the central positions of the 9 areas (indicated in Fig. 9) which were obtained by spatial averaging of the IQ data. The in-orbit R^2 variation does not resemble the predicted distribution with a minimum R^2 at the centre of the FPA. The in-orbit variation in ellipticity shows better resemblance to the predicted distribution and the magnitude of the relative variation is similar.

5.3.3. Transient performance: IQ time histories

The time histories of R^2 , ellipticity components e_1 and e_2 , as well as the stellar colours averaged over 12 h time bins over the periods addressed in Sect. 5.2 are presented in Figs. 10 and 11. The selection of stars based on limited colour and magnitude suitable to derive the IQ parameters provided sufficient statistics only for time intervals of 12 h, shorter time intervals did not improve time resolution. The values of R^2_{norm} were obtained by normalising the timelines per FPA area by dividing the values by their mean value over the duration of the period. Subsequently, we derived the mean normalised R^2 of the 9 FPA areas weighted by the variances per time bin, see Appendix E for details. The

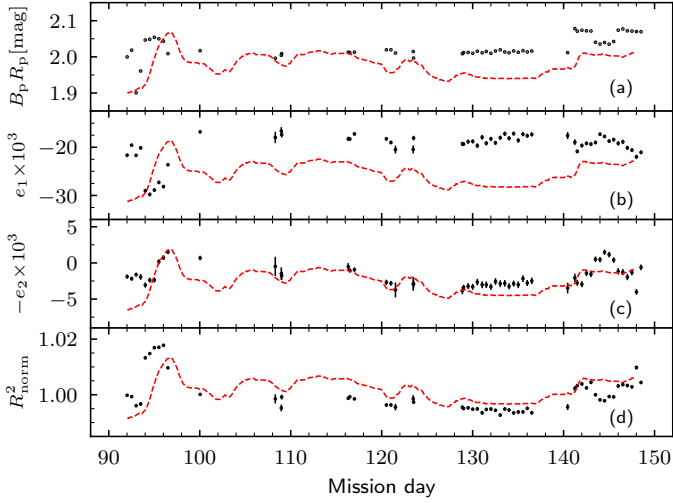


Fig. 10. Time history of the observed IQ metrics for the period shown in Fig. 7, before the start of the nominal mission, from 2023-10-01 (mission day 92) until 2023-11-26 (mission day 148). (a): the variation of the average *Gaia* colour $B_p R_p$ per 12 h bin. (b): the variation of the ellipticity component e_1 . (c): the variation of the ellipticity component e_2 . To underscore the correlation with baseplate temperature $-e_2$ is displayed. (d): the variation of R^2 . The red dashed lines indicate baseplate temperature T_{BP} (THM05) with arbitrary scaling for reference.

values for e_1 and e_2 are the average of the 9 FPA areas without further normalisation. The stellar colours serve as an indicator for possible colour dependencies on the IQ metrics. We plotted $-e_2$ to show the correlation better with temperature T_{BP} .

R_{norm}^2 exhibits a peak-to-peak variation of less than $\approx 2.2\%$ in both time histories. Variations in ellipticity components are substantially greater, up to a factor of 1.7 for e_1 and 3.6 for e_2 in the 2024 period. The ellipticity component e_1 is more than 3 times larger than e_2 , making it the dominant component of e . During the 2023 period, e_2 changed sign, indicating that e_2 can be zero during periods of time. In addition, the mean values of the ellipticity components over the two time periods are different, which could be due to differences between pre-launch and nominal survey conditions. Despite the observed variations, the FPA-averaged R^2 and $|e|$ remain within the required values assuming $R^2 = R_{\text{norm}}^2 \times R_{\text{in-orbit}}^2$, where $R_{\text{in-orbit}}$ is the in-orbit value provided in Table 5.

The time history of the 2024 period (Fig. 11) allows us to assess the IQ metrics during the routine operations over a long time interval without gaps in the VIS observations as large as seen in Fig. 10. T_{BP} exhibits 5 prominent peaks, which all can be associated with observations of the southern calibration field for spectroscopy (CPC-South), requiring large SAA rotations affecting T_{BP} . The variations of R_{norm}^2 in Fig. 11d roughly match those seen in T_{BP} , with higher R_{norm}^2 when T_{BP} is higher. Four of the five prominent peaks in T_{BP} are accompanied by peaks in R_{norm}^2 . However, higher values of R_{norm}^2 can also be associated with survey periods in which the stars are redder. During mission days 467 to 482 *Euclid* observed a region in the Deep Field North followed by a nearby survey patch at low Galactic latitude (close to $b = +30^\circ$) where the selected stars are on average ≈ 0.06 mag redder per bin than during the earlier and later observations. The period 510–515 can be associated with observations of the COSMOS calibration field, where the selected stars are ≈ 0.04 mag redder. During these periods R_{norm}^2 is higher by 0.3–0.5%. In addition, the second and third temperature peaks, occurring on mission days ≈ 442 and ≈ 456 , are accompanied by

