

Classical Cepheids in the Galactic thin disk.

I. Abundance gradients through NLTE spectral analysis

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ABSTRACT

Classical Cepheids (CCs) have long been considered excellent tracers of the chemical evolution of the Milky Way’s young disk. We present a homogeneous, Non-Local Thermodynamical Equilibrium (NLTE) spectroscopic analysis of 401 Galactic CCs, based on 1,351 high-resolution optical spectra, spanning Galactocentric distances from 4.6 to 29.3 kpc. Using PySME with MARCS atmospheres and state-of-the-art grids of NLTE departure coefficients, we derive atmospheric parameters and abundances for key species tracing multiple nucleosynthetic channels (O, Na, Mg, Al, Si, S, Ca, Ti, Mn, Fe, Cu). Our sample—the largest CC NLTE dataset to date—achieves high internal precision and enables robust modeling of present-day thin-disk abundance patterns and radial gradients. We estimate abundance gradients using three analytic prescriptions (linear, logarithmic, bilinear with a break) within a Bayesian, outlier-robust framework, and we also apply Gaussian Process Regression to capture non-parametric variations. We find that NLTE atmospheric parameters differ systematically from LTE determinations. Moreover, iron and most elemental abundance profiles are better described by non-linear behavior rather than by single-slope linear models: logarithmic fits generally outperform simple linear models, while bilinear fits yield inconsistent break radii across elements. Gaussian Process models reveal a consistent outer-disk flattening of [X/H] for nearly all studied elements. The [X/Fe] ratios are largely flat with Galactocentric radius, indicating coherent chemical scaling with iron across the thin disk, with modest positive offsets for Na and Al and mild declines for Mn and Cu. Finally, Cepheid kinematics confirm thin-disk orbits for the great majority of the sample. Comparison with recent literature shows overall agreement but highlights NLTE-driven differences, especially in outer-disk abundances. These results provide tighter empirical constraints for chemo-dynamical models of the Milky Way and set the stage for future NLTE mapping with upcoming large spectroscopic surveys.

Key words. Classical Cepheids – Galaxy formation and evolution – Spectroscopy – Chemical Abundances – ...

1. Introduction

Radial abundance gradients play a key role in constraining Galactic chemo-dynamical models (Palla et al. 2020; Spitoni et al. 2023; Lemasle et al. 2022). The radial trends observed in the present-day chemical abundances of the thin disk, when

compared with theoretical predictions across different Galactocentric distances and stellar ages, allow us to constrain its chemical enrichment history (Chiappini et al. 2001). Furthermore, radial gradients and their steady flattening as a function of time soundly support stellar radial migrations in shaping the current metallicity distribution function of the thin disk (Hou et al. 2000;

Prantzos et al. 2023). Moreover, they also provide an independent constraint on the star formation efficiency when moving from the innermost to the outermost disk regions (Spitoni et al. 2023).

Radial gradients across the thin disk have been investigated by using different stellar tracers: open clusters (Friel et al. 2002; Magrini et al. 2023; Myers et al. 2022; Carbajo-Hijarrubia et al. 2024; Otto et al. 2025; Dal Ponte et al. 2025), planetary nebulae (Maciel & Costa 2013; Stanghellini & Haywood 2018), HII regions (Fernández-Martín et al. 2017), and OB stars (Daflon & Cunha 2004; Nieva & Przybilla 2012; Bragança et al. 2019). Classical Cepheids (CCs) have also been extensively used to investigate metallicity distributions and chemical gradients in the Galactic thin disk (Yong et al. 2006; Sziládi et al. 2007; Romaniello et al. 2008; Pedicelli et al. 2010; Luck et al. 2011; Luck & Lambert 2011; Lemasle et al. 2013, 2017; Genovali et al. 2013, 2014, 2015; Korotin et al. 2014; Martin et al. 2015; Andrievsky et al. 2016; da Silva et al. 2016, 2022, 2023; Proxauf et al. 2018; Luck 2018; Inno et al. 2019; Kovtyukh et al. 2022; Ripepi et al. 2022; Trentin et al. 2023, 2024a, 2024b, 2025). Cepheids are radially pulsating variable stars that serve as excellent distance indicators, since their individual distances can be estimated with an accuracy of 1-3%. They also trace young stellar populations (Bono et al. 2024). CCs are central helium-burning, intermediate-mass stars younger than a few hundred million years, with pulsation periods ranging from days to roughly one year. They are ubiquitous across the thin disk and the Galactic center (Matsunaga et al. 2011; Bono et al. 2024).

Since CCs are typically distributed across the thin disk, their intrinsic colors are affected by uncertainties due to reddening corrections. This is the main reason why atmospheric parameters are estimated directly from the spectra. Several approaches have been proposed to address this issue. In particular, Proxauf et al. (2018) and da Silva et al. (2022) adopted the line depth ratio method (Kovtyukh 2007) based on empirical calibrations of different element pairs to derive effective temperature (Kovtyukh 2007; Elgueta et al. 2024). A new temperature scale for Galactic CCs based on a data-driven, machine-learning technique applied to observed spectra, was recently suggested by Lemasle et al. (2020) in which the flux ratios of different spectral features were tied to the effective temperatures derived using the infrared surface-brightness method (Hanke et al. 2018).

A significant fraction of abundance analysis of CCs is based on Local Thermodynamic Equilibrium (LTE) atmosphere models, such as those of Castelli & Kurucz (2004, ATLAS) and Gustafsson et al. (2008, MARCS). However, both theoretical predictions and empirical evidence show that LTE modeling can be prone to possible systematics (Lind & Amarsi 2024; Bergemann & Hoppe 2025). These systematics become more relevant when moving from dwarf stars to giants, and from metal-rich to metal-poor populations (e.g., Thévenin & Idiart 1999; Idiart & Thévenin 2000; Bergemann et al. 2012; Hansen et al. 2013; Fabrizio et al. 2021). Moreover, NLTE corrections are element- and line-dependent, with neutral lines more affected than ionized lines (Kiselman 2001; Collet et al. 2005; Merle et al. 2011; Tautvaišienė et al. 2015; Duffau et al. 2017).

Chemical abundances for CCs based on a NLTE approach have been provided by Andrievsky et al. (2016) for all their analyzed elements, by Luck et al. (2011) for CNO elements, by Martin et al. (2015) for Fe and O, and by Korotin et al. (2014) for O. However, in these investigations the stellar atmospheric parameters were derived by using an LTE approach.

Recently, Amarsi et al. (2020) provided NLTE departure coefficients for H, Li, C, N, O, Na, Mg, Al, Si, K, Ca, Mn and

Ba. Additionally, NLTE departure coefficients are provided by Mallinson et al. (2024) for Ti, by Amarsi et al. (2022) for Fe, by Caliskan et al. (2025) for Cu and by Amarsi et al. (2025) for S. These departure coefficient are computed on a grid of 3756 1D MARCS model atmospheres (Gustafsson et al. 2008) covering $3000 < T_{\text{eff}}/\text{K} < 8000$, $-0.5 < \log(g/\text{cm s}^{-2}) < 5.5$, and $-5 < [\text{Fe}/\text{H}] < +1$ dex. These comprehensive grids allow proper spectral modeling and NLTE analysis.

In this study, we apply a complete NLTE approach—deriving both atmospheric parameters and chemical abundances—for Galactic CCs by using high-resolution optical spectra. Although the statistical sampling in the inner (Galactocentric distance, $R_{\text{GC}} \lesssim 5\text{kpc}$) and in the outer ($R_{\text{GC}} \gtrsim 20\text{kpc}$) disk is sparse, there is solid empirical evidence for radial gradients for most of the elements that have been investigated. More specifically, our dataset spans a wide range in Galactocentric distances ($R_{\text{GC}}=5\text{--}29\text{ kpc}$) and provides homogeneous NLTE abundances for light (O), odd-Z (Na, Al, Cu), α (Mg, S, Si, Ca, Ti), and iron peak (Mn, Fe) elements. The current analysis is a stepping stone for exploiting thousands of new high-resolution spectra for CCs that will be collected from upcoming large spectroscopic surveys in optical (de Jong et al. 2016, 4MOST) and near-infrared (Cirasuolo et al. 2012, MOONS). This paper is organized as follows: in § 2, we present the spectral dataset and the line list we adopted for the spectral analysis. In § 3 we show the results concerning the atmospheric parameters and the chemical abundances, including the gradients, and the kinematic properties. The summary of the results, the discussion and a brief outline of the near future developments of this project are presented in § 4. The paper also includes several appendices. Appendix A presents the tables included in the paper, while Appendix B discusses the new distances of two CCs. Appendix C provides details of the adopted spectral synthesis, and Appendix D presents the validation of NLTE atmospheric parameters and chemical abundances.

2. Dataset and method

The current spectroscopic sample includes 401 CCs distributed across the thin disk, as shown in Fig. 1¹. Among them, 379 CCs have already been collected and discussed in da Silva et al. (2023) (hereinafter *ds23*). We added 66 High-Resolution (HR) spectra for 22 CCs (proprietary plus public archives): 10 with $R_{\text{GC}} \lesssim 6\text{ kpc}$ and 4 with $R_{\text{GC}} \gtrsim 10\text{ kpc}$.

2.1. Spectral sample

In the spectroscopic sample presented in *ds23*, individual heliocentric distances were estimated using the following methods: Gaia DR3 parallaxes (72%), W1-band Period-Luminosity (PL) relations (25%), K-band PL (2%), and J-band PL (1%). This ranking was adopted in *ds23* to favor purely geometrical distances over distances from PL relations. On the other hand, the ranking of PLs is W1-K-J to minimize the impact of extinction, which is less severe at redder passbands. The *ds23* sample includes 1,285 HR spectra and more than 80% have a signal-to-noise ratio per pixel greater than 100. For a comprehensive description of the *ds23* spectroscopic sample the reader is referred to da Silva et al. (2022) and *ds23*.

The *ds23* dataset was complemented with 22 CCs for which 40 HARPS-N, 10 UVES and 16 ESPaDOnS spectra were available from public archives: 35 HARPS-N are proprietary spec-

¹ NASA/JPL-Caltech/R. Hurt (SSC/Caltech), <https://milkyway-plot.readthedocs.io/en/>

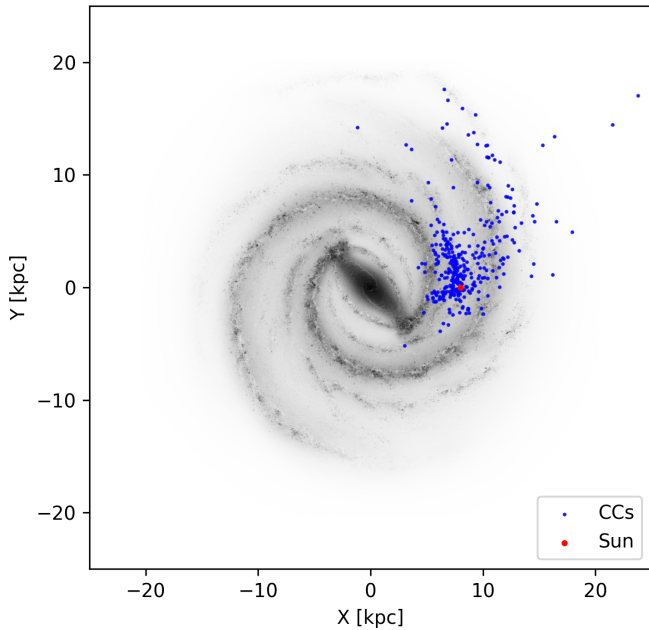


Fig. 1: Distribution of the sample of CCs across the thin disk in a Galactocentric reference frame. The face-on map of the Milky Way was made with the Python library `mw-plot`.

tra collected in the Stellar Population Astrophysics² (Origlia et al. 2019) program, while the other 5 were available in the TNG database and analyzed in Ripepi et al. (2021). The additional UVES and ESPaDOnS spectra have been analyzed in Trentin et al. (2023, 2024a), Martin et al. (2015), Andrievsky et al. (2016) and Kovtyukh et al. (2022). Details of the spectroscopic sample, such as the spectrographs, the resolution, and the spectral coverage, are provided in Table A.1. We estimated the individual distances by using the same approach adopted by dS23: Gaia EDR3 parallaxes for 20 CCs and W1-band PL for 2 CCs. For two CCs (ASAS 181024–2049.6 and ASAS SN J065046.50–085808.7) we estimated a different distance from Andrievsky et al. (2016) and Trentin et al. (2023). The new distance estimates are discussed in more detail in the Appendix B.

2.2. Spectral synthesis method

We utilized a Python version of Spectroscopy Made Easy (SME, Piskunov & Valenti 2017) for spectral synthesis, i.e. PySME³ (Wehrhahn 2021). PySME generates synthetic spectra based on a given set of atmospheric parameters, specified spectral intervals, and spectral resolution. It determines the optimal values for the selected atmospheric parameters by fitting the synthetic spectra to the observed data, taking into account the data uncertainties. The free parameters in the fitting process can include one or more stellar parameters, specific elemental abundances, and parameters related to atomic transitions in the line list. PySME provides two uncertainties according to two different methods. The first is based on SME statistics, i.e. uses a metric based on the distribution of the derivatives for each free parameter, esti-

imating the cumulative distribution function of the generalized normal distribution (Piskunov & Valenti 2017). The second is based solely on the least-squares fit. Further details on PySME are presented in Appendix C.

We adopted the line list presented in Table A.2 to estimate the atmospheric parameters, including Fe I, Fe II, Ti I and Ti II lines. We first tested this line list in a NARVAL solar spectrum with resolution $R \sim 68,000$ and in HARPS and NARVAL spectra of 16 Gaia benchmark stars, both dwarfs and giants (Blanco-Cuaresma et al. 2014, Casamiquela et al. 2025). Spectral lines of CCs are broader with respect to dwarfs and to red giants because of extended convective envelopes and pulsation, which are parameterized by higher microturbulence (v_{mic}) and macroturbulence (v_{mac}) values (da Silva et al. 2023; Luck & Lambert 1981, 1985; Bersier & Burki 1996). Consequently, spectral lines that show minimal blending in the Sun may experience more significant blending effects in CCs. Fig. C.2 shows the difference between the atmospheric parameters derived in this work and the reference values for the Sun and the benchmark stars provided by Blanco-Cuaresma et al. (2014).

The atmospheric parameters are estimated according to the following steps:

- Definition of the line masks, with information on the spectral coverage (segments), lines and continuum regions, and lines used for the RV adjustment;
- Definition of the continuum and RV optimization, we optimized the RV and the continuum level in each segment (`sme.vrad_flag = "each"`, `sme.cscale_flag = "linear"`, `sme.cscale_type = "match + mask"`);
- Definition of the atmosphere model, the abundance scale and the NLTE grids. We used the grid of MARCS model atmospheres (Gustafsson et al. 2008), that PySME-package defines as "marcs2012". We used solar abundances from Asplund et al. (2009) and the most updated versions of the NLTE grids⁴ (Amarsi et al. 2020). In order to have super-solar [N/Fe] and subsolar [C/Fe], typical of stars that have experienced mixing episodes during their evolution such as CCs, we used $A(\text{N})=8.23$ and $A(\text{C})=8.2$ for $[\text{Fe}/\text{H}]=0.0$;
- Spectral fitting in two different steps. In the first step, the free parameters are T_{eff} , $\log(g)$, $[\text{Fe}/\text{H}]$, v_{mac} and $A(\text{Ti})$, while in the second step the free parameters are $[\text{Fe}/\text{H}]$, v_{mic} , v_{mac} and $A(\text{Ti})$.

We focus on the elements for which NLTE grids are available in PySME, that are C, N, O, Na, Mg, Al, Si, S, K, Ca, Ti, Mn, Fe, Cu, Ba (Amarsi et al. 2020, 2022, 2025; Mallinson et al. 2024; Caliskan et al. 2025). The line list highlighted in Table A.3 is built from Gaia-ESO database according to the two quality parameters *gfflag* and *synflag* (further details in Appendix C). We used a NARVAL solar spectrum and three UVES spectra of Red Giant Branch (RGB) stars of M67, in order to validate our line list on giant stars for which the chemical abundances are well established (Carbajo-Hijarrubia et al. 2024). In PySME, we use the abundances of the specific element and the macroturbulence as free parameters. The final $[\text{X}/\text{Fe}]$ value is the median value of the $[\text{X}/\text{Fe}]$ values of each line. Table A.4 includes the $[\text{X}/\text{Fe}]$ values of the calibrating stars.

We estimated the sensitivity of each spectral line to the respective chemical abundance, computed as the mean abundance variation over 10 CCs spectra, in response to changes in the atmospheric parameters. We varied T_{eff} of ± 50 K and ± 100 K, $\log(g)$ of ± 0.1 and ± 0.2 dex, $[\text{Fe}/\text{H}]$ of ± 0.05 and ± 0.1 dex, and v_{mic} of ± 0.2 and ± 0.4 km s⁻¹. The results are listed in Table A.5

² The Stellar Population Astrophysics is an ongoing project based on a Large Program conducted at the Telescopio Nazionale Galileo (TNG) using the HARPS-N and GIANO B echelle spectrographs for about 74 nights from 2018 to 2021

³ `pysme-astro=0.6.20`, <https://github.com/MingjieJian/SME>

⁴ <https://zenodo.org/records/3888394>

3. Abundance Analysis

This section deals with the estimate of atmospheric parameters and the comparison with similar estimates available in the literature. Furthermore, we also discuss elemental abundances and their radial gradients with a comparison of the NLTE and LTE results. Finally, we present the kinematic properties of the sample.