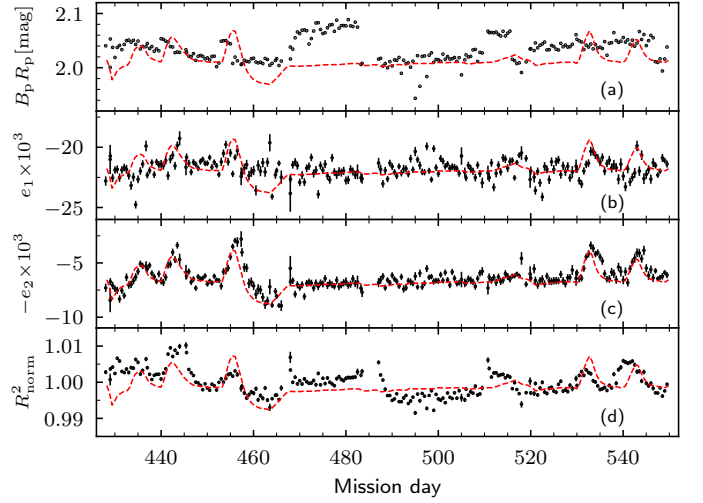


Fig. 11. Similar as Fig. 10 but for the IQ metrics observed in the period shown in Fig. 8, covering 4 months of nominal mission, from 2024-09-01 (mission day 428) until 2024-12-31 (mission day 549).

redder $B_p R_p$ of 0.02 and 0.03 mag, respectively. We infer that our preselection of stars in a limited colour range does not prevent a bias in the values of R_{norm}^2 due to observations of survey patches containing significantly redder stars with respect to the observations taken before and after. The colour variations could be reduced by a narrower range in $B_p R_p$ for the selection of stars, but then the statistical scatter in the IQ timelines would become too large. Fortunately, the temperature peaks detected in December 2024 on mission days ≈ 533 and ≈ 542 are not accompanied by significant colour excursions. Therefore, we attribute the peaks in R_{norm}^2 to variations in temperature T_{BP} .

The time history of e_2 in Fig. 11c shows a striking correlation with T_{BP} and no colour dependence. The ellipticity component $-e_2$ traces all the peaks in T_{BP} with the same relative amplitudes. It also closely follows the course of the baseline of T_{BP} . The correlation is less tight in Fig. 10c, and there might be an indication that e_2 is affected by the colour of the stars; see mission days 144 and 145. For e_1 , we observe that the strongest temperature peaks are accompanied by 10–15% higher values of e_1 , with no obvious colour dependency. However, Fig. 11b suggests a weaker correlation with T_{BP} than was observed for R_{norm}^2 and e_2 .

5.3.4. Other IQ metrics: trefoil

In addition to the quadrupole moments metrics, other third-order metrics were also visible, such as the trefoil, producing a “triangular” shaped PSF. A degree of trefoil was predicted as a result of the introduction of a special mechanical device to correct for a slight astigmatism by M1 during telescope development. The updated CODE V model tuned for sky observations predicts a wavefront error (WFE) contribution of the trefoil Zernike terms (z_{10} and z_{11}) of less than 14 nm rms in the worst position. This value is negligible in the WFE budget (≈ 50 nm) required to be well within the IQ performance in the 9 reference field points. We did not analyse the on-orbit dependency of the trefoil with variations in T_{BP} , we refer to Whittam et al. (in preparation) for a detailed investigation.

Table 6. Estimated correlations based on the analysis of the in-orbit data. In-orbit R^2 and e were extracted from data obtained in the period September to December 2024, averaged over intervals of 12 h. Temperature sensor THM05 was used for T_{BP} .

relation	in orbit	STOP prediction	comment
$\Delta T_{BP}/\Delta SAA$ (mK deg ⁻¹)	5–15	2.6	SAA >90° for STOP
$\Delta R_{norm}^2/\Delta T_{BP}$ (% K ⁻¹)	3–6	3.7	in-orbit value from peak on mission day 533
$\Delta R_{norm}^2/\Delta SAA$ (% deg ⁻¹)	0.03–0.05	0.01	estimated 30% uncertainty in R^2 due to colour SAA >90°; in-orbit peak on mission day 533
$\Delta e_{norm}/\Delta T_{BP}$ (% K ⁻¹)	50–70	70	estimated 30% uncertainty in R^2 due to colour variation of mean normalised ellipticity $ e $
$\Delta e_1/\Delta T_{BP}$ (K ⁻¹)	0.01	-	estimated from Fig. 11b
$\Delta e_2/\Delta T_{BP}$ (K ⁻¹)	-0.013	-	estimated from Fig. 11c
$\Delta e_{norm}/\Delta SAA$ (% deg ⁻¹)	0.5–0.7	0.2	SAA >90° for STOP;
Baseplate thermal time constant (h)	20	12–15	the larger STOP value was derived from system parameters updated close to launch
M1 thermal time constant (h)	33	33–35	

6. Discussion

6.1. Synthesis of thermal and structural features from the STOP analysis

Thermal analysis revealed the importance of the baseplate as the thermal regulator of the telescope. We note that PLM heaters are not used for the regulation of temperature variations but are controlled by open-loop settings to maintain its temperature constant. The temperatures of the main optical elements show a linear dependency with the average temperature of the baseplate, suggesting that the dominant temperature-driven telescope aberration is a global effect and not due to local distortions in the optical train or by individual mirrors. These findings confirm the purpose of the all-SiC homothetic design, where the entire telescope structure expands and contracts homogeneously with temperature. However, the linear dependence with the baseplate temperature is not uniform, indicating that the effect is not completely homothetic still causing optical aberration albeit on a lower level than for a non-homothetic design.

In the analysis, the Sun’s incidence angles, AA and SAA, are the temperature drivers. Within the stated limits of these angles, the maximum temperature fluctuation was found to be ≈ 0.3 K, while during the 7-year lifetime, the mean temperature was predicted to vary by ≈ 1.5 K. The EOL condition assumes a 7-year degradation of the surfaces responsible for thermal regulation, as well as the extremes of the seasonal variation. Hence, if all of the variation is attributed to a gradual degradation process, the STOP analysis predicts a continuous increase of the baseplate temperature of ≈ 200 mK/year. Temperature variation at this level can be monitored with sufficient accuracy since the temperature sensors have a resolution better than 10 mK.

Transient analysis revealed that the telescope temperature stability on the 11 000 s time scale has a long-term and a short-term component. The dominant component, with a long time constant of 12 hours, describes the heating/cooling of the PLM following the variation of the Solar aspect angle and of the internal dissipation. The fast (about 3 h) component is revealed in the derivative of the temperature variation of the baseplate. We understand the two terms as associated with the different time constants of the heavily insulated all-SiC PLM and the more exposed SVM, which experiences a rapid and large temperature excursion as the Sun moves below the horizontal plane. The fast component was shown to have a counterpart in the displacement of M2 and other mirrors, acting as a damped oscillation, with a

quick rise and a slower exponential return. Later optical analysis showed that this phenomenon under extreme conditions approaches the upper limit of the IQ stability boundary.

6.2. Observed thermal properties in-orbit

As to the spacecraft attitude, SAA changes affect the baseplate temperature, whereas the effect of AA is hardly noticeable if present at all. A large AA step of 5° did not provide evidence for a related T_{BP} variation (Sect. 5.2.1). Since the operational range in AA was reduced to 5°4, which is only 30% of the range considered in the STOP analysis, we do not expect significant baseplate temperature changes due to AA variations alone.

For the SAA, we observe an increase in T_{BP} when the SAA increases above 90°. For constant ΔSAA , the temperature increases approximately following a logistic function, reaching a plateau after about 4 days. Short-duration visits to high SAA of less than 12 h have little impact. The STOP analysis predicted an effect of SAA on baseplate temperature; however, the magnitude of the effect was 1–2 mK for $\Delta SAA = 10^\circ$, while an average of about 5–15 mK/deg is observed (see Table 6). This suggests a larger than predicted heat input from the SVM to the baseplate, possibly due to a slight overestimation of thermal insulation by the MLI devoted to the protection of the SVM bottom part against the Sun.

The operation of the instruments causes significant baseplate temperature variations. In the case of VIS, a higher cadence of the exposures leads to an increase of T_{BP} . For NISP, filter and grism wheel actuations are a source of power dissipation leading to T_{BP} fluctuations at a level of a few 10 mK. These effects were not studied before launch because the instrument dissipation was assumed constant. After they were detected, the STOP model was employed to verify the heat exchange between the VIS instrument components, the VIS radiator, and baseplate. To simulate the baseplate temperature variation, we applied a power reduction step to the FPA electronics (−0.5 W) and RSU (−0.2 W), corresponding to the extra power associated to a readout event. The subsequent evolution of the temperatures was monitored for six days. The model provided time constants and temperature variations in qualitative agreement with the orbit data.

6.3. Image quality

The first VIS measurements after the telescope alignment showed that the IQ parameters R^2 , FWHM, and ellipticity e are in line with the STOP predictions and meet the design requirements; see Table 5. The in-orbit IQ values averaged over the VIS FPA given in Table 5 indicate comfortable margins, without correcting for the spectral energy distribution of the stars (colours) and the known detector effects.

The STOP analysis indicated that the contribution of the structural deformation of the truss to image quality should be small, less than 0.5% (Sect. 4.3.1). The main cause of variations is the deformation of the mirrors due to temperature variations. From the analysis of the Zernike terms in the transient case it was found that the telescope aberration is mainly the defocus term (Sect. 4.3.2) and is proportional to the variation of the temperature of the baseplate (Sect. 4.1).

The time histories of the IQ metrics were determined from VIS frames obtained during early operations and during a period of four months during the nominal mission of *Euclid*. The IQ values are not directly comparable with the required values; therefore, relative comparisons were made (Sect. 5.3.2). The predicted R^2 distribution over the FPA is found to be significantly different from the observed distribution (Fig. 9). This suggests that the in-orbit wavefront error differs from the predicted WFE which was based on on-ground measurements of the telescope's optical parameters. The observed ellipticity distribution in the FPA shows a better match with the notion that the in-orbit e is dominated by e_1 . For both IQ parameters, the observed relative variations over the FPA are smaller and within the peak-to-peak variations predicted by STOP suggesting that the deviations are within the requirements.

The in-orbit variations of the IQ parameters R^2_{norm} , e_1 , and e_2 exhibit a dependence on T_{BP} (Figs. 10 and 11). We found that our measurement of R^2_{norm} is biased by the mean colour of the stars in the field, an effect we could not remove by applying a narrower colour selection without losing statistical significance, but we can identify R^2 variations without colour bias. The tight correlation between e_2 and T_{BP} and to a lesser extent the correlation of R^2 and e_1 with T_{BP} highlight the role of T_{BP} and the thermo-optical time constants in IQ, as predicted by the STOP analysis. The estimated dependencies are given in Table 6. Assuming that during the wide survey the instrument activities have a constant cadence, variations in image quality of science data can mainly be attributed to variations in SAA, when considering stars with the same colour. We note that all significant baseplate temperature transients during the nominal survey period investigated by us are triggered by deep-field or calibration observations that require large SAA rotations and observation periods longer than 12 h. Large SAA excursions with durations shorter than 6 h, seen during spacecraft maintenance periods, caused no significant IQ variations.