3.1. Atmospheric parameters

We estimated the atmospheric parameters in LTE once the NLTE grids were turned off, leaving the rest of the code unchanged. Fig. 2 shows the difference of the atmospheric parameters and of iron and titanium abundances derived under NLTE and LTE assumption, with histograms. Effective temperature, surface gravity and microturbulence velocity are the most affected by the departures, while Fe and Ti display minimal variations, indeed the differences are, on average smaller than -0.03 dex. Although, the difference between NLTE and LTE abundances is modest and the standard deviation is, on average, smaller than 0.08 dex, the individual differences can be of the order of $\pm 0.4/0.5$ dex. Individual differences are in several cases larger than random errors, thus suggesting that the NLTE approach gives narrower metallicity distribution functions (smaller standard deviations).

Previous investigations addressing NLTE effects on the atmospheric parameters of FGK-type stars (see Ruchti et al. 2013; Kovalev et al. 2019) show that NLTE analyses generally yield higher T_{eff} , $\log g$ and $[\text{Fe}/\text{H}]$ values, particularly at low metallicities, as well as systematically higher microturbulent velocities of about $0.1/0.2 \text{ km s}^{-1}$ for $[\text{Fe}/\text{H}] \approx -1$. Cepheids are supergiants and the v_{mic} is intrinsically much higher. The increase is supported by data plotted in Fig. 2, where v_{mic} is the only parameter that shows a statistically significant offset ($-0.15 \pm 0.11 \text{ km s}^{-1}$), suggesting that microturbulence has a substantial impact on the determination of the other atmospheric parameters, which explains why we also find negative NLTE-LTE differences in T_{eff} , $\log g$ and $[\text{Fe}/\text{H}]$. Further discussion and a comparison with previous studies is presented in Appendix D.

To constrain possible differences with similar investigations we also compared our NLTE atmospheric parameters and abundances with those provided by dS23 by using the same spectra (Fig. 3). The LTE atmospheric parameters provided by dS23 were estimated by using pyMOOGi (Adamow 2017), a Python version of the MOOG code (Snedden 1973, 2002), atmosphere models from ATLAS (Castelli & Kurucz 2004), and the ARES code (Sousa et al. 2007, 2015) to measure the equivalent widths (EWs). In dS23, the authors used a large line list composed of 521 Fe I and Fe II lines based on different datasets and validated on calibrating Cepheids (Proxauf et al. 2018; da Silva et al. 2022). In contrast, our line list consists of 29 Fe I, Fe II, Ti I and Ti II, and was built in such a way as to be able to estimate atmospheric parameters via synthesis (see Sec 2.2 and Appendix C). Moreover, they derived T_{eff} by using the line depth ratio method (Kovtyukh 2007; Proxauf et al. 2018), and $\log(g)$, v_{mic} , and $[\text{Fe}/\text{H}]$ following the classical approach, i.e. consecutive iterations to reach ionization equilibrium for Fe I and Fe II lines and no trends in the Fe I abundances versus the EWs. We found positive trends in the difference of the atmospheric parameters, and indeed the slopes range from $+0.12 \pm 0.01$ for the effective temperature (panel a) to 0.45 ± 0.03 for the surface gravity (panel b) and to 0.43 ± 0.02 for the microturbulence (panel c), but the mean differences are well within 1σ . The difference in iron abundance also shows a positive slope, which can be attributed

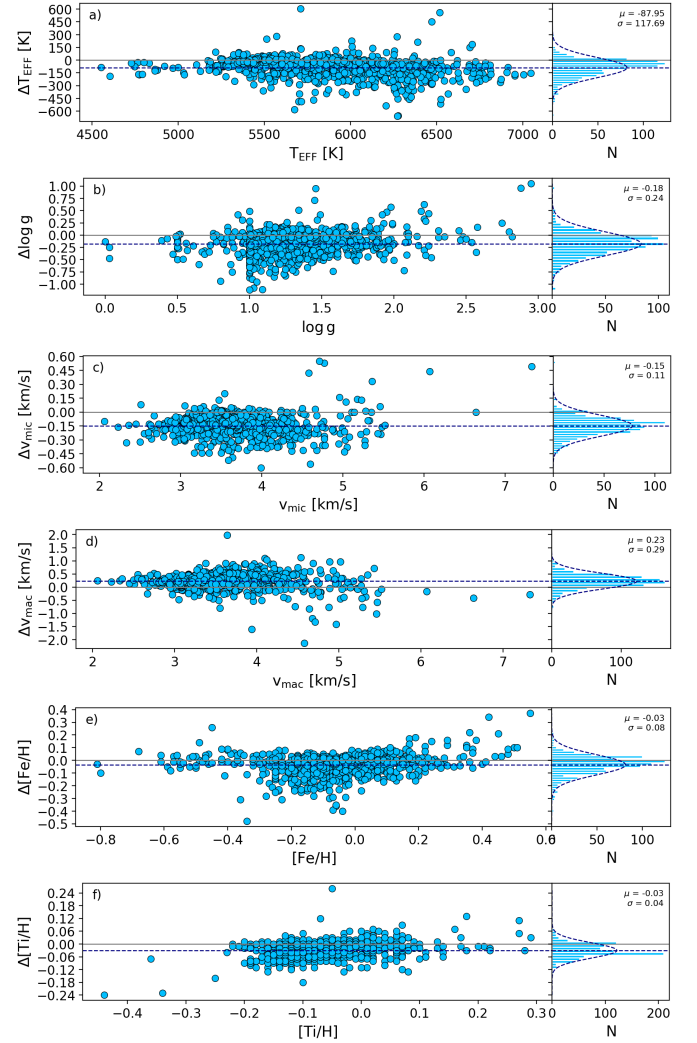


Fig. 2: Comparison of atmospheric parameters derived under NLTE and LTE assumptions. Each panel displays the difference (NLTE – LTE) as a function of the corresponding parameter on the left, and the distribution of the difference on the right. A blue-dotted Gaussian fit is overplotted on each histogram, the mean and standard deviation are labelled in the top-right corner. The blue-dashed horizontal line shows the mean difference, while the gray solid line the null difference.

both to differences in spectral analysis and to the fact that, as expected, the more metal-poor Cepheids appear even more metal-poor when analysed using a NLTE approach compared to LTE (Bergemann et al. 2012; Hansen et al. 2013). However, the mean difference is quite small (-0.07 ± 0.13).

Furthermore, we compared our iron abundance with those from previous studies, specifically we do have 101 Cepheids in common with Trentin et al. (2024a) and 209 Cepheids in common with Luck (2018). The top panel of Fig. 4 shows that the comparison with Trentin et al. (2024a) covers a broad range in iron abundance ($-0.85 \leq [\text{Fe}/\text{H}] \leq 0.40$ dex) and the difference does not show any trend, while the comparison with Luck (2018) (bottom panel of the same figure) covers a narrower range in iron abundance ($-0.50 \leq [\text{Fe}/\text{H}] \leq 0.50$ dex), but the difference displays a clear positive trend. The mean differences are quite small 0.02 ± 0.18 dex (Trentin et al. 2024a) and -0.08 ± 0.11 dex (Luck 2018), but the individual differences, once again, reach ± 0.4 dex.

The reader interested in a more detailed discussion concerning the atmospheric parameters is referred to appendix C.

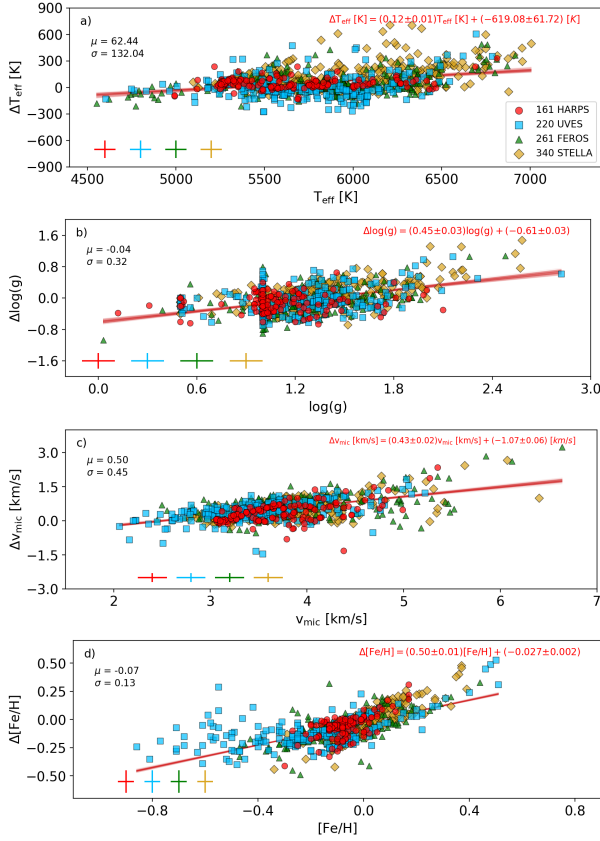


Fig. 3: Panel a) – Difference between the T_{eff} from current estimates and [dS23](#), plotted as a function of the current T_{eff} . Spectra from different datasets are marked with different colors. The error bars are plotted in the bottom left corner. The red line shows the linear fit and the coefficients are displayed on top of each panel. The mean and the standard deviation are also labeled. Panel b) – Same as Panel a), but for the surface gravity. Panel c) – Same as Panel a), but for the microturbulent velocity (km s^{-1}). Panel d) – Same as Panel a), but for the iron abundance (dex).

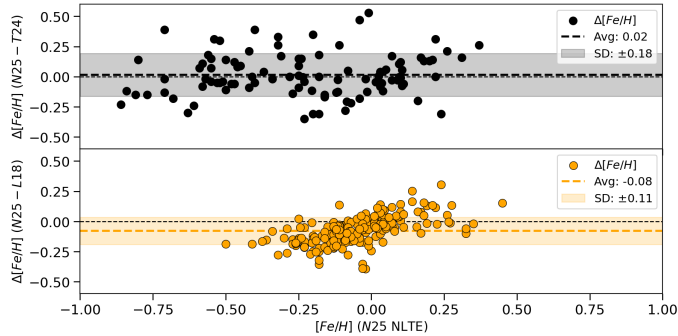


Fig. 4: Comparison of our metallicity with [Trentin et al. \(2024a\)](#) (top panel) and with [Luck \(2018\)](#) (bottom panel).

3.2. Iron radial gradient

The Galactocentric distances covered by our sample range from 4.6 kpc to 29.3 kpc, with four CCs located in the outskirts of the thin disk ($R_{\text{GC}} > 20$ kpc).

To investigate the iron radial gradient we adopted three different analytical fitting functions (Fig. 5) by taking into account errors on both axes: $[\text{Fe}/\text{H}] \propto R_{\text{GC}}$, $[\text{Fe}/\text{H}] \propto \log R_{\text{GC}}$, and a bilinear function. Moreover, we adopted a Bayesian inference

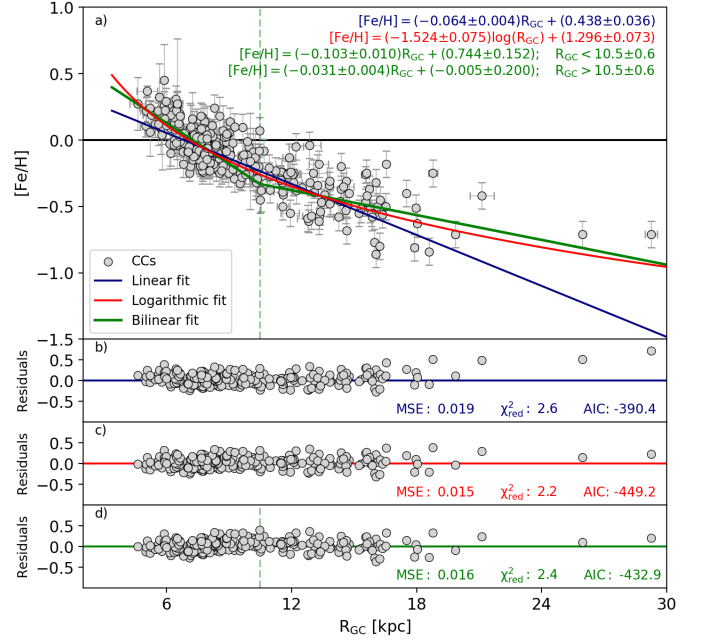


Fig. 5: Panel a) – Iron radial gradient for classical Cepheids as a function of the Galactocentric distance. The linear, the logarithmic and the bilinear fits are shown respectively in blue, red and green. The vertical green dashed line shows the “knee” of the bilinear fit. The coefficients of the fits are labelled. Panel b) – residuals of the linear fit. The values of the mean square error, the reduced chi-squared and the AIC score are labelled. Panel c) – Same as panel b), but for the logarithmic fit. Panel d) – Same as panel b), but for the bilinear fit.

framework to estimate model parameters and their uncertainties. Specifically, given data D and parameters θ , the posterior probability is:

$$p(\theta|D) \propto \mathcal{L}(D|\theta)\pi(\theta),$$

where $\pi(\theta)$ is the prior and $\mathcal{L}$ the likelihood. We work with the log-posterior $\log p(\theta|D) = \log \mathcal{L}(D|\theta) + \log \pi(\theta) + \text{const}$. We sampled the posterior with the affine-invariant Markov Chain Monte Carlo (MCMC) sampler implemented in the `emcee` library to obtain parameter point estimates (MAP) and credible intervals (the 68% credible region from the marginal posterior). In particular, we adopted mixture models ([Foreman-Mackey 2014](#)), for which the per-data-point likelihood is given by a weighted sum of a foreground (signal) and a background (outlier) component. Writing Q for the mixing fraction (the prior probability that a datum belongs to the background component), the contribution of datum x_i is:

$$\mathcal{L}_i = (1 - Q)\mathcal{L}_{fg}(x_i, y_i, \sigma_{x,i}, \sigma_{y,i}|\theta_{fg}) + Q\mathcal{L}_{bg}(x_i, y_i, \sigma_{x,i}, \sigma_{y,i}|\theta_{bg}),$$

and the total log-likelihood is $\log \mathcal{L} = \sum_i \log \mathcal{L}_i$. The parameter Q therefore controls the relative weight of the background/outlier model and robustifies the inference against points that are inconsistent with the foreground model. In Gaussian implementations the components are typically normal distributions with their own means and variances (optionally including an extra “jitter” term added in quadrature to account for underestimated measurement errors). Numerical evaluation of $\log \mathcal{L}$ should use stable log-sum-exp arithmetic to avoid underflow when one component is much smaller than the other. Convergence of the MCMC chains was verified via autocorrelation-time estimates and visual inspection of trace plots.

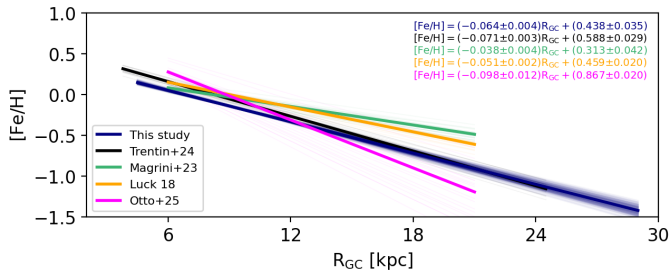


Fig. 6: Comparison between this study, Trentin et al. (2024a), Magrini et al. (2023), Luck (2018) and Otto et al. (2025) in the $[\text{Fe}/\text{H}]$ vs R_{GC} plane.

We provide the mean square error, the reduced chi-squared and the Akaike's information criterion (AIC) scores for each fit. The AIC is a method for evaluating and comparing statistical models, which provides a measure of the quality of a statistical model's estimate, taking into account both the goodness of the fit and the complexity of the model. It is defined as $\text{AIC} = 2k - 2 \log L$, where k is the number of parameters of the model and L is the maximized value of the total likelihood. Models with lower AIC should be preferred.

Data plotted in the top panel of Fig. 5 display quite clearly that the bilinear and the logarithmic fit reproduce quite well the radial gradient of CCs when moving from the innermost to the outermost disk regions. Note that the agreement for the bilinear fit is expected, since this fit has more degrees of freedom when compared with the linear and the logarithmic fit. However, the bilinear fit is also prone to possible systematics, since the edge of the two different intervals in Galactocentric distances is arbitrary. The maximized likelihood is obtained at a value of the "knee" of 10.5 kpc (green dashed vertical line). In any case, the logarithmic fit is the most accurate analytical representation of the radial gradient, since the statistical parameters (MSE, reduced chi-squared, AIC score) attain their smallest values when compared with the other fits.

Fig. 6 shows the comparison between our linear iron radial gradient with those available in the literature, revealing a slope of -0.064 dex/kpc, which is steeper than the slopes provided by Magrini et al. (2023, -0.038) by using open clusters and by Luck (2018, -0.051) by using CCs, and shallower than the slopes provided by Trentin et al. (2024a, -0.071) based on CCs and by Otto et al. (2025, -0.098) by using open clusters.