We were unable to verify the stability requirements in-orbit to the required stability time resolution of 11 000 s due to the statistical noise in our measurements over these time intervals. The minimum time interval considered by us is 12 h. The observed variations after large SAA steps indicate that the resulting peak values of R^2 , e_1 and e_2 are reached after ≈ 48 hours. Assuming a pessimistic $\Delta T_{\text{BP}}/\Delta \text{SAA} = 15 \text{ mK/deg}$ (Table 6), we obtain $\Delta e_1/\Delta \text{SAA} = 2 \times 10^{-4}$. An instantaneous SAA step of 15° would result in a continuous e_1 increase of 3×10^{-3} over 48 hours. During that period, the stability science requirement is violated (2×10^{-4} over 11 000 s, see Table 1), but not the spacecraft stability requirement (2×10^{-3} over 11 000 s). For $\Delta R^2/R^2$, the

science requirement over time intervals of 11 000 s would also be violated for a slew of 15° , but not the spacecraft requirement, which would be violated if the change happened in 24 h instead of 48 h. The estimated variations of e_2 and R^2 could be more severe in some intervals, but we do not have the time resolution. Our estimation shows that the margin in the R^2 stability requirement is good to within a factor of 2.

7. Conclusions

We have presented a mathematical model, which includes thermal, structural, and optical components, to ensure that the design of the *Euclid* spacecraft can support stringent image quality requirements to meet its scientific objectives. The model was adapted to the development of the satellite and was refined to predict the in-orbit performance of the final design of the satellite, confirming that the requirements would be fulfilled during the lifetime of the mission. The model was used to investigate the response of the telescope to extremes in the expected thermal environment, also accounting for ageing during the mission. It showed that the defocus is the dominant aberration, and is due to the overall deformation of the optical train and mirror surfaces, with a linear dependency on the temperature of the telescope baseplate.

We compared prelaunch predictions with in-orbit performance data obtained from housekeeping telemetry and VIS exposures. Changes in SAA are confirmed to be the dominant cause of baseplate temperature fluctuations. We observe a temperature increase as large as 400 mK when SAA moves to 120° and remains there for 4 days or more. AA rotations cause no obvious temperature variations as long as they are confined to a smaller range in the negative AA domain, as is the case in the present survey. The SAA dependency is predicted by STOP but with a much smaller amplitude. We attribute the observed behaviour to a possible under-performance of the MLI attached to the bottom of the SVM. However, it should be noted that the deviation from the prediction is extremely small in absolute terms and completely compatible with the uncertainties associated with a thermal analysis based on state-of-the-art mathematical models used in the space industry.

The in-orbit time constants are found to be around 1 day for the baseplate and 33 hours for the primary mirror M1. These values are significantly longer than predicted by the STOP model and are likely due to numerous worst-case assumptions, including less efficient thermal insulation between the top side of the SVM and baseplate, adopted to make the design robust.

In addition to changes in solar attitude, activities by the VIS instrument are the second dominant cause of baseplate temperature variations, which can be up to 100 mK under nominal survey operations. These changes can be associated with the cadence of the readouts. NISP activities, which were monitored by the instrument's filter wheel movements, have a lower impact on the baseplate temperature. We note that time-dependent NISP activities were not considered in the STOP analysis.

Monitoring in-orbit image quality parameters R^2 and ellipticity e over two periods during the mission, including a continuous survey period of four months, confirms the dependence of both parameters on the baseplate temperature, as predicted by the STOP model. Although the ellipticity varies according to the prediction and remains within the requirements, the dependence of R^2 on SAA appears significantly higher. Applied steps in SAA of more than $\approx 10^\circ$ – 20° , could cause a violation of the IQ stability requirement for periods of 1 to 2 days, which corresponds to 20–40 survey observations. We were unable to ver-

ify the predicted compliance of the stability requirement during shorter periods, due to the statistical noise in our measurement method.

We conclude that the STOP analysis proved useful during the design phase in correctly estimating the observed behaviour, to the extent needed during development. In this paper, we provide the first quantitative assessment of the measured effects of environmental parameters on image quality. The observed in-orbit deviations, while confirming the good design choices, deserve further investigation.

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References

- Amara, A. & Refregier, A. 2008, MNRAS, 391, 228
- Bougoin, M., Castel, D., & Levallois, F. 2012, in SPIE Conference Series, Vol. 10564, International Conference on Space Optics (ICSO) 2012, ed. B. Cugny, E. Armadillo, & N. Karafolas, 1056410–1
- Bougoin, M., Mallet, F., Levenac, J., et al. 2018, in SPIE Conference Series, Vol. 11180, International Conference on Space Optics (ICSO) 2018, ed. Z. Sodnik, N. Karafolas, & C. Bruno, 111801P
- Cropper, M., Hoekstra, H., Kitching, T., et al. 2013, MNRAS, 431, 3103
- Cuillandre, J.-C., Bertin, E., Bolzonella, M., et al. 2025, A&A, 697, A6
- Euclid Collaboration: Cropper, M., Al-Bahlawan, A., Amiaux, J., et al. 2025, A&A, 697, A2
- Euclid Collaboration: Jahnke, K., Gillard, W., Schirmer, M., et al. 2025, A&A, 697, A3
- Euclid Collaboration: Mellier, Y., Abdurro’uf, Acevedo Barroso, J., et al. 2025, A&A, 697, A1
- Euclid Collaboration: Scaramella, R., Amiaux, J., Mellier, Y., et al. 2022, A&A, 662, A112
- Gaspar Venancio, L., Laureijs, R., Racca, G. D., et al. 2022, in SPIE Conference Series, Vol. 12777, International Conference on Space Optics (ICSO) 2022, ed. K. Minoglou, N. Karafolas, & C. Bruno, 127776H
- Gottero, M., Costa, G., Ferrero, S., et al. 2024, in International Conference on Environmental Systems ICES, Vol. 53, ICES-2024-234
- Laureijs, R. J., Vavrek, R., Racca, G., et al. 2024, in SPIE Conference Series, Vol. 13092, Space Telescopes and Instrumentation 2024: Optical, Infrared, and Millimeter Wave, ed. L. E. Coyle, S. Matsuura, & M. D. Perrin, 130920M
- Lorenzo Alvarez, J., Metselaar, H., Amiaux, J., et al. 2016, in SPIE Conference Series, Vol. 9911, Modeling, Systems Engineering, and Project Management for Astronomy VII, ed. G. Z. Angeli & P. Dierickx, 99110C
- Meneghetti, M., Melchior, P., Grazian, A., et al. 2008, A&A, 482, 403
- Paulin-Henriksson, S., Amara, A., Voigt, L., Refregier, A., & Bridle, S. L. 2008, A&A, 484, 67
- Paulin-Henriksson, S., Refregier, A., & Amara, A. 2009, A&A, 500, 647
- Racca, G. D., Laureijs, R., Stagnaro, L., et al. 2016, in SPIE Conference Series, Vol. 9904, Space Telescopes and Instrumentation 2016: Optical, Infrared, and Millimeter Wave, ed. H. A. MacEwen, G. G. Fazio, M. Lystrup, N. Batalha, N. Siegler, & E. C. Tong, 99040O
- Schneider, P., Kochanek, C., & Wambsganss, J. 2006, in Saas-Fee Advanced Course, Vol. 33, Gravitational Lensing: Strong, Weak & Micro, ed. G. Meylan, P. Jetzer, & P. North
- Vavrek, R. D., Laureijs, R., Lorenzo Alvarez, J., et al. 2016, in SPIE Conference Series, Vol. 9911, Modeling, Systems Engineering, and Project Management for Astronomy VII, ed. G. Z. Angeli & P. Dierickx, 991105

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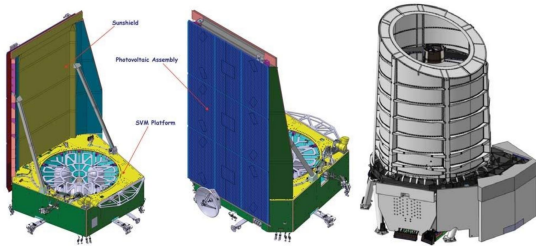


Fig. A.1. *Euclid* spacecraft overview. From left to right: back view of the SVM; front view of the SVM; the integrated PLM (not on the same scale as the SVM). The sunshield is mounted to the SVM structure and shields the PLM structure from direct sunshine for all allowed spacecraft attitudes. The side of the PLM exhibits three thermal radiators for the cooling of the VIS, instrument electronics harness, and NISP, from left to right, respectively. The telescope baffle is also used as thermal radiator. Three bipods are used to attach the PLM onto the SVM thereby providing maximum conductive thermal insulation.

Appendix A: *Euclid* design and development summary

In this paper, we consider three major satellite components: the service module (SVM) without sunshield, the payload module (PLM), and the sunshield (Fig. A.1). By design, the sunshield is part of the SVM and it is mounted to the SVM with no mechanical contact with the PLM. The PLM contains the Korsch three-mirror anastigmat (TMA) telescope and the two science instruments VIS and NISP as part of the optical train. All main PLM components are mounted on the baseplate, which acts as a two-sided optical bench for the optical components and telescope baffle. Using a system of thermal radiators and heaters, including the telescope baffle that functions as a radiator, the PLM instrument cavity is maintained at a required temperature of 135 K.

At the design level, many countermeasures were taken to isolate the PLM, where the telescope and instruments are located, from the SVM, where the ancillary subsystems reside and where most of the heat from the spacecraft is dissipated. The isolation was achieved by using high-efficiency thermal insulation and employing a set of three identical bipods, each of them composed of two glass fibre reinforced plastic struts, ensuring high stiffness to withstand the launch environment and low thermal conductance and thermal expansion.