Fig. 7 displays the comparison of the current iron radial gradient with that of *ds23* in a log-log plane. The slope of the current gradient is steeper (-1.53 vs. -0.91). The two radial profiles nearly overlap for $R_{\text{GC}} \lesssim 10$ kpc, but the current one becomes steeper in the outer disk. This difference is due to several reasons: the different spectral analysis method (including a different line list), the increased number of CCs at $R_{\text{GC}} \approx 15 - 20$ kpc, which were not included in the *ds23* sample, and NLTE effects, since the outer disk Cepheids are the most metal-poor stars in the current sample.

To overcome possible problems in the choice of the adopted analytical function to describe the iron radial gradient, we decided to use the Gaussian Process Regression (GPR) to model the data without assuming a predefined functional form. The GPR is a Bayesian non-parametric method for modeling an unknown function $f(x)$ in a regression setting, starting from a Gaussian process prior:

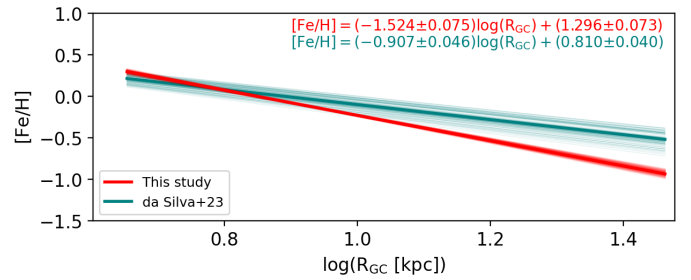


Fig. 7: Comparison between the current iron radial gradient and that of *ds23* in a log-log plane.

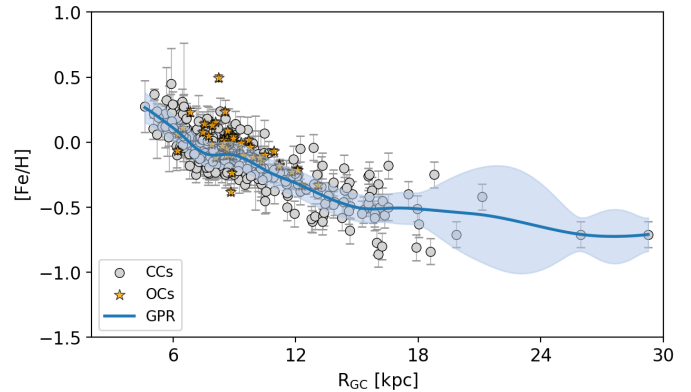


Fig. 8: In gray iron abundances of CCs as a function of R_{GC} . The orange symbols display open clusters younger than 400 Myr from Otto et al. 2025. GPR applied to the CCs distribution is highlighted in blue, while the shaded region highlights the uncertainties on the model.

$$f(x) \sim GP(m(x), k(x, x'))$$

defined by a mean function $m(x)$ and a covariance (kernel) function $k(x, x')$. Observational data $y = f(x) + \epsilon$ (with noise $\epsilon \sim N(0, \sigma^2)$) induce a predictive posterior distribution for f at new input points. This posterior yields both a predictive mean (the best estimate of the function) and a predictive variance (quantifying uncertainty in the estimate). The kernel function encodes assumptions about smoothness, length scales and correlation structure of the function. Hyper-parameters of the kernel are typically inferred by maximizing the marginal likelihood (or by Bayesian integration). Therefore, GPR adaptively fits complex, non-linear relationships while providing principled uncertainty quantification. In short, GPR offers a flexible and rigorous framework for regression when uncertainty estimates are required and when one prefers to avoid strong parametric assumptions on the functional form.

We employed GPR by using Gpy⁵. After testing different kernel functions, we adopted GPy.kern.Matern32 according to the distribution dispersion and different sampling across the x-axis. Furthermore, we optimized the variance and the length-scale of the kernel function.

Data plotted in Fig. 8 display several interesting features worth being discussed in more detail. The GPR fits quite well the distribution of CCs as a function of the Galactocentric distance and this outcome applies to both the inner and the outer disk Cepheids. The GPR shows a particular trend in the data across

⁵ <https://gpy.readthedocs.io/en/deploy/>

the Solar circle, a "shoulder" at $R_{GC} \sim 8$ kpc. Moreover, GPR also shows a steady change in the slope for $R_{GC} \sim 14 - 16$ kpc. This evidence indicates that the linear fit fails to take account for the radial gradient variations of the CCs in the outer disk, thus suggesting that the change is an intrinsic feature when moving into the outer disk. The radial distribution of open clusters younger than 400 Myr collected by [Otto et al. \(2025\)](#) agrees quite well with the radial distribution of Cepheids. This agreement is expected, since CCs are stellar tracers younger than ~ 300 Myr. Unfortunately, the range in Galactocentric distances covered by this sample is too small when compared with CCs, but the agreement is very promising.

3.3. Radial gradients of chemical abundances

Distributions of the chemical abundance ratios are shown in Fig. 9, together with overlaid Gaussian fits (dashed red lines). The present results indicate Solar abundance ratios, within 1σ , for four α -elements (O, Mg, Si, S) and for one light metal (Al). In contrast, [Na/Fe] ratio shows a clear overabundance of $+0.30 \pm 0.12$ dex, while [Cu/Fe] appears underabundant, with a mean value of [Cu/Fe] = -0.23 ± 0.20 dex. The Na overabundance is expected, since Cepheids are intermediate-mass stars and the low-mass tail experiences the first dredge-up along the red giant branch (see Fig. C13 in [Bono et al. 2024](#)). A similar Na overabundance was also reported by [Trentin et al. \(2024a\)](#), who found a mean value of [Na/Fe] = $+0.39 \pm 0.16$ dex, whereas their [Cu/Fe] measurements yielded a mean value of [Cu/Fe] = $+0.10 \pm 0.27$. [Otto et al. \(2025\)](#), analyzing open clusters who derived abundances for both main-sequence and giant stars, obtained mean values of [Na/Fe] = -0.16 ± 0.72 dex and [Cu/Fe] = -0.04 ± 0.45 dex. In passing, an overabundance in [Na/Fe] was also found by [Genovali et al. \(2015\)](#).

Fig. 10 shows the radial abundance gradient of ten out of the eleven investigated chemical elements as a function of R_{GC} , with blue solid lines representing the GPR models. Similar to iron, a clear flattening of the gradients at $R_{GC} \approx 14 - 16$ kpc is obtained for all the elements, but for oxygen. Unfortunately, O abundances are only limited to CCs located within 14 kpc, since our Oxygen-determination is based on the triplet at $7770\text{\AA}$ which is only covered by FEROS, STELLA and ESPaDOnS spectra in our spectroscopic dataset.

Flattening of the gradients is also supported by data plotted in Fig. 11, where the linear, logarithmic and bilinear fits are shown. The AIC scores support the logarithmic fit for Al, Si, S, Ca, Ti, Mn and Cu among the three models. Moreover, the linear fit is always the least suitable model of the three, while Na and Mg favor the bilinear model with respect to the logarithmic one, but with a different knee position (9.8 vs 15.3 kpc).

To further investigate the radial variation of the abundance gradients, Fig. 12 presents the different [X/Fe] ratios as a function of Galactocentric distance. The radial trends remain approximately constant for most of the investigated elements across the entire disk. The abundance ratio slopes are vanishing, suggesting that these elements follow the same radial behavior as iron. Three elements, however, deviate from this pattern: Al, Mn, and Cu. Mn and Cu display well-defined negative gradients (-0.018 ± 0.003 , -0.009 ± 0.005 dex kpc^{-1} , respectively), while Al shows a positive gradient ($+0.012 \pm 0.003$ dex kpc^{-1}). The negative trend observed for Mn supports the behavior found for the iron-peak elements (Mn, Co, Ni) in OCs by [Otto et al. \(2025\)](#), whereas the nearly flat [X/Fe] ratios for O, Mg, Si, S, and Ca (all with slopes < 0.01 dex kpc^{-1}) further confirm the overall homo-

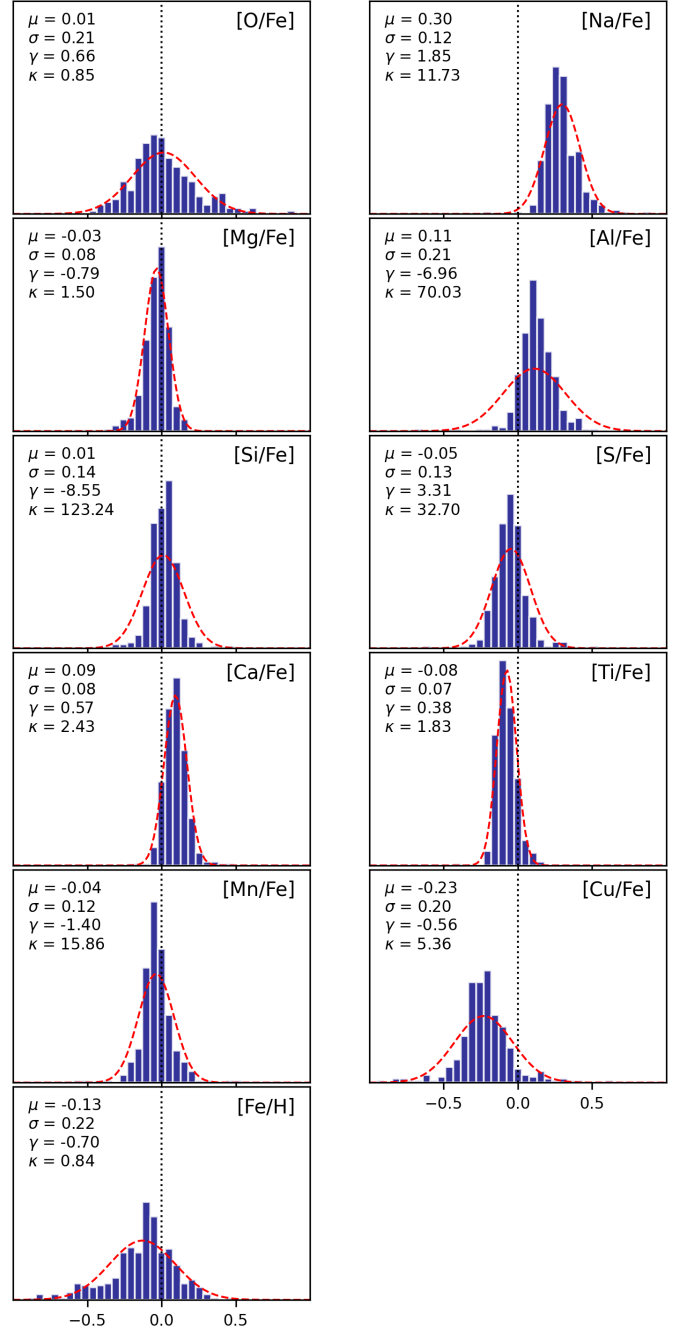


Fig. 9: Distribution of chemical abundance ratios ([X/Fe]). The red dashed line shows the Gaussian fit. Mean value (μ), standard deviation (σ), skewness (γ) and kurtosis (κ) are labeled on the top left of each panel. The bottom left panel shows the iron distribution function.

geneity of α -element enrichment across the thin disk. [Genovali et al. \(2015\)](#) also reported nearly constant trends for Na, Al and Si. In contrast, they found mild evidence of positive slopes for Mg (0.015 ± 0.006 dex kpc^{-1}) and Ca (0.028 ± 0.004 dex kpc^{-1}). However, their results were based on a smaller Cepheid sample, characterized by higher dispersion and a more limited spatial coverage.

Fig. 13 shows for the investigated elements the abundance ratios as a function of $\log(P)$ along with a linear fit. CCs obey to a well defined Period-Age relation ([Bono et al. 2005](#), and references therein). The individual age steadily increases when

moving from long- to short-period Cepheids. The abundance ratio of all the elements remain approximately constant across the full age range of the Cepheids in the sample. Only O, Na, Si, Ti, and Cu show a weak positive trend, 0.11 ± 0.20 , 0.10 ± 0.03 , 0.07 ± 0.02 , 0.09 ± 0.02 , 0.08 ± 0.05 , respectively. For oxygen, the higher dispersion and the shorter age coverage may affect the result. In other words, age-related effects are weak or negligible in the radial gradients and abundance ratios presented above. [Genovali et al. \(2015\)](#) found a slight negative trend for Ca (-0.15), in contrast with the null slope obtained in our study. In [dS23](#), they found a slope of 0.08 ± 0.02 in the $[S/Fe]$ vs $\log P$ plane, while we obtained a weaker positive trend of 0.04 ± 0.02 . This suggests that the overall consistency of the abundances in our sample, which reduces the dispersion, allows for a more accurate assessment of whether a trend is present.

In Fig. 14 the $[X/Fe]$ vs $[Fe/H]$ planes are shown. We emphasize the positive trend observed for Mn. Manganese is mainly produced in thermonuclear supernovae (Th-SNe), with yields that increase with metallicity ([Badenes et al. 2008](#); [Kobayashi & Nomoto 2009](#)), which explains the positive trend. Sodium, copper and aluminum are predominantly produced in core-collapse supernovae (CC-SNe); however, we observe a slight positive trend for Cu and a negative trend for Na and Al, reflecting their different metallicity dependences. [Kobayashi et al. \(2020\)](#) predict a constant or slightly negative trend for Na and Al, and a slight positive trend for Cu over the metallicity range we cover ($-1 < [Fe/H] < 0.5$ dex), in agreement with our findings. Moreover, for Na and Al, the slight positive trend in $[Na/Fe]$ and $[Al/Fe]$ vs R_{GC} corresponds to a negative trend in the $[Na/Fe]$ and $[Al/Fe]$ vs $[Fe/H]$ plane. Similarly, the negative trends of $[Cu/Fe]$ and $[Mn/Fe]$ with Galactocentric distance correspond to positive trends in the $[Cu/Fe]$ and $[Mn/Fe]$ vs $[Fe/H]$ planes. For oxygen, magnesium, silicon, sulfur, calcium, and titanium, we expect a roughly constant trend in the metal-poor region of the sample and a decreasing trend at higher metallicities. This behavior arises because these elements are predominantly produced in CC-SNe, whereas metal-rich stars are increasingly affected by Th-SN enrichment, which reduces their $[X/Fe]$ ratios. Since Si, S, Ca and Ti also receive contributions from Th-SNe, the negative trend in the metal-rich region is weak. In contrast, O and Mg are more affected and show a higher decreasing trend.

3.4. Radial gradients: LTE vs NLTE

We measured the chemical abundances under the LTE assumption and performed the logarithmic fit of the LTE abundance gradients. Fig. 15 reveals that for almost all the elements the NLTE slopes of the logarithmic fit are shallower. Only S and Cu show a similar slope, which suggests negligible NLTE effects for the lines employed for these two elements. This is expected for S, since we used an S I triplet that is weakly prone to NLTE effects ([Duffau et al. 2017](#)). NLTE have a minor impact on Cu I lines in stars with $[Fe/H] \geq -1$ dex. [Shi et al. \(2018\)](#) found a mean abundance increase of about $+0.07$ dex when applying NLTE corrections, a result previously suggested by [Yan et al. \(2016\)](#).

The histograms in Fig. 16 illustrate the comparison of the NLTE and LTE abundances where, on average, lower NLTE abundances for O (-0.74 ± 0.21 dex), Na (-0.11 ± 0.08 dex), Al (-0.07 ± 0.10 dex), Si (-0.08 ± 0.15 dex) and Cu (-0.13 ± 0.17 dex), and more centered distributions for Mg (-0.03 ± 0.08 dex), S (-0.02 ± 0.10 dex), Ca (0.06 ± 0.09 dex) and Mn (0.02 ± 0.14 dex). Our results indicate that the oxygen lines used in this study (OI 777 nm triplet) are strongly affected by NLTE, consistent with [Amarsi et al. \(2016\)](#), who reported corrections of about 0.6 dex

or higher at solar metallicities. For the other elements, the NLTE corrections are more moderate and agree with the findings of [Amarsi et al. \(2020\)](#). In appendix D we further discuss the NLTE abundances for Mn and O.

In Fig. 17, we compare the slopes of our chemical gradients with those reported in previous studies. A general agreement is observed in the inner disk for most elements, although a slight underabundance is found for sulfur and copper when comparing the GPR profiles to the linear trends from other works. Moreover, the GPR profiles provide additional information on the behavior of Na, Mg, Al, S, Ca, Ti, Mn and Cu, revealing a shoulder at the solar circle similar to that obtained in the iron GPR profile. The GPR profiles also suggest a change of slope around 14-16 kpc, resulting in a flatter gradient in the outer disk compared to the linear trends of previous studies, particularly for Na, Mg, Al, Si, S, Ca, Ti, Mn and Cu.

In Fig. 18, we emphasize the agreement with the [dS23](#) logarithmic profiles for oxygen and sulfur. According to [dS23](#), the radial gradient of sulfur is steeper than that of iron, and sulfur is on average under-abundant compared to oxygen. This discrepancy points to the need for a revision of chemical evolution models, particularly regarding sulfur yields from massive stars, in order to reproduce observed trends in the outer Galaxy. However, our results confirm the general behavior of sulfur but show a profile that is more consistent with iron and the other elements, rather than reproducing the steeper trend reported by [dS23](#).