The PLM telescope structure is based on an all-SiC concept; see [Euclid Collaboration: Mellier et al. \(2025\)](#). Sintered Silicon Carbide offers special properties for the construction of a stable telescope ([Bougoin et al. 2012, 2018](#)). In particular its high steady state stability, i.e. the material ratio between the thermal conductivity and the coefficient of thermal expansion allows, in principle, the construction of a telescope that behaves homothetically under temperature changes. The high thermal conductivity ($180 \text{ W m}^{-1} \text{ K}^{-1}$ at room temperature) minimises the thermal gradients across the structure and the low coefficient of thermal expansion (CTE; 2.2 ppm K^{-1} at room temperature and 0.3 ppm K^{-1} at operational temperature) ensures that no optical aberrations occur for limited temperature changes. This isothermality is also enhanced by maximising thermal diffusivity (thermal conductivity over heat capacity) as is the case for SiC. In practice, the optical elements and some instrument parts are connected to the SiC structure by interposed elements made of Invar and glue, whose CTE is similar but not identical to that of SiC. This realisation together with a non-uniform temperature distribution

introduces some residual optical aberrations, which are the subject of this study.

Another important element for stability is the control of the heat input into the PLM, which is isolated from the SVM by design. The two sources of power into the PLM instrument cavity are the VIS and NISP instruments. The VIS focal plane is mounted directly at the position of the telescope focal plane ([Euclid Collaboration: Cropper et al. 2025](#)). The heat dissipated in the focal plane read-out electronics is transferred to an external 1 m^2 radiator, while the 36 CCDs are coupled to the SiC baseplate via thermal straps. The other VIS dissipating unit is the Readout Shutter Unit (RSU) which dumps its internal heat directly into the baseplate via radiation and conductance. NISP is instead a stand-alone instrument mounted on the baseplate with three bipods ([Euclid Collaboration: Jahnke et al. 2025](#)) mounted directly on the telescope baseplate. The average internal dissipation during the actuation of the filter and grism wheels is low and is damped directly to the baseplate through the bipods. The detection electronics instead requires high-efficiency thermal coupling to the baseplate, due to significant dissipation (4.8 W) and low thermal gradient to be maintained with the baseplate; therefore, methane heat pipes are used. The NISP focal plane is the coldest part of the PLM, which must be kept below 100 K. This is achieved by thermal coupling to the 2 m^2 NISP radiator with 4 thermal straps. The heat dissipated by the instruments and its impact on the telescope stability is determined by the instrument operation cadence and the thermal design performance and is analysed in this paper.

The PLM temperature control of the instrument cavity during science observations is carried out by setting the appropriate power of four heater groups in open loop to achieve the required temperatures at the specified instrument interface locations. The baseplate is not a controlled item, but its temperature is determined by the heater power settings and the environmental conditions of the spacecraft.

The 3σ relative pointing error is required to be better than 25 mas for a fiducial VIS exposure time of 700 s. This is achieved with the fine guidance sensor (FGS) ([Racca et al. 2016](#)) using a cold gas micro-propellant system (MPS) for the actuation, which mainly provides a timely correction against the Solar pressure. Spacecraft vibrations are minimised by forcing the reaction wheels to stop spinning during exposures. These design measures ensure stable IQ during an exposure. Longer term IQ variations can only be caused by structural deformations due to changes in the thermal environment of the telescope or by degradation of the mirror surfaces. The latter effect is expected to occur on much longer timescales.

The thermal and structural design of the integrated spacecraft was verified by developing a structural thermal model (STM) of the *Euclid* satellite. This mechanical model of the spacecraft was partly assembled from the qualification models of the spacecraft components (SVM, PLM, sunshield, and instruments) and tested in 2019. The STM underwent more extreme and risky thermal and structural testing conditions, which could not be applied to the flight model. The tests were adapted for the correlation with the mathematical models.

The PLM flight model was assembled in Airbus Defence and Space (ADS) in Toulouse and tested in the Centre Spatial de Liège (CSL) 2021, where it was placed under thermal vacuum to simulate the ambient conditions in orbit ([Gaspar Venancio et al. 2022](#)). The PLM thermal design with the aid of the thermal model was verified, and the in-orbit mathematical telescope model was derived from the model correlations obtained during this campaign. The thermal analysis performed before launch

predicted an average $T_{BP} = (131 \pm 4)$ K, including ± 3 K uncertainty. During the in-orbit phases examined in this paper the average T_{BP} has been (131.9 ± 0.3) K. This achievement shows that the PLM thermal control design performs extremely well and according to the predictions.

The *Euclid* development consisted of a number of major common development milestones for the different system components. In this paper we refer to milestones for the spacecraft prime contractor: the Preliminary Design Review (PDR) held in Dec. 2015, the Critical Design Review (CDR) held in Nov. 2018, and the Qualification Review (QR) held in Dec. 2022. These milestones are accompanied by reviews where the requirements, implemented design, and expected performance are evaluated by independent panel members.

After the launch of *Euclid* on 1 July 2023, the plan was to devote the first month to the spacecraft commissioning phase followed by a 2-month performance verification phase before the start of the scientific programme. In practice various issues required the extension to 1.5 months of the spacecraft commissioning phase, followed by four months of performance verification phase that comprised also an extensive test campaign to verify the system software changes that were necessary to improve the pointing performance when observing sparse stars fields and in the presence of cosmic rays. This meant that scientific observations began only in December 2023. During the spacecraft commissioning phase, the telescope was put in focus by adjusting M2. The IQ parameters were verified during the performance verification phase.