3.5. Kinematic properties

The kinematic properties of CCs provide fundamental insights to constrain their true nature of thin disk stellar population. Their orbits were integrated by using the *galpy* code, a Python library dedicated to Galactic dynamics ([Bovy 2015](#)). Positional parameters (RA, DEC, distance) and proper motions are mainly based on Gaia DR3. Individual distances have been discussed in more detail in § 2.1, while radial velocities are mainly based on current spectroscopic measurements. The galactic potential employed is MWPotential2014.

To improve the sampling across the four quadrants, the current sample was complemented with CCs analysed by [Luck \(2018\)](#) and by [Trentin et al. \(2024a\)](#) for which homogeneous elemental abundances are available. We did not include the 16 CCs located at 3–5.6 kpc and analyzed by using high resolution NIR spectra by [Matsunaga et al. \(2023\)](#), because of possible systematics between NIR and optical spectroscopy. Moreover, the quoted samples have a sizable number of objects in common with the current sample and they have been adopted to move their chemical abundances into our metallicity scale.

Fig. 19 shows nine different kinematic planes with the orbital properties of CCs compared with a sample of Type II Cepheids (TIICs). These planes are based on kinematic diagnostics that are quite useful to characterize Galactic stellar populations, and in particular to identify stars that are associated to the different Galactic components ([Lane et al. 2022](#); [Bonifacio et al. 2024](#); Bono et al. submitted). According to a Galactocentric cylindrical frame, V_R , V_T and V_z are the radial, the tangential and the vertical velocity, while L_z and E are the angular momentum component along the Galactic Z-axis and the total orbital energy. J_R , J_ϕ , and J_z are the radial, the azimuthal, and the vertical actions, while J_{tot} is the total action, defined as $J_{tot} = |J_\phi| + J_R + J_z$. λ_z is the normalized angular momentum (circularity), defined as $\lambda_z = J_z/J_{max}(E)$ where $J_{max}(E)$ is the angular momentum of a circular orbit with the same maximum binding energy. Finally, Z_{max} denotes the maximum height along the orbit above the Galactic

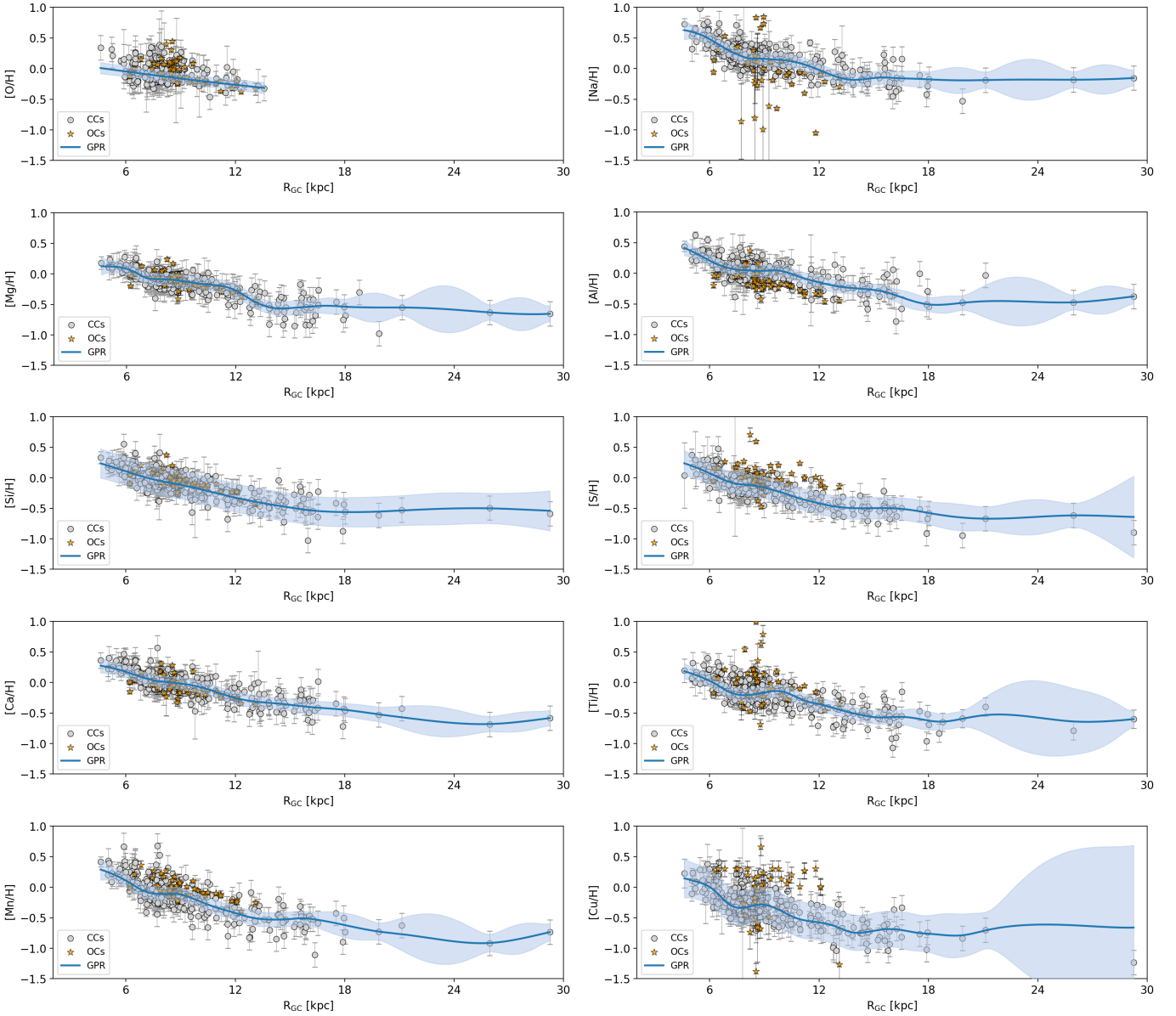


Fig. 10: Chemical gradients with GPR modeling for O, Na, Mg, Al, Si, S, Ca, Ti, Mn and Cu. Colors and symbols are the same of Fig. 8.

plane. The bulk of the kinematic sample has cold orbits typical of thin disk stars, as stated by the coherent distribution in the 9 planes. In particular, in the Toomre diagram (panel a) our sample is grouped in the bottom-right corner, characterized by $V_T \simeq 250 \text{ km s}^{-1}$ and low perpendicular velocities, typical of the thin disk component. On the other hand, TIICs cover a broad range in transversal velocity and a fraction of them also attain negative transversal velocities suggesting that they are on retrograde orbits. The stark difference between the two samples is fully supported by the Lindblad plane (panel c) in which CCs are distributed along the tiny region typical of thin disk stars, while TIICs move from the region typical of hot stellar components with radial orbits and random motions. The same outcome applies to the action-diamond diagram (panel f), since CCs are grouped in the right corner, since the major contribution to the action is given by their L_z . Furthermore, CCs show high circularities ($\lambda_z \gtrsim 0.8$, which means almost circular orbits and strong rotation), low eccentricities and low Z_{max} values (panels g, h, i). The TIICs display a more complex behavior, since they move

from a warm stellar component more typical of thick disk stars to hot stellar components and to retrograde orbits. A more quantitative analysis will be provided in a forthcoming paper (Nunnari et al. in preparation).

The left panel of Fig. 20 shows the radial distribution of CCs projected onto the Galactic plane. The symbols are color-coded according to the azimuth angle Θ . We note that stars with $\Theta \lesssim 180^\circ$ typically reach smaller Galactocentric radii than those at $\Theta \gtrsim 180^\circ$. This trend may arise from the intrinsic spatial distribution of stars in the sample (for example, due to the presence of spiral arms), from the selection function of our dataset, or from a combination of both effects. Data plotted in the central and right panel of the same figure display that for Galactocentric distances larger than 11-12 kpc the height above the Galactic plane steadily decreases. This finding is expected, since the variation is mainly caused by the Galactic warp. Specifically, it is known from previous observations that the Galactic midplane is bent downwards (i.e. $Z < 0$) in the portion of the disk covered by our dataset. The reader interested in a more detailed discussion

concerning the Galactic warp is referred to [Chen et al. \(2019\)](#), [Skowron et al. \(2019\)](#), [Lemasle et al. \(2022\)](#) and [Poggio et al. \(2025\)](#).

It is worth mentioning that the change in the slope of the radial metallicity gradient occurs at Galactocentric radii where the Galactic warp begins to become significant (i.e. approximately 11-12 kpc), tentatively suggesting that the two aspects might be related. It is important to note, however, that the spiral arms and bar-driven radial mixing can also potentially generate variations in the radial metallicity gradient (see e.g., [Buder et al. 2025](#)).

The orbital properties (radial velocity and circularity) of these objects fully support their association to the thin disk. However, more photometric and spectroscopic data are required to address this issue on a more quantitative basis.

4. Summary and conclusions

The present study represents one of the most comprehensive NLTE spectral analyses performed on CCs spanning an unprecedented range in Galactocentric distance (4.6–29.3 kpc) and elemental coverage including light, α , odd-Z, and iron peak elements. Such an approach is crucial because CCs, as young (ages younger than a few hundred Myrs), luminous super-giant stars, sample the current chemical enrichment of the Galactic thin disk with minimal contamination from older populations. We used the largest high-resolution optical spectral sample, for which we estimated the atmospheric parameters and chemical abundances. We constrained the iron and the chemical gradients for O, Na, Mg, Al, Si, S, Ca, Ti, Mn and Cu.

Previous studies available in the literature have modeled the chemical gradients of the thin disk, by assuming either a linear ([Luck 2018](#); [Genovali et al. 2014](#); [Trentin et al. 2024a](#)) or a log-arithmetic ([da Silva et al. 2023](#)) function. The fits based on a linear function overestimate the observed gradient in the inner disk ($R_{GC} \lesssim 7$ kpc) and underestimate the observed gradient in the outer disk ($R_{GC} \gtrsim 18$ kpc). The difference with previous investigations, based on a linear fit, is mainly due to the larger range in Galactocentric distances covered by the current sample. This means that we are dealing with 22 CCs with R_{GC} smaller than 6 kpc and 33 CCs with a R_{GC} larger than 15 kpc⁶.

A flattening of the abundance gradient in the outer disk has been reported by several studies (e.g., [Luck 2018](#); [Kovtyukh et al. 2022](#); [da Silva et al. 2023](#)). A similar behavior has also been observed in open clusters ([Magrini et al. 2023](#); [Otto et al. 2025](#)), where a bilinear fitting approach has been adopted. Nevertheless, the AIC test performed by [Otto et al. \(2025\)](#) indicated that introducing two additional parameters in a bilinear model does not significantly improve the fit compared to a simple linear relation. In any case, such modeling inherently implies a break at a certain Galactocentric radius, raising a theoretical question about the physical origin of this discontinuity. For our sample, the best-fitting bilinear model applied to the iron gradient (Fig. 5) suggests a break, or "knee", at $R_{GC} = 10.5 \pm 0.4$ kpc. This model provides a better fit than the linear one, as indicated by smaller residuals and a lower AIC score. However, when the same bilinear approach is applied to the other elements, the inferred knee positions vary considerably, ranging from 8 to 15 kpc. The bilinear model yields the lowest AIC score only for Mg compared to alternative models (Fig. 11), although in this case the knee oc-

curs at $R_{GC} = 15.3 \pm 0.7$ kpc, a value inconsistent with the other elements and with previous studies.

In [dS23](#), a large sample of 356 CCs, including four at $R_{GC} \gtrsim 20$ kpc, was analyzed using a log-log relation for Fe, O, and S. In Fig. 5, we applied the same functional form to fit the iron gradient. As expected, this approach yields lower mean squared error, chi-squared, and AIC values compared to the linear and bilinear models, except in the case of Mg, as noted above. Furthermore, as shown in Fig. 7, the slope we obtain differs from that reported in [dS23](#), likely due to differences in spectral analysis, such as the line list and the use of NLTE grids for estimating atmospheric parameters, and to the increased number of Cepheids located in the outskirts of the disk. Overall, the log-log fit appears to be the most reasonable among the three analytical models, although sodium and magnesium are better modeled by the bilinear fit (Fig. 11), suggesting that a single analytical function may not fully describe all gradients. For this reason, we also employed Gaussian Process Regression, which does not require any predetermined functional form. This modeling confirms that the radial abundance gradients in the thin disk generally flatten with increasing Galactocentric distance across all studied elements.

The flattening of the gradients is a feature that extends to extragalactic studies. Works on other disk galaxies have shown that, although the inner regions show a negative gradient, the radial distribution of metallicity flattens to a virtually constant value beyond the isophote radius. This behavior has been found in the disk galaxy M83 ([Bresolin et al. 2009, 2016](#)), and subsequently confirmed in NGC 1512 and NGC 3621 ([Bresolin et al. 2012](#); [Kudritzki et al. 2014](#)), where the oxygen abundance remains homogeneous and approaches an almost constant value in the outer regions. These observations suggest that flattening is a common feature of the metallicity gradient in spiral galaxy disks.

The current metallicity of CCs can be compared with that of B-stars located in the solar neighborhood. According to [Nieva & Przybilla \(2012\)](#), early B-type stars within 500 pc show a mean iron abundance of 7.52 ± 0.03 dex, consistent with the solar value. This result is based on a NLTE spectral analysis of 29 early B-type stars. In the present work, the mean iron abundance of the 11 CCs located within 500 pc of the Sun is 7.49 ± 0.13 dex, in perfect agreement with [Nieva & Przybilla \(2012\)](#). Differently, the 93 CCs located at the solar circle (i.e. all Cepheids with $7.7 \lesssim R_{GC} \lesssim 8.7$ kpc), exhibit a slightly iron underabundance of 7.42 ± 0.10 dex, but still consistent within 1σ .

Our results for chemical abundance ratios as a function of R_{GC} reveal that $[X/Fe]$ ratios remain approximately constant across the entire thin disk. The near-flat distributions, with average values within 0.10 dex and minimal slopes, indicate that most elemental abundances scale closely with iron throughout the Galactic disk. This consistency suggests that these elements trace the gas distribution in the thin disk in a coherent manner. Slight overabundances in Al and Na, underabundance in Cu, and the small but measurable slopes for Mn, and Cu point to subtle variations, yet they do not compromise the overall trend of roughly constant abundance ratios. The negative slopes observed for Mn and Cu are in agreement with previous findings for iron-peak elements ([Otto et al. 2025](#)).

Chemical gradients from previous studies, such as [Genovali et al. \(2014\)](#) and [dS23](#), indicate that the outer disk exhibits a higher dispersion. This effect can be partially attributed to the Galactic warp ([Kerr et al. 1957](#); [Oort et al. 1958](#); [Lemasle et al. 2022](#); [Poggio et al. 2025](#)), a large-scale distortion of the disk that affects the estimated Galactocentric distances. In regions where the warp is pronounced, Cepheids appear closer in Galactocen-

⁶ The distance of the Cepheid ASAS J062939-1840.5 was estimated by using a mid-infrared (W1-band) PL relation and we found $R_{GC} = 29.3 \pm 0.3$ kpc. The distance of this object was also estimated by [Trentin et al. \(2023\)](#) using period-Wesenheit relations and they found $R_{GC} \approx 25$ kpc.

tric distance than they would in a non-warped disk. However, [Lemasle et al. \(2022\)](#) noted that the dispersion introduced by the warp is too small to fully account for the larger spread around the mean metallicity gradient reported by [Genovali et al. \(2014\)](#) and [da Silva et al. \(2023\)](#). Instead, this increased dispersion likely reflects larger uncertainties associated with distant stars, whose spectra typically have lower signal-to-noise ratio. In our study, the gradients do not show a significant increase in dispersion (Fig. 10), in agreement with [Lemasle et al. \(2022\)](#). It is worth noting that our sample contains fewer Cepheids in the outer disk ($R_{GC} > 12$ kpc), where the warp is stronger, than in the inner regions. Nevertheless, the trend visible in Fig. 20 appears to be a direct consequence of the Galactic warp. Similar trends have also been observed using larger kinematic samples of Cepheids ([Lemasle et al. 2022](#), Fig. A.3; [Poggio et al. 2025](#)).

Results for the sample of the youngest open clusters (OCs) from [Otto et al. \(2025\)](#), age < 400 Myrs) overlaps with that of Cepheids. Sodium is an exception, since Cepheids are on average overabundant in this element. It was suggested by [Takeda et al. \(2013\)](#) and confirmed by [da Silva et al. \(2015\)](#) in a comparison between giant and dwarf field stars, that this overabundance is mainly due to stellar evolution effects: the first dredge-up brings some sodium from the interior to the atmosphere, altering its initial chemical composition ([Ventura et al. 2013](#); [Lagarde et al. 2012](#)). In contrast, the chemical abundances of OCs have been measured in both giant and dwarf stars, with the latter still unaffected by significant mixing processes. Comparisons with other stellar populations, such as open clusters, generally reveal similar radial trends for iron-peak and alpha elements, further supporting the homogeneity of chemical enrichment in the thin disk, at least in the region between 5 and 14 kpc. These findings are consistent with previous studies ([Magrini et al. 2023](#); [Yong et al. 2012](#); [Otto et al. 2025](#)).

This study opens multiple possibilities for future investigation. In addition to the natural need of increasing the sample size, especially in the outer regions of the disk, we emphasize the following points:

- The chemical abundances of other elements, such as heavy elements produced by the s- and r-processes, have yet to be determined. As discussed in [Trentin et al. \(2024a\)](#), some elements—such as Scandium, Zinc, and Zirconium—show a bifurcation at the low-metallicity tail. This anomalous behavior may arise from variations in spectral line analyses or from the presence of distinct stellar populations associated with different spiral arms. Since our sample includes Cepheids located in various spiral arms, a future study focusing on these elements could help clarify this issue. However, NLTE grids for these elements are not yet available in PySME.
- In the inner disk, for $R_{GC} < 5$ kpc, a double sequence can appear, one that follows the logarithmic profile and the other showing a change in the slope. Note that a significant decrease in the metallicity of the young stellar population, confined to the central/inner regions, has been found in barred spiral galaxies such as NGC 1365 and M83 ([Sextl et al. 2024, 2025](#)). Furthermore, [Andrievsky et al. \(2016\)](#) discussed a change in the slope at $R_{GC} < 5$ kpc, suggesting that it might be caused by a reduced star formation rate compared to the outer disk regions. To account for such a change of slope in the radial gradients, it has also been suggested that the inner disk regions have also been affected by dilution effects caused by the infall of more metal-poor gas ([Genovali et al. 2015](#), and references therein). Further extensive sampling at small R_{GC} is crucial to investigate this feature of the gradi-

ent and its implications for the chemical enrichment of the Galactic bar.

- The $[X/H]$ gradients exhibit similar slopes among the studied elements, and the $[X/Fe]$ ratios remain approximately constant across the entire range of R_{GC} explored. This behavior may indicate that these elements could serve as potential tracers of the underlying gas density profile. However, further validation is required to confirm these trends, and it remains to be established whether such a result can be extended to the disks of external galaxies.
- The line list employed in this work serves as a tool for estimating atmospheric parameters and has been specifically adapted to optical spectra of Galactic Classical Cepheids. However, further investigation is required to extend the applicability of this method to the near-infrared (NIR) domain to fully exploit the NIR spectroscopic data and reach the more extinct regions. In addition, the metallicity range over which the selected lines have been tested should be broadened toward lower values, as the current analysis only covers the metallicity interval typical of the Galactic thin disk (-0.8 to $+0.4$ dex). This extension is essential to enable the application of the line list to more metal-poor Cepheids of Magellanic Clouds and the Local Group.

In conclusion, our results suggest that a single analytic function is not sufficient to model all the gradients. For this reason, we provided the behavior of each element gradient with a GPR modeling. However, when using an analytical function, a log-log profile should be preferred. Our results confirm that the radial abundance gradients in the thin disk flatten with radius across all studied elements. This is consistent with theoretical expectations of inside-out disk formation and continuous chemical enrichment, where inner regions exhibit higher metallicity and alpha-enhancements due to more rapid star formation. Nevertheless, the outer disk is still poorly sampled to give a definitive conclusion.

The adoption of NLTE modeling throughout, from atmospheric parameters to elemental abundances, represents an important step forward relative to prior studies that often combined LTE atmospheric parameters with NLTE corrections. The method employed in this study could be extended to data that come from large spectroscopic surveys (4MOST⁷, [de Jong et al. 2016](#)) and from new instruments (MOONS@VLT⁸, [Cirasuolo et al. 2012](#); MAVIS@VLT⁹, [McDermid et al. 2020](#)), which will allow to increase sample sizes and extend spatial coverage, allowing for finer chemical gradient mapping, especially in regions that we are now poorly sampling such as the inner and outer disk, and improved constraints on Galactic formation models.

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⁷ <https://www.4most.eu/cms/home/>

⁸ <https://vltmoons.org/>

⁹ <https://mavis-ao.org/mavis/>

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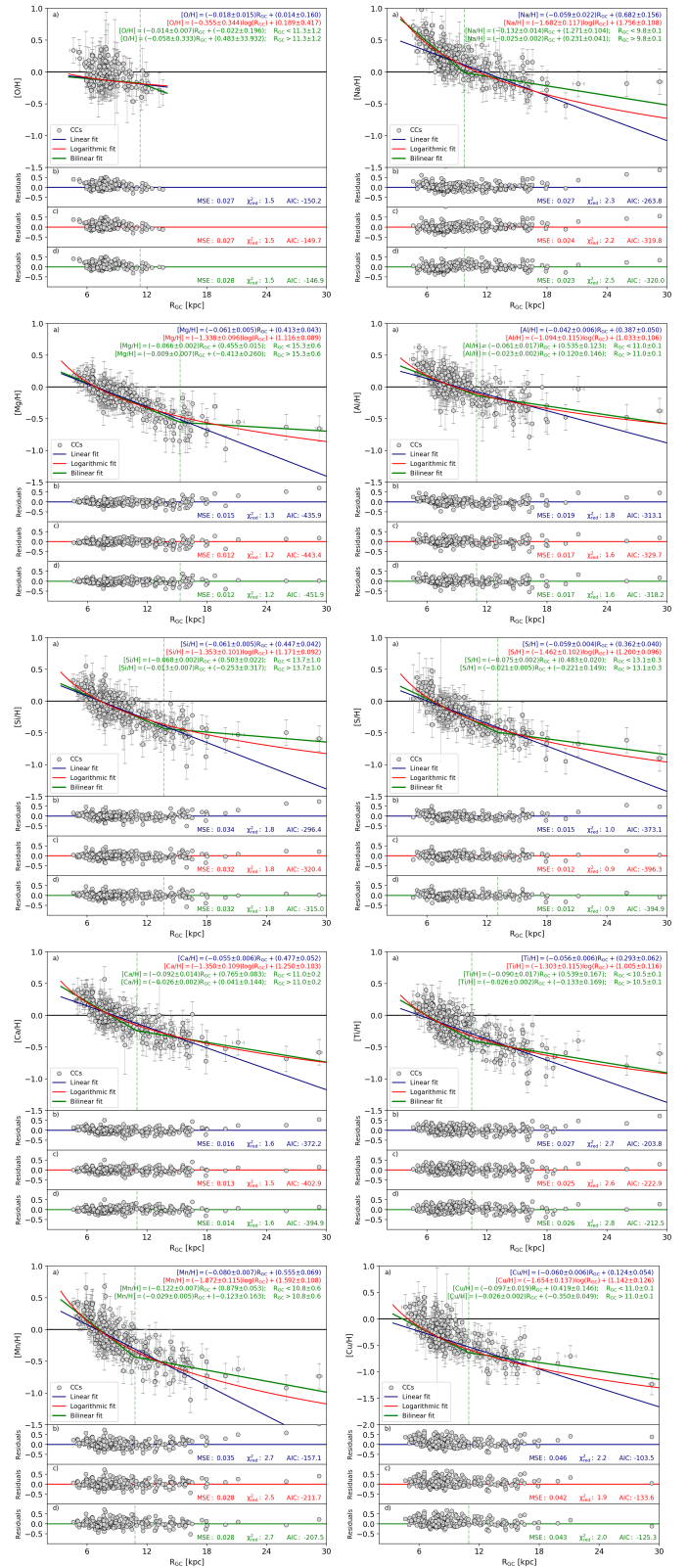


Fig. 11: Chemical gradients with linear, logarithmic and bilinear fits for O, Na, Mg, Al, Si, S, Ca, Ti, Mn and Cu. Colors and symbols are the same of Fig. 5.

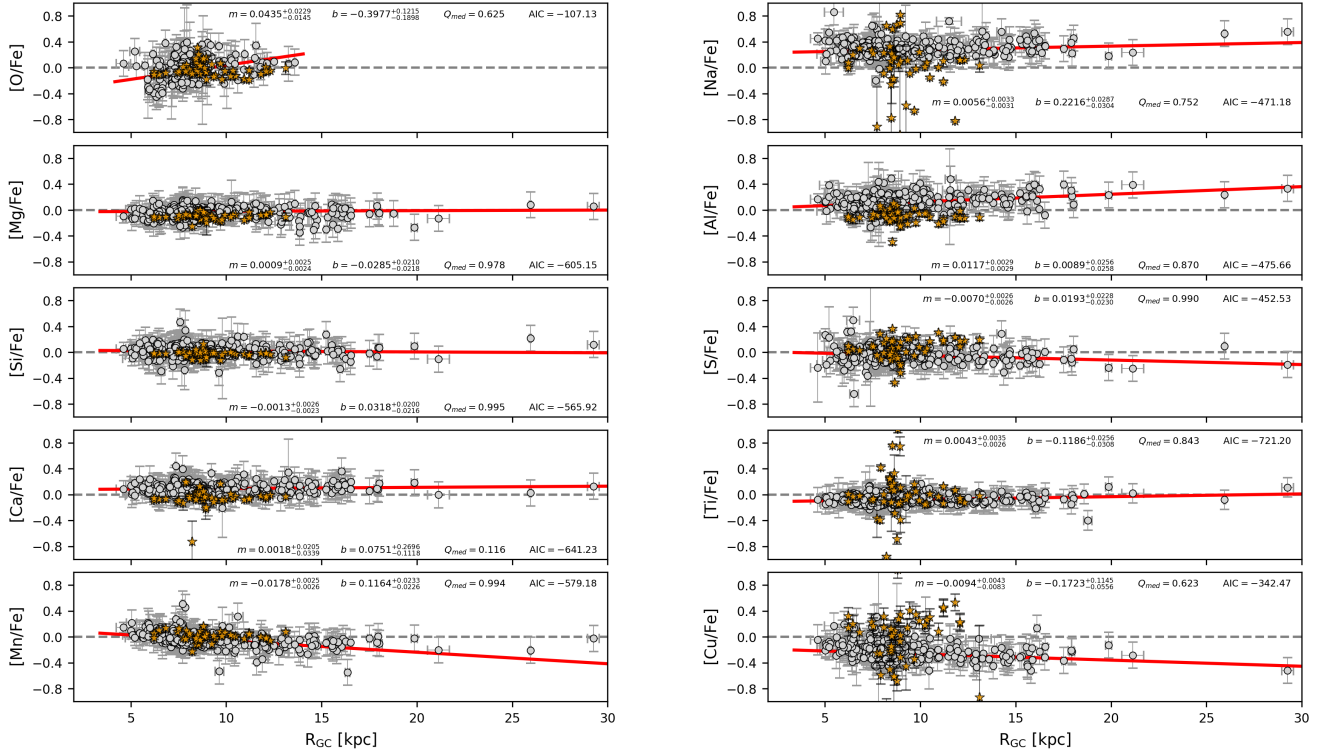


Fig. 12: Element over iron gradients for O, Na, Mg, Al, Si, S, Ca, Ti, Mn and Cu. In gray iron abundances as a function of RGC for CCs, while orange symbols depict open clusters younger than 400 Myr from Otto et al. (2025). The red profile represents the linear fit with parameters given in the top-right box for each panel.

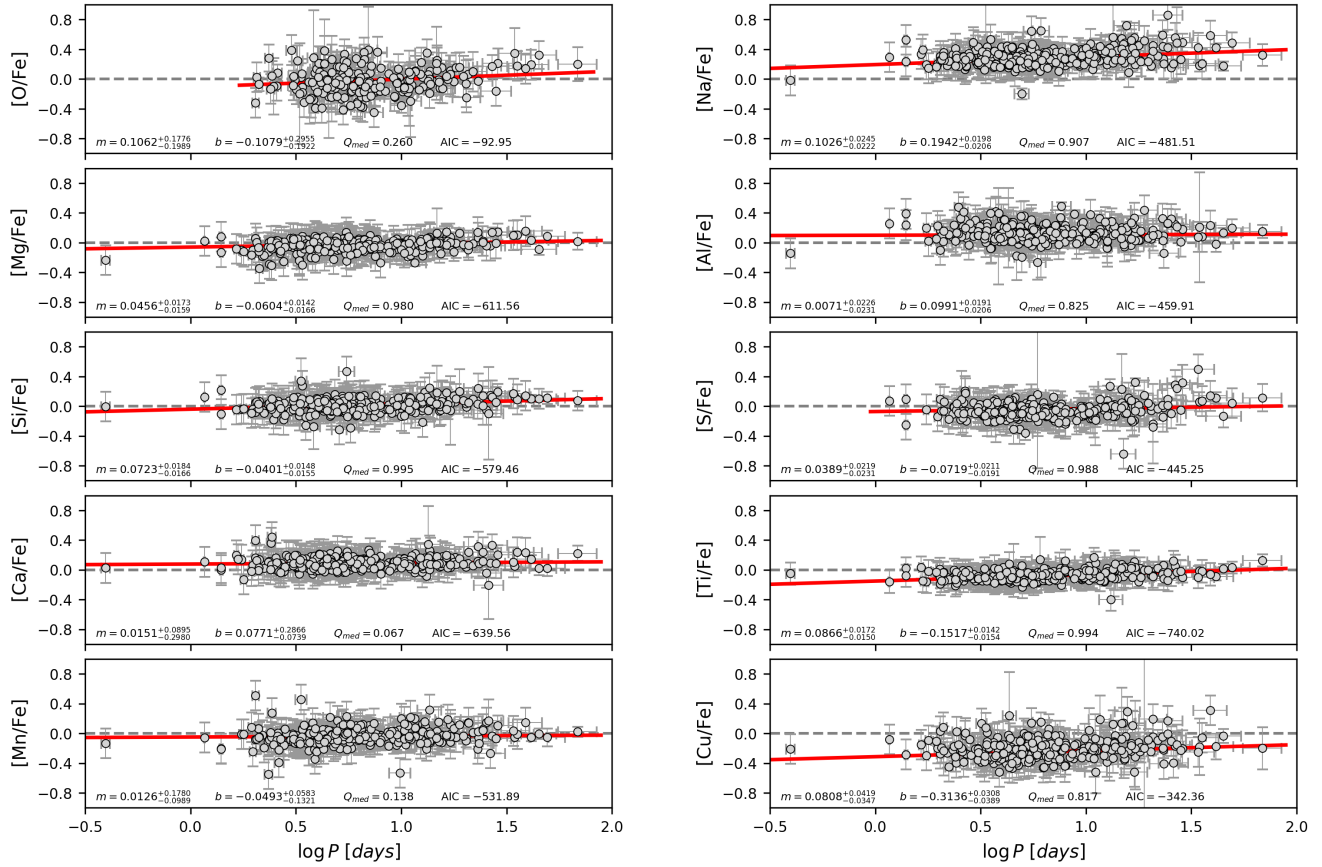


Fig. 13: Abundance ratios $[X/Fe]$ as a function of the logarithmic period for O, Na, Mg, Al, Si, S, Ca, Ti, Mn, and Cu. The red lines display the linear fit and their coefficient are plotted in the bottom left box of each panel.

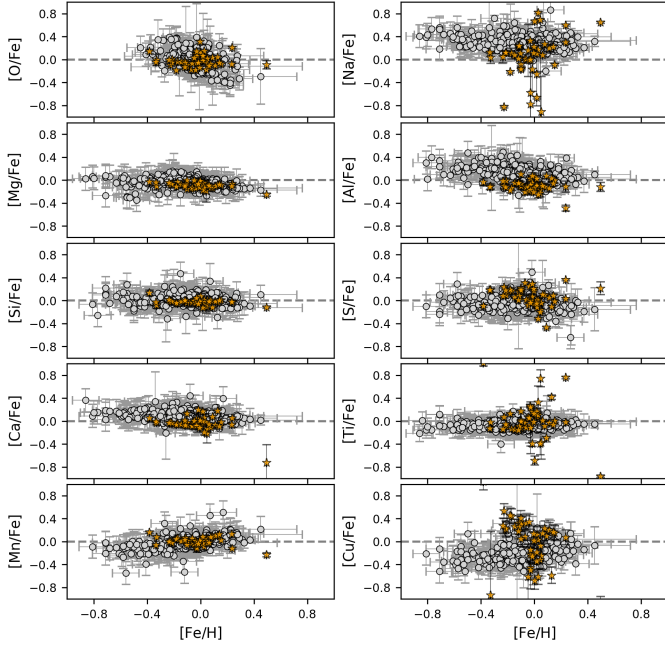


Fig. 14: Chemical planes for O, Na, Mg, Al, Si, S, Ca, Ti, Mn and Cu. Colors and symbol are the same as in Fig. 8.

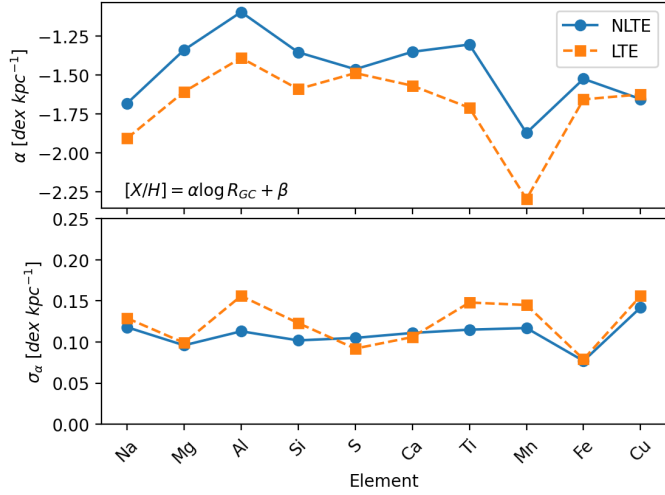


Fig. 15: NLTE (blue line) and LTE (orange line) chemical gradients slopes (α) of the logarithmic fit for each element (top panel) and their uncertainties (bottom panel).