Appendix B: Software applications to facilitate STOP analysis

B.1. MaREA

Two different approaches are available in MaREA for transient thermo-elastic analysis. In the direct approach, an interpolation function $F(\text{TMM}, \text{FEM})$ is created. At each time step, the temperatures are transferred from the TMM to the FEM, applying the F function, and a NASTRAN simulation is performed to calculate the displacements. The second approach makes use of an analytical procedure called TEMASE (Temperature Mapping Sensitivity). A fixed number of unit- δT thermo-elastic load cases, equal to the number of thermal nodes, is run to create the so-called “influence matrix”. A time sequence of displacements is obtained by multiplying the influence matrix by the matrix collecting all temperature time histories pertaining to the thermal nodes. The direct method is used once, to compute the elements of the influence matrix, while the matrix multiplication method is used to calculate the displacement time histories. Synthesising, the process:

- acquires temperature data from thermal analysis;
- establishes the correspondence between thermal and structural nodes and writes temperatures to the latter;
- loads each structural part with a unit- δT , performs thermo-elastic analysis, and calculates the influence matrix elements;
- calculates the displacement time histories by matrix multiplication;
- transfers the computed displacements, runs the CODE V optical model, and performs post-processing.

The most critical steps in the entire STOP process are the temperature interpolation from TMM to FEM and the transfer of displacements from FEM to CODE V. Both processes are executed

automatically as part of the workflow. The temperature interpolation is based on an inverse distance weighting algorithm. To minimise the interpolation error, both TMM and FEM were split into several substructures, and the interpolation algorithm was applied to each one of them. After this step, each thermal node has a finite-element model counterpart.

B.2. Using CODE V for the optical analysis

IQ variations are caused by the combination of the displacements of the optical surfaces attachment points and the thermal expansion of the curved mirrors. The first effect was introduced into the CODE V model of the telescope as output from the NASTRAN FEM run under several thermal conditions, as explained in Sect. 4. For the second effect, a transfer function was defined, allowing us to derive the WFE variations from the mean thermal variation of each optical surface. In synthesis, the definition of the transfer function is based on the FEM of the telescope, including the detailed FEM for optics and all SiC and Invar parts. A thermal load was applied to create a uniform temperature increase as large as 100 K to get a significant change in order to avoid FEM computation noise. The displacement of each optical surface and their deformations were integrated in the CODE V model. Once the global check was completed, the Zernike coefficients were calculated for all optical surfaces of interest. The resulting delta Zernike was then divided by 100 to get the sensitivity coefficient normalised to 1 K. The second step was to add to the CODE V model the WFE variation caused by the mirror temperature changes. The combined WFE variation was obtained from the first 36 Zernike coefficients (using Fringe’s indexing):

$$dz = (dz_1, \dots, dz_{36}). \quad (\text{B.1})$$

A quadratic model function of the Zernike variations was derived. This model was built by generating PSFs from telescopes with different WFE content using the CODE V model. The IQ metrics for each of these PSFs were computed and their variations as a function of Zernike variation were derived by a polynomial fitting. The outcome is one matrix of 36×36 coefficients per IQ metric. The variation of a given metric M (R^2 , Q_{xx} , Q_{xy} , Q_{yy} , and FWHM) is computed from

$$dM = \sum_{i=1}^{36} \sum_{j=1}^{36} a_{i,j}(M) (2 z_i dz_j + dz_i dz_j), \quad (\text{B.2})$$

with $a_{i,j}$ the metric sensitivity coefficients, z_i the baseline Zernike coefficients, computed with the CODE V model, and dz_i the variation of the Zernike coefficient. The final IQ metric is obtained from:

$$M = M_0 + dM, \quad (\text{B.3})$$

where M_0 is the metric baseline value, computed from the PSF generated by the CODE V model. The assessment of dM remains accurate for low aberration. The error assessment of this method was evaluated with a Monte-Carlo analysis.

Appendix C: Micro-propulsion boom redesign

Excluding the thrusters, the boom initial design was 470 mm long and would be constantly exposed to the Sun. For $\text{SAA} \leq 90^\circ$, the MLI of the boom reached temperatures above 20°C and illuminated directly the VIS radiator. The resulting maximum T_{BP} excursion was 1.4°C assuming EOL conditions (see

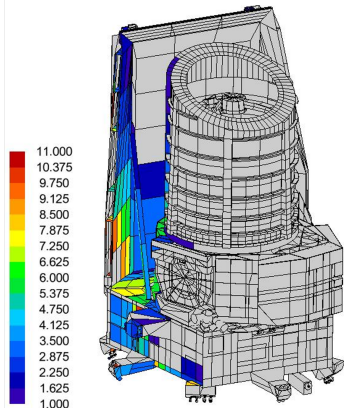


Fig. C.1. Temperature differences between Sun azimuth -8° and $+8^\circ$ at constant SAA. Grey areas are outside the colour code range. Modified MPS boom configuration.

Fig. 3a). Taking into account a 5% change in ellipticity per degree of T_{BP} , this would cause an ellipticity change of approximately 7%, violating the IQ requirement for ellipticity.