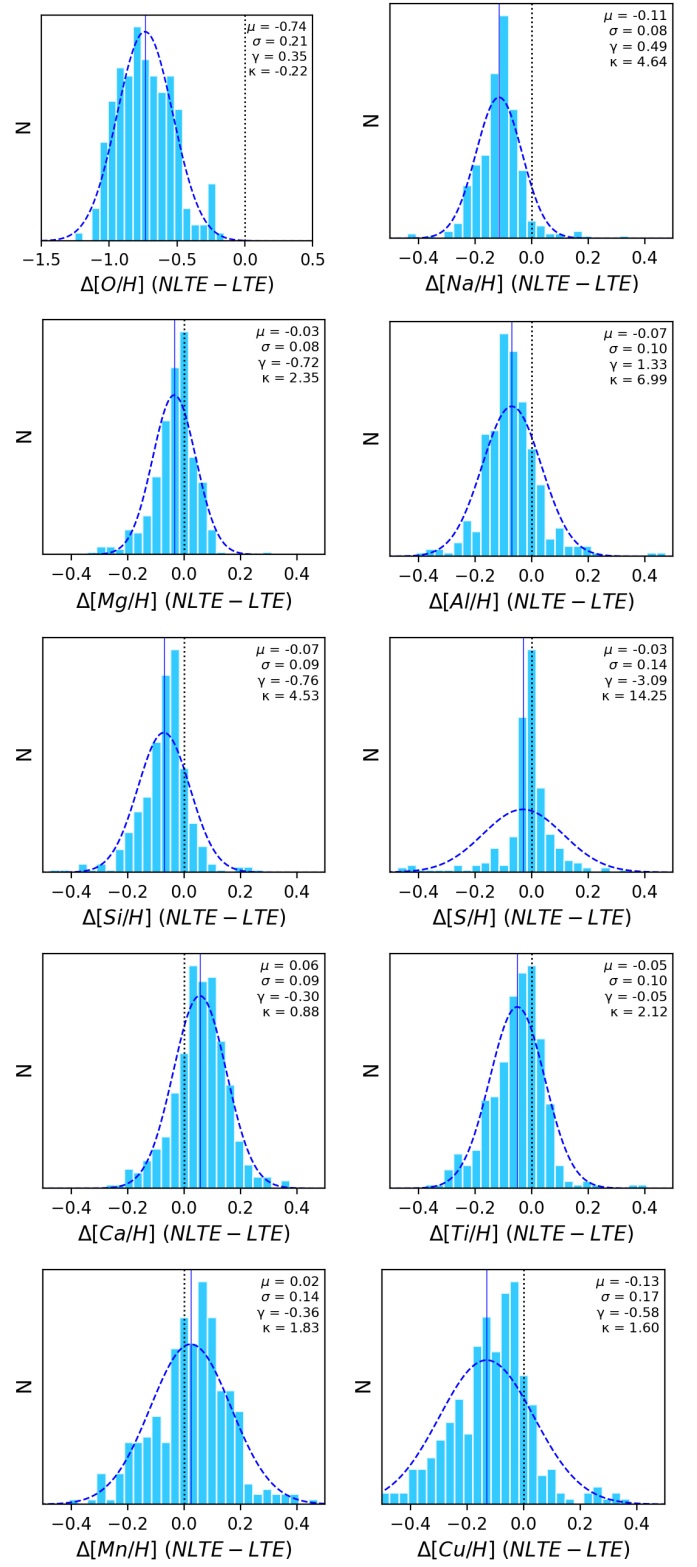


Fig. 16: Histograms of the difference between NLTE and LTE abundances. The x-axis spans the range -0.5 to $+0.5$ dex for all elements, except for oxygen, for which the range is -1.5 to $+0.5$ dex. Mean value (μ), standard deviation (σ), skewness (γ) and kurtosis (κ) of the overlaid gaussian fit (blue dashed line) are indicated in the top-right corner of each panel.

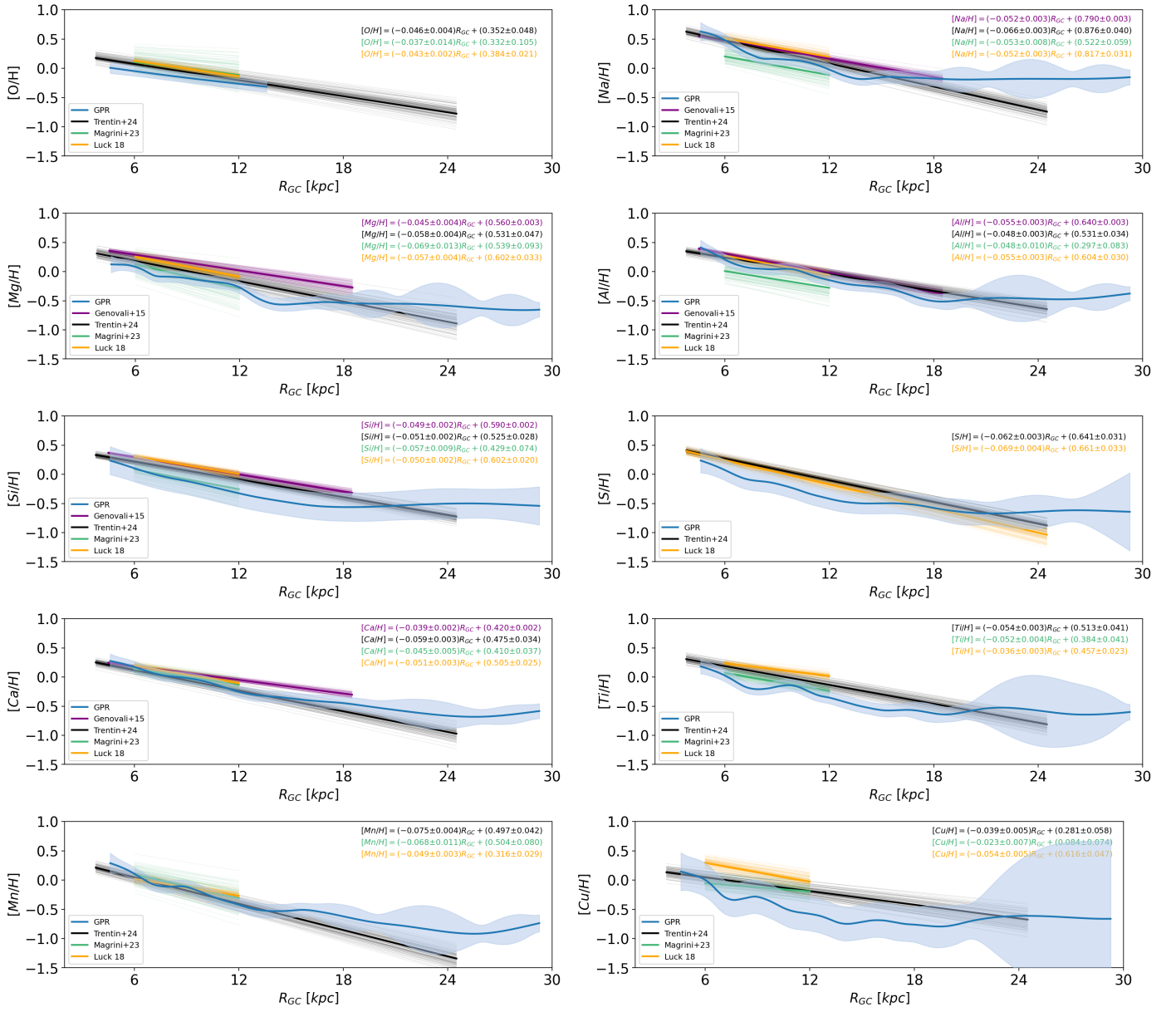


Fig. 17: Comparison of the current GPR profiles with similar findings available in the literature based either on CCs by Genovali et al. (2015), Luck (2018) and Trentin et al. (2024b) or on open clusters by Magrini et al. (2023).

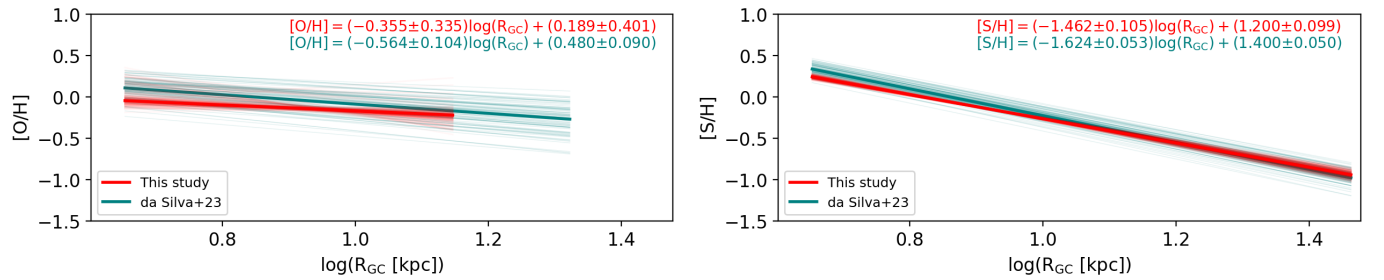


Fig. 18: Comparison of the current oxygen and sulphur radial gradients (logarithmic scale) and ds23 investigation.

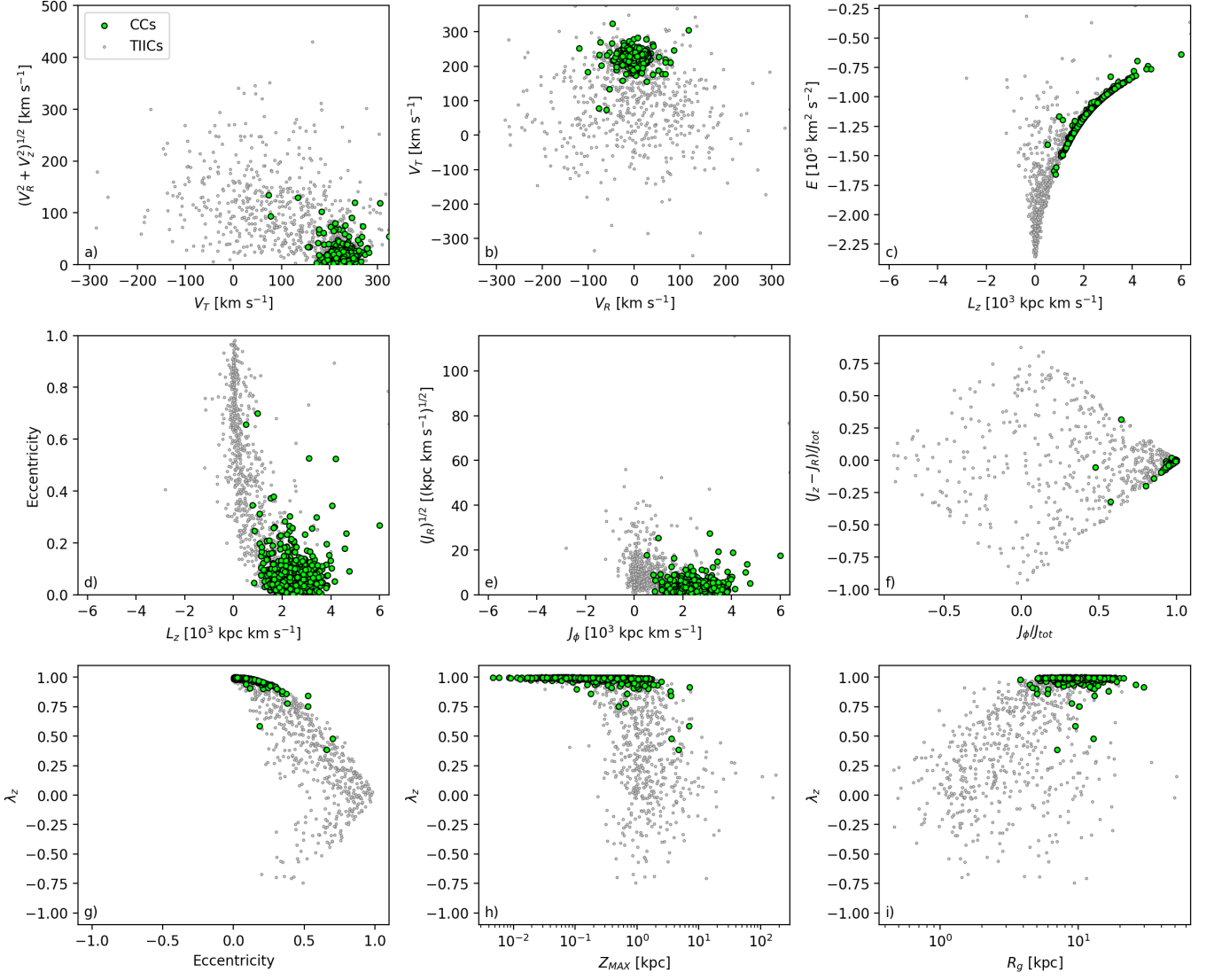


Fig. 19: Kinematic planes for CCs (green circles) and a sample of Galactic Type II Cepheids (TIICs, grey dots). Panel a)-[Toomre diagram]: V_T as a function of the sum in quadrature of V_R and V_Z . Panel b): V_T as a function of V_R . Panel c)-[Lindblad diagram]: E as a function of L_Z . Panel d): eccentricity as a function of L_Z . Panel e): the square root of J_R as a function of J_ϕ . Panel f)-[Action diamond diagram]: $(J_Z - J_R)/J_{\text{tot}}$ vs J_ϕ/J_{tot} , where $J_{\text{tot}} = |J_\phi| + J_R + J_Z$. Panel g): circularity as a function of eccentricity. Panel h): circularity as a function of Z_{MAX} . Panel i): circularity as a function of R_{GC} .

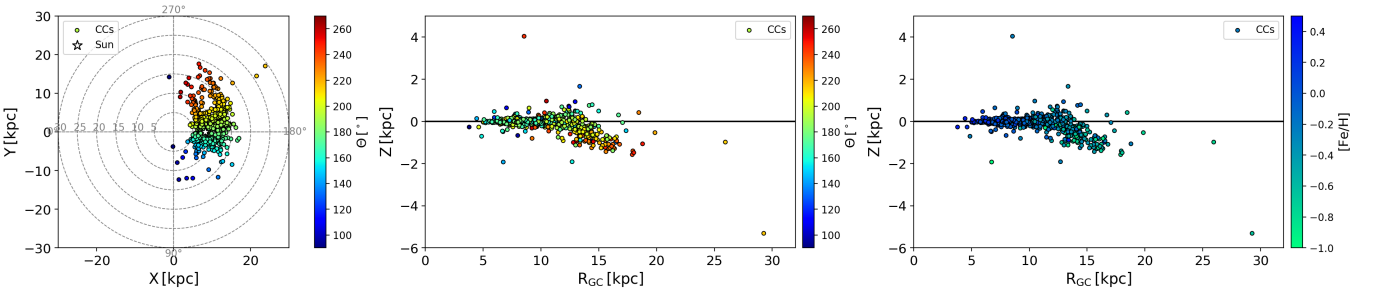


Fig. 20: Left – Distribution of the total sample of CCs onto the Galactic plane color coded according to the azimuth angle. Central – Distance perpendicular to the Galactic plane as a function of the Galactocentric distance color coded according to the azimuth angle. Right – Distance perpendicular to the Galactic plane as a function of the Galactocentric distance color coded according to the metallicity.

Appendix A: Tables

Table A.1 lists the properties of the spectroscopic sample. Lines used for atmospheric parameters, iron and titanium abundance estimation are listed in Table A.2, while those used for chemical abundances are highlighted in Table A.3. Table A.4 shows the chemical abundances of the 3 RGB stars of M67 and the Sun. Finally, table A.5 lists the values of $\Delta[X/H]$ for each spectral line computed as the mean abundance variation over 10 spectra, in response to changes in the atmospheric parameters.

Appendix B: New distance estimates for two Cepheids

The targets ASAS 181024 – 2049.6 and ASAS SN J065046.50 – 085808.7 have been estimated respectively at a distance of 5.64 kpc (Andrievsky et al. 2016) and of 10.04 kpc (Trentin et al. 2023).

Andrievsky et al. (2016) applied the method of the colour excess determination based on the use of the $E(B - V) - EW(DIB)$ calibrating relation from Friedman et al. (2011). Here $EW(DIB)$ is the equivalent width of the 6613 Å diffuse interstellar band (DIB), which is also seen in Cepheid spectra. Using this relation, they found that the $E(B - V)$ value for ASAS181024 – 2049.6 is equal to 1.41. Combining the approximate relation $A_V = E(B - V)/3.2$, the "absolute magnitude–pulsational period" relation of Gieren et al. (1998), and the mean V magnitude, they found a heliocentric distance of 5641 pc, and $R_{GC} = 2.53$ kpc. Given the $parallax_over_error < 10$, consistent with the rest of the sample, we used the Gaia parallax to obtain a heliocentric distance of 2.7 ± 0.3 kpc and $R_{GC} = 5.5$ kpc (external.gaiaedr3_distance table, Bailer-Jones et al. 2021). Given this value so different from that obtained by the other authors, we used the W1-band PL relation, calibrated from Galactic CCs by Wang et al. (2018), obtaining a distance of ~ 3.0 kpc, consistent with the value that is obtained by the parallax method.

Trentin et al. (2023) found a heliocentric distance of 10.0 ± 0.6 kpc for ASAS SN J065046.50 – 085808.7 using PW relations calibrated on Galactic CCs by Ripepi et al. (2020) and on LMC CCs by Ripepi et al. (2022). This measurement is in contrast with that of Skowron et al. (2019), who found 7.6 ± 0.3 kpc using mid-IR PL relations. In a coherent manner as described in Sec. 2.1, we found a heliocentric distance of 6.2 ± 0.7 kpc using the Gaia DR3 parallax, which is consistent with that of Skowron et al. (2019) within 2σ .

Appendix C: Spectral synthesis details

The required steps for computing atmospheric parameters by using PySME are the following:

1. Observed spectrum: with a specific spectral resolution.
2. Atomic linelist: in a specified format, i.e. VALD.
3. Atmosphere model: we utilized MARCS12 (Gustafsson et al. 2008).
4. NLTE grids: we used Amarsi et al. (2020) grids.
5. Initial chemical abundances: we used those from Asplund et al. (2009).
6. Initial set of atmospheric parameters.
7. Spectral intervals (Masks) around the spectral lines used for parameter estimates with an identification of continuum and line regions.

PySME is designed to work in conjunction with the Vienna Atomic Line Database (VALD, Ryabchikova et al. 2015), which provides a comprehensive collection of atomic and molecular transition parameters of astronomical interest. In order to have good estimators we selected spectral lines with accurate atomic parameters. We updated the $\log(gf)$ values for the lines used as estimators by referring to the Gaia-ESO database (Heiter et al. 2021). Gaia-ESO provides two quality parameters for each line: *gfflag*, which assesses the accuracy of $\log(gf)$, and *synflag*, which evaluates the blending resolution between closely spaced lines. These parameters are rated as "y" (recommended use), "n" (not recommended), or "u" (undecided quality). For the line selection we need a relatively large number of estimators, possibly both neutral and ionized lines to be sensitive to the gravity variations. We selected Fe I, Fe II, Ti I, and Ti II lines with *gfflag* and *synflag* rated as "y" for their high quality. This selection includes 21 Fe I lines, 2 Fe II lines, 6 Ti I lines, and 1 Ti II line. To expand the Fe II line list, we also included lines with *gfflag* rated as "u", adding 3 additional lines. Similarly, for the Ti II line list, we included lines with *synflag* rated as "u", adding 3 more lines. Fig. C.1 displays an example of the fit of a Ti I line in a Cepheid spectrum. The complete line list is presented in Table A.2.

We compared the atmospheric parameters derived from our method and the reference values for the Sun and 16 benchmark stars in Fig. C.2.

We point out a limitation given by the atmosphere model used, since the atmospheric parameters are estimated at the edge of the grid for 22% of the spectra, as shown in Figure C.3, where a "staircase structure" artifact is in place. This implies that for the objects at the edge of the grid the surface gravity is in general overestimated, and/or the effective temperature is underestimated. As shown in Fig. C.4, we compared our values of $\log(g)$ and T_{eff} with those of *ds23* to quantify this effect. We found an average difference of -0.20 dex in $\log(g)$ and 5 K in T_{eff} . While for the temperature the underestimation is completely negligible, for the surface gravity the average difference is comparable with the errors. For this reason, our final estimations of the surface gravity uncertainties take into account this effect.

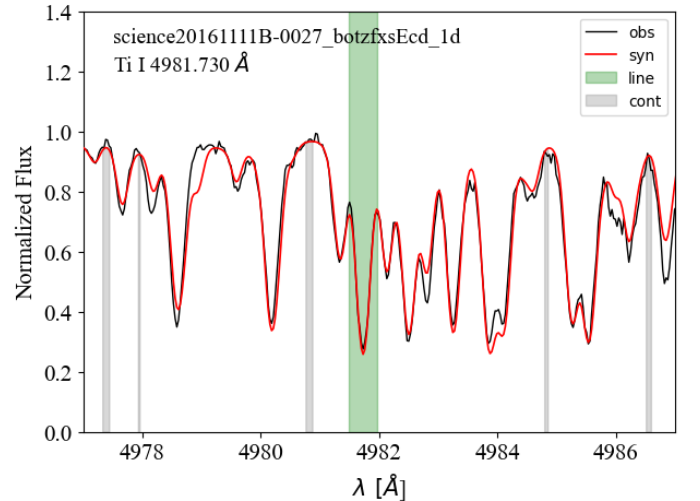


Fig. C.1: Fit of a Ti I line for a Cepheid spectrum. The normalized observed spectrum is shown in black, while the synthetic spectrum, derived by estimating the atmospheric parameters through line fitting, is displayed in red. The continuum regions are marked by gray boxes, while the green box highlights the line region where the best fit is performed.

Spectrograph	N. of spectra	N. of stars	Resolution (R)	Coverage [Å]
HARPS ¹	231	29	~ 115000	[3790-6900]
SES ²	400	64	~ 55000	[3860-8800]
FEROS ³	339	161	~ 48000	[3500-9200]
UVES ⁴	373	224	~ 40000	[3800-9400]
HARPS-N ⁵	40	29	~ 115000	[3780-6910]
ESPaDOnS ⁶	16	13	~ 81000	[3700-10500]

Table A.1: Spectroscopic sample.

¹ High Accuracy Radial velocity Planet Searcher (HARPS, [Mayor et al. 2003](#))

² STELLA échelle Spectrograph (SES, [Strassmeier et al. 2004, 2010](#))

³ The Fiber-fed Extended Range Optical Spectrograph (FEROS, [Kaufer et al. 1999](#))

⁴ Ultraviolet and Visual Echelle Spectrograph (UVES, [Dekker et al. 2000](#))

⁵ High Accuracy Radial velocity Planet Searcher for the Northern hemisphere (HARPS-N, [Cosentino et al. 2012](#))

⁶ Echelle SpectroPolarimetric Device for the Observation of Stars (ESPaDOnS, [Donati et al. 2006](#))

A double check on possible trends among the different parameters has been performed, and in Figs. C.5 the iron and titanium abundances are shown as a function of the atmospheric parameters.

Furthermore, to provide a more quantitative analysis of the differences between the NLTE and LTE approaches, we made use of the large number of HR spectra collected by our group for 20 calibrating Cepheids ([da Silva et al. 2022](#)). For these objects, the HR spectra cover the entire pulsation cycle, making them an excellent benchmark for validating spectroscopic approaches. The reason is twofold. a) Classical Cepheids with moderate to large amplitudes show variations along the pulsation cycle of approximately 1000 K in T_{eff} , about half a dex in $\log g$, and a factor of two in v_{mic} (from 2 to 4 km s⁻¹). However, the chemical abundances for each object are the same. This makes them a robust laboratory for constraining the physical assumptions involved when transitioning from LTE to NLTE analyses. b) Classical Cepheids trace a loop in both the color-magnitude diagram and the spectroscopic colour-magnitude diagram (or Kiel diagram; T_{eff} vs. $\log g$; see [Kudritzki et al. 2020, 2024](#)). As a consequence, during the pulsation cycle a Cepheid can reach, at a fixed surface gravity, two different effective temperatures: one on the rising branch and one on the declining branch. This provides a unique opportunity to test the impact of T_{eff} on NLTE analysis while keeping both surface gravity and chemical composition fixed.

We focused our analysis on three Cepheids for which we have very detailed phase coverage: ζ Gem, (fundamental mode; Fig. C.6), δ Cep (fundamental mode; Fig. C.7), and FF Aql (first overtone; Fig. C.8). The bottom four panels of these figures show the differences in atmospheric parameters and iron abundances between the NLTE and LTE analyses. These panels reveal overall consistency between the two approaches, as no significant trends or systematic offsets are observed. The same conclusion holds when comparing the LTE results from [da Silva et al. \(2022\)](#) with our current NLTE analysis (top four panels of the same figures).

Appendix D: Validation of the NLTE parameters and abundances

We compared our NLTE atmospheric parameters and chemical abundances with previous studies that investigated the impact of NLTE analysis, focusing in particular on the behavior of NLTE-LTE differences as a function of metallicity. [Ruchti et al. \(2013\)](#) introduced an optimized methodology, termed NLTE-

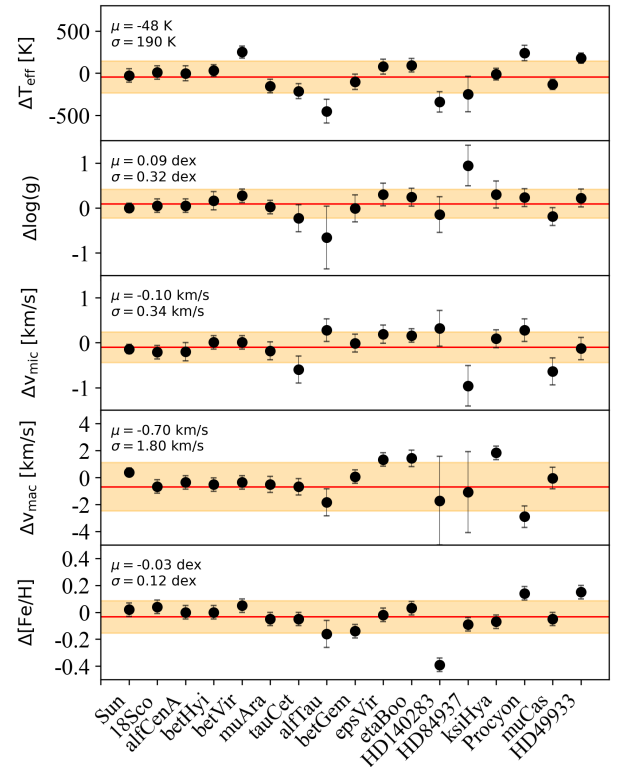


Fig. C.2: Comparison of the atmospheric parameters derived with our method and the reference values for the Sun and 16 Gaia benchmark stars from [Blanco-Cuaresma et al. \(2014\)](#). The black dots represent the differences, with error bars indicating uncertainties. The shaded regions are $\pm 1\sigma$ around the mean difference value μ (red solid line).

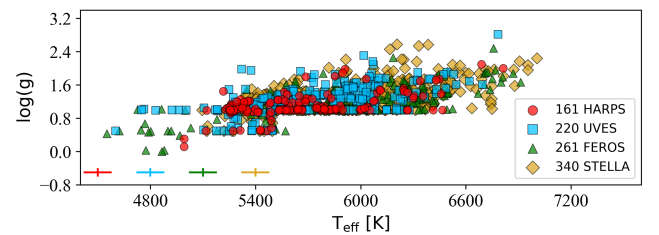


Fig. C.3: Surface gravity as a function of the effective temperature. The "staircase structure" shows the limit of the grid of the atmospheric model.

El	Ion. state	λ [Å]	gfflag	synflag	E_{low} [eV]
Fe	1	5217.389	y	y	3.211
Fe	1	5364.871	y	y	4.446
Fe	1	5569.618	y	y	3.417
Fe	1	5576.089	y	y	3.430
Fe	1	5638.262	y	y	4.220
Fe	1	5679.023	y	y	4.652
Fe	1	5705.464	y	y	4.301
Fe	1	5778.453	y	y	2.588
Fe	1	5855.076	y	y	4.608
Fe	1	5883.816	y	y	3.960
Fe	1	5916.247	y	y	2.453
Fe	1	5956.694	y	y	0.859
Fe	1	6056.005	y	y	4.733
Fe	1	6151.617	y	y	2.176
Fe	1	6165.360	y	y	4.143
Fe	1	6240.646	y	y	2.223
Fe	1	6246.318	y	y	3.603
Fe	1	6315.811	y	y	4.076
Fe	1	6336.823	y	y	3.686
Fe	1	6380.743	y	y	4.186
Fe	1	6494.980	y	y	2.404
Fe	2	4993.350	u	y	2.807
Fe	2	5264.802	y	y	3.231
Fe	2	6084.102	u	y	3.200
Fe	2	6149.246	u	y	3.889
Fe	2	6432.676	y	y	2.891
Ti	1	4913.613	y	y	1.873
Ti	1	4981.730	y	y	0.848
Ti	1	5219.702	y	y	0.021
Ti	1	5689.460	y	y	2.297
Ti	1	5978.541	y	y	1.873
Ti	1	6091.171	y	y	2.267
Ti	2	4874.009	y	u	3.095
Ti	2	5211.530	y	u	2.590
Ti	2	5381.022	y	u	1.566
Ti	2	5418.768	y	y	1.582

Table A.2: Selected linelist used to estimate the atmospheric parameters of the CCs sample. The column labeled 'Ionization state' indicates whether the line corresponds to neutral atoms (value "1") or first-ionized species (value "2").

Opt, specifically designed to mitigate the systematic biases typical of standard 1D LTE analyses for FGK-type stars. In this approach, T_{eff} is determined from Balmer lines profiles, while $\log g$ and $[\text{Fe}/\text{H}]$ are derived from the NLTE ionization balance

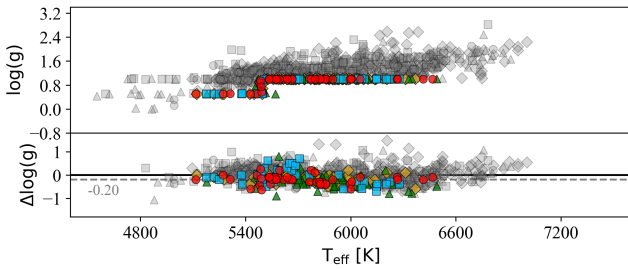


Fig. C.4: The top panel shows the same distribution of Fig. C.3 in gray. Values at the limit of the grid are shown with the same color-code of Fig. C.3. The bottom panel highlights the difference between this study and $\Delta \log g$ values. The gray dotted-line highlights $\Delta \log g = -0.20$ dex.

El	Ion. state	λ [Å]	gfflag	synflag	E_{low} [eV]
O	1	7771.944			9.146
O	1	7774.166			9.146
O	1	7775.388			9.146
Mg	1	5711.088	u	y	4.346
Mg	1	6319.237	u	y	5.108
Si	1	5645.613	y	u	4.930
Si	1	5665.555	y	u	4.920
Si	1	5684.484	y	u	4.954
Si	1	5701.104	y	y	4.930
Si	2	6371.371	y	u	8.121
Ca	1	5260.387	y	y	2.521
Ca	1	5588.749	y	y	2.526
Ca	1	5590.114	y	y	2.521
Ca	1	5867.562	y	y	2.933
Ca	1	6166.439	y	y	2.521
Ca	1	6455.598	y	y	2.523
Ca	1	6471.662	y	y	2.526
Ca	1	6499.650	y	y	2.523
S	1	6743.540	u	u	7.866
S	1	6757.150	u	u	7.870
Mn	1	5394.683	y	y	0.000
Mn	1	5420.328	y	y	2.143
Mn	1	6021.771	y	y	3.075
Al	1	5557.063	u	u	3.143
Al	1	6696.023	u	u	3.143
Al	1	6698.673	u	y	3.143
Na	1	5688.205	u	y	2.104
Na	1	6154.226	u	y	2.102
Na	1	6160.747	u	y	2.104
Cu	1	5105.530	y	u	1.389
Cu	1	5218.195	y	u	3.817
Cu	1	5220.065	y	u	3.817
Cu	1	5700.170	y	u	1.642

Table A.3: Selected linelist used to estimate the chemical abundances of the CCs sample. The column labeled 'Ionization state' indicates whether the line corresponds to neutral atoms (value "1") or first-ionized species (value "2"). For oxygen lines, the information on "gfflag" and "synflag" flags is not available.

of Fe. Fig. D.1 shows the systematic offsets in stellar parameters as a function of metallicity based on standard LTE analysis when compared with NLTE-Opt estimates. These trends are generally attributed to the inability of the LTE Fe I excitation balance to provide an accurate T_{eff} scale. As shown by the black circles representing the results of Ruchti et al. (2013), these differences decrease moving toward higher metallicities (see also Kovalev et al. 2019). Our Cepheids (in blue) follow the same general behavior, displaying a good agreement with Ruchti et al. (2013), particularly in ΔT_{eff} (panel a), $\Delta \log g$ (panel b) and $\Delta [\text{Fe}/\text{H}]$ (panel d). However, the panel c) of the same figure shows the difference in microturbulence differences, revealing an offset of $0.20\text{--}0.30 \text{ km s}^{-1}$. Since Cepheids typically exhibit higher micro- and macroturbulent velocities, both the observed trends and the intrinsic scatter should be interpreted with caution. We also note that v_{mic} is the only parameter for which we find a statistically significant offset between the NLTE and LTE determinations ($-0.15 \pm 0.11 \text{ km s}^{-1}$, see Fig. 2). Finally, our fitting method simultaneously optimizes all atmospheric parameters, which means that the contribution of v_{mic} significantly affects the determination of the others.