We reconfigured the MLI on the topside of the boom such that it remained in shadow even at minimum SAA. This caused the topside boom temperature to remain nearly constant when moving the Sun from the $-Y_{SC}$ side to the $+Y_{SC}$ side. However, the sides of the boom facing $+X_{SC}$ experienced significant temperature variations. T_{BP} was found to be almost independent of SAA, but continued to show a dependence on AA. The maximum T_{BP} excursion was halved to 0.7°C still violating the ellipticity requirement. The Sun no longer illuminates the boom topside, but the sides facing the Sun reach temperatures of about 200°C . In the enclosure formed by the sunshield $+Y + Z$ wing, the PLM side facing it, and the SVM topside (all surfaces with a non-zero view factor to the Sun-facing side of the boom), the MLI temperature increases by about 10°C , see Fig. C.1. The PLM baffle strip on the $+Y$ side, unprotected by MLI, undergoes a temperature raise of about 0.5°C . As the PLM interior is directly coupled to the baffle, about the same effect is detected for the baseplate.

The next option was to reduce the length of the MPS boom, the trade-off being the reduction of torque authority around the $+Z_{SC}$ axis. The minimum feasible length was determined to be 200 mm. We still found dependencies on both AA and SAA (Fig. 3b) but the maximum T_{BP} excursion at EOL was further reduced to 0.6°C . The simulation of the case without boom resulted in a maximum T_{BP} excursion of 0.5°C at EOL. This showed that the half-length boom comes very close to cancelling the effects of the boom altogether. Based on this study, the design of the spacecraft was updated, incorporating the short boom.

Appendix D: Time dependent IQ computations

The 70 h variation of the IQ metrics in the field points is shown in Fig. D.1 for the ellipticity and in Fig. D.2 for R^2 . As explained in Sect. 4.3.2, the long-term evolution driven by defocus and rapid initial variation driven by coma and astigmatism are the dominant components of the derived variations.

Appendix E: IQ determination from in-orbit VIS data

To overcome the high glitch rate due to cosmic rays hitting the VIS sensors, we decreased the standard deviation of the Gaussian weighting function from $\sigma = 0''.75$ to $0''.25$ when calculating

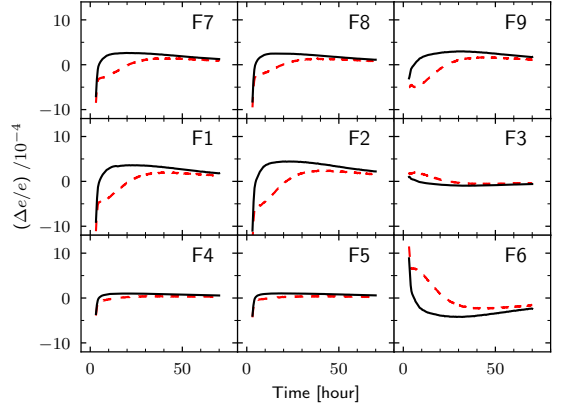


Fig. D.1. 11 000 s variation of ellipticity δe in the nine field points as function of time. Dashed lines: telescope only; solid lines: telescope and mirrors.

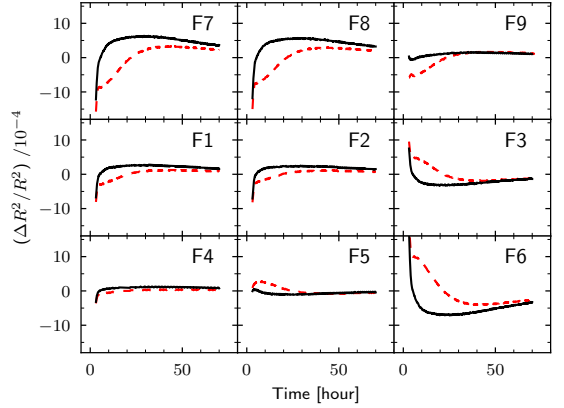


Fig. D.2. 11 000 s variation of $\delta R^2/R^2$ in the nine field points as function of time. Dashed lines: telescope only; solid lines: telescope and mirrors.

the quadrupole moments for the derivation of the image quality parameters (see Sect. 2) for stars in the VIS frames. This suppression of the wings of the PSF causes a systematically smaller value of R^2 . The IQ values we derived were still sensitive to environmental changes, so we can use them as diagnostic for changes in IQ. In our analysis, we removed VIS frames that were affected by strong solar flares using a merit function derived from e_1 , e_2 , and R^2 covariances. However, we found that some areas in the FPA are more affected by solar X-rays than others because of selective shielding by the sunshield (Laureijs et al. 2024). For the value of R^2 averaged over the FPA we therefore took the weighted mean per FPA area. This weighting was not performed to obtain the average value of ellipticity components due to the large intrinsic variation of the ellipticity (a factor of approximately 5) over the FPA.

The large width of the VIS band causes a strong PSF dependency on the colour of the stars in the field. From the inspection of individual VIS frames we estimate a $\approx 4\%$ increase of R^2 per magnitude of $G_{BP}R_p$. Furthermore, the PSF is affected by the brighter-fatter effect, inherent in the usage of CCD sensors. To minimise these effects we only considered stars that were also detected by *Gaia* within the colour range $1.5 < B_p R_p < 2.5$ providing approximately flat spectral energy distribution in the VIS band, and with a *Gaia* magnitude in the range $18.9 < m_G < 19.9$ to constrain IQ variations due to stellar brightness dependencies. The PSF measurements were binned in 9 equal areas on the VIS FPA to investigate the in-orbit IQ variations over the focal plane.