Target			Mg	eMg	Si	eSi	Ca	eCa	S	eS
08512898+1150330			0.01	0.08	-0.02	0.14	-0.08	0.05	0.04	0.10
08511268+1152422			-0.01	0.07	-0.04	0.05	-0.01	0.04	0.08	0.10
08512283+1148015			0.01	0.08	-0.01	0.14	-0.03	0.06	-0.01	0.10
Sun			0.04	0.03	0.00	0.02	0.00	0.03	-0.05	0.10
	Ti	eTi	Mn	eMn	Al	eAl	Na	eNa	Cu	eCu
08512898+1150330	-0.07	0.05	-0.14	0.02	-0.09	0.07	-0.08	0.06	0.02	0.08
08511268+1152422	-0.04	0.05	-0.10	0.05	-0.07	0.08	-0.11	0.05	0.00	0.01
08512283+1148015	-0.09	0.06	-0.14	0.03	-0.09	0.06	-0.08	0.03	-0.07	0.06
Sun	-0.04	0.04	0.00	0.01	-0.01	0.02	-0.03	0.05	-0.04	0.02

Table A.4: $[X/Fe]$ values with errors of 3 RGB stars of M67 and the Sun.

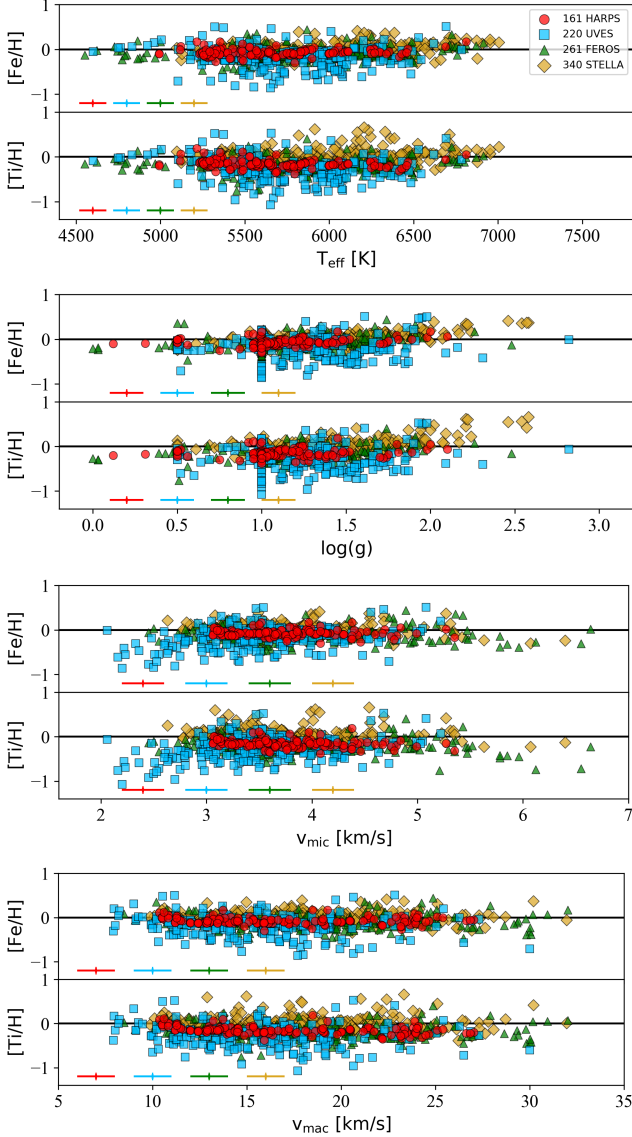


Fig. C.5: Iron and titanium abundances as a function of the atmospheric parameters. Colors and symbols are the same of Fig. 3.

We independently verified our NLTE abundances by applying the line-by-line NLTE abundance corrections provided by the MPIA database¹⁰ (Kovalev et al. 2018). Specifically, we focused on Mn and O. The atomic model for Mn is the same as

¹⁰ <https://nlte.mpia.de/>

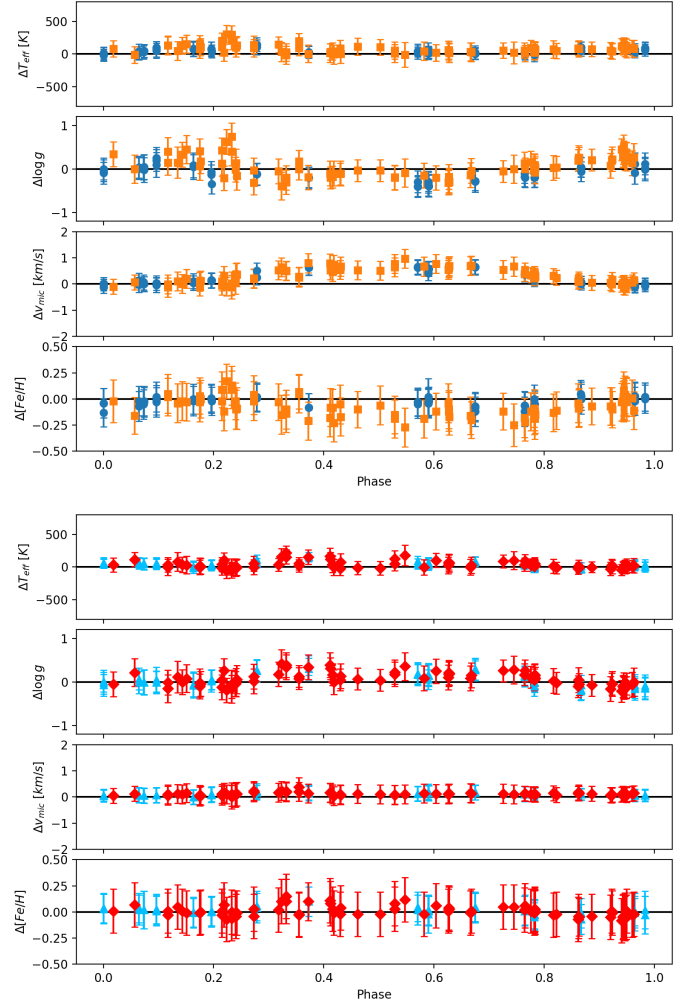


Fig. C.6: Top four panels – Difference in atmospheric parameters between the current analysis and da Silva et al. (2022) as a function of pulsation phase for the calibrating Cepheid ζ Gem. Dark symbols indicate STELLA spectra, whereas orange symbols mark HARPS spectra. Black solid lines display zero difference. Bottom four panels – same as the top four panels, but the difference is between NLTE and LTE estimates of the atmospheric parameters. Light blue symbols indicate STELLA spectra, whereas red symbols mark HARPS spectra.

that used in our study from Bergemann et al. (2019), while for O, we used the atomic model from Amarsi et al. (2018), while the MPIA database relies on Bergemann et al. (2021). We performed this comparison for five sample spectra whose atmospheric parameters are representative of the whole population.

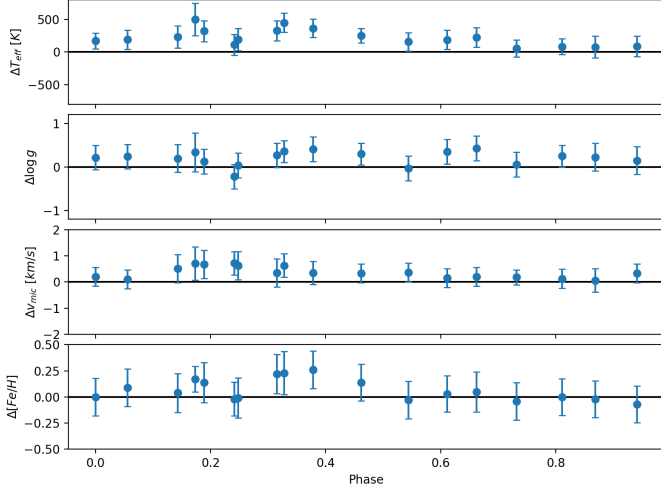


Fig. C.7: Same as Fig. C.6, but for δ Cep.

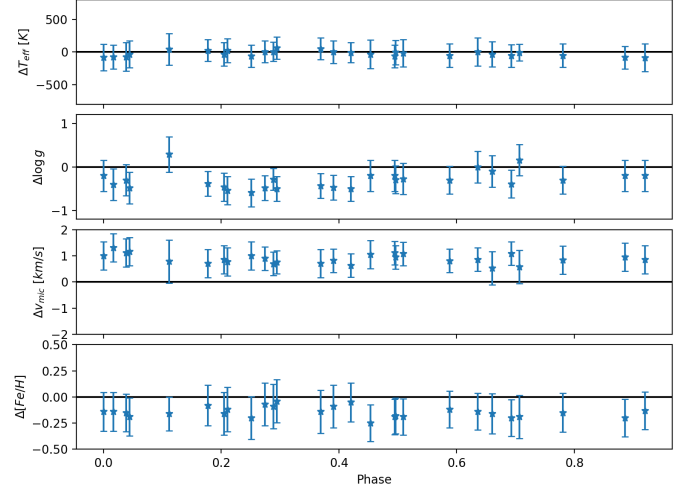


Fig. C.8: Same as Fig. C.6, but for FF Aql.

We downloaded the NLTE abundance correction from the MPIA database for Mn (5394.68 Å, 5420.33 Å and 6021.77 Å) and O (7771.94 Å, 7774.17 Å and 7775.39 Å). As shown in Fig. D.2, the comparison indicates good agreement for all Cepheids, especially for Mn. For oxygen the agreement is also good (differences within 0.1 dex) for three Cepheids, while the remaining two show discrepancies of approximately 0.2 dex. This is well within our observational uncertainties. Finally, the decreasing trend observed in the plane $\langle \Delta[\text{O}/\text{H}] \rangle$ vs T_{eff} has also been obtained in Vasilyev et al. (2019) for the 777.4 nm line with both 1D and 2D models, and previously observed in Cepheids in Luck et al. (2013).

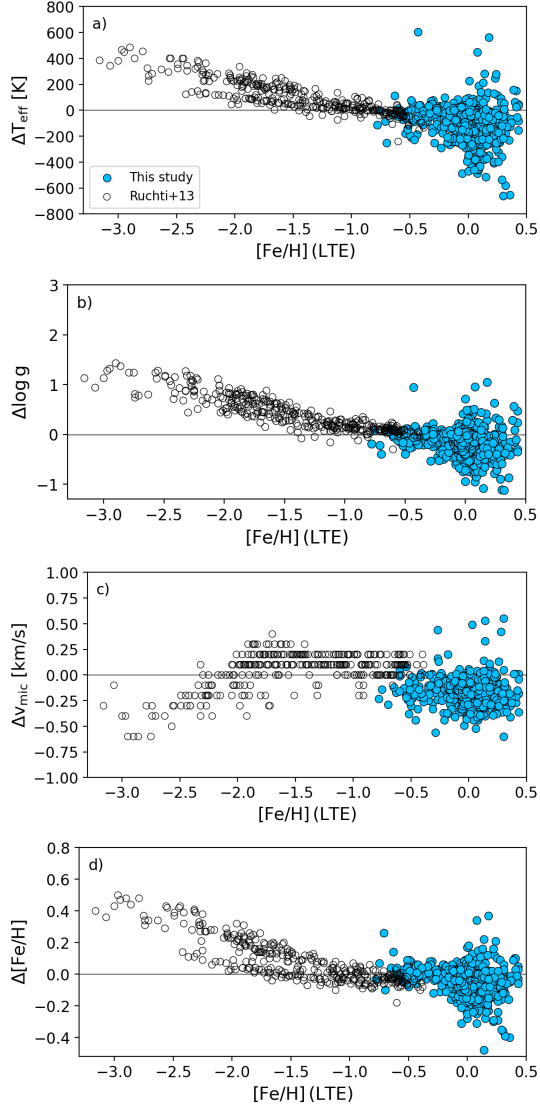


Fig. D.1: Differences between NLTE and LTE estimates of T_{eff} (panel a), $\log g$ (panel b), v_{mic} (panel c) and $[\text{Fe}/\text{H}]$ as a function of LTE-metallicity. Ruchti et al. (2013) results are shown with black circles, while this work's results are highlighted in blue.

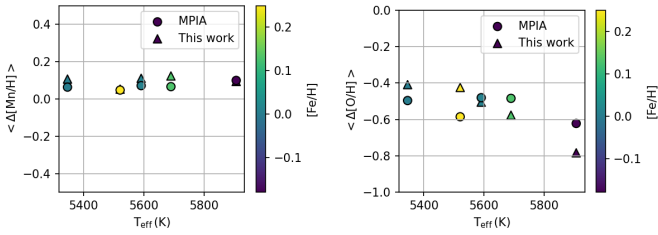


Fig. D.2: Comparison of the mean NLTE abundance corrections provided by the MPIA database (<https://nlte.mpia.de/index.php>) and this study's results ($\Delta[X/H] = [X/H]_{\text{NLTE}} - [X/H]_{\text{LTE}}$) for manganese (left panel) and oxygen (right panel) as a function of T_{eff} and color-coded by $[\text{Fe}/\text{H}]$.

λ [Å]	$\Delta[X/H]$	$\Delta\log(g)$ [dex]			$\Delta[Fe/H]$ [dex]			ΔT_{eff} [K]			$\Delta\xi$ [km s ⁻¹]						
		-0.1	-0.2	+0.1	+0.2	-0.05	-0.1	+0.05	+0.1	-100	-50	+100	+50	-0.2	-0.4	+0.2	+0.4
4874.009	Ti	0.111	0.070	0.256	0.291	0.280	0.338	0.162	0.105	0.192	0.207	0.252	0.237	0.300	0.375	0.154	0.095
5105.530	Cu	-0.002	-0.007	0.009	0.017	0.048	0.092	-0.034	-0.074	-0.074	-0.037	0.095	0.046	0.033	0.063	-0.017	-0.041
5211.530	Ti	0.117	0.087	0.190	0.220	0.201	0.238	0.119	0.082	0.149	0.153	0.174	0.166	0.212	0.267	0.112	0.078
5218.195	Cu	-0.012	0.001	-0.006	-0.010	0.055	0.125	-0.072	-0.126	-0.088	-0.043	0.062	0.027	-0.004	0.034	-0.018	-0.014
5219.702	Ti	0.150	0.151	0.161	0.156	0.167	0.168	0.150	0.129	0.099	0.137	0.177	0.172	0.164	0.163	0.157	0.155
5220.065	Cu	-0.057	-0.162	0.173	0.249	0.227	0.321	0.009	-0.165	-0.009	0.036	0.316	0.183	0.121	0.124	-0.013	-0.075
5260.387	Ca	-0.024	-0.043	0.012	0.026	0.054	0.109	-0.080	-0.157	-0.031	-0.026	0.019	0.010	-0.003	-0.011	-0.010	-0.016
5381.022	Ti	-0.223	-0.239	-0.138	-0.129	-0.135	-0.125	-0.178	-0.200	-0.165	-0.159	-0.143	-0.148	-0.129	-0.113	-0.201	-0.239
5394.683	Mn	-0.013	-0.018	0.005	0.012	0.030	0.061	-0.043	-0.084	-0.112	-0.054	0.099	0.053	-0.001	-0.002	0.007	0.000
5418.768	Ti	-0.039	-0.072	0.043	0.078	0.052	0.093	-0.033	-0.073	-0.016	-0.005	0.030	0.020	0.088	0.137	-0.061	-0.119
5420.328	Mn	-0.003	-0.007	0.006	0.005	0.053	0.102	-0.049	-0.098	-0.083	-0.041	0.094	0.049	0.006	0.010	-0.001	-0.004
5557.063	Al	-0.035	-0.038	-0.011	-0.004	0.041	0.091	-0.062	-0.110	-0.057	-0.042	0.044	0.007	-0.020	-0.016	-0.013	-0.009
5588.749	Ca	-0.005	-0.006	0.005	0.014	0.060	0.120	-0.066	-0.127	-0.066	-0.036	0.067	0.032	0.142	0.304	-0.130	-0.227
5590.114	Ca	0.031	0.033	0.024	0.033	0.086	0.140	-0.027	-0.084	-0.045	-0.017	0.115	0.070	0.070	0.130	-0.035	-0.055
5645.613	Si	0.008	0.011	0.004	0.002	0.067	0.126	-0.057	-0.098	-0.059	-0.028	0.063	0.037	0.017	0.026	-0.003	-0.005
5665.555	Si	0.000	0.001	-0.001	0.000	0.047	0.096	-0.039	-0.073	-0.046	-0.024	0.048	0.024	0.008	0.016	-0.004	-0.006
5684.484	Si	-0.027	0.005	0.010	0.017	0.054	0.085	-0.065	-0.134	-0.067	-0.030	0.054	0.031	0.035	0.068	-0.028	-0.057
5688.205	Na	-0.002	0.006	-0.001	0.002	0.085	0.165	-0.066	-0.124	-0.064	-0.036	0.054	0.025	0.062	0.154	-0.061	-0.113
5700.170	Cu	0.005	-0.031	0.008	0.028	0.127	0.250	-0.153	-0.306	-0.103	-0.065	0.123	0.059	0.019	0.023	0.013	0.005
5701.104	Si	0.007	0.006	0.002	0.001	0.051	-0.012	-0.032	-0.072	-0.034	-0.014	0.046	0.025	0.011	0.019	-0.002	-0.006
5711.088	Mg	1.010	1.009	0.709	0.709	0.749	0.791	0.670	0.631	0.673	0.690	0.744	0.725	0.732	0.761	0.683	0.664
5867.562	Ca	0.018	0.021	0.007	0.003	0.057	0.093	-0.018	-0.042	-0.020	-0.008	0.061	0.035	0.010	0.013	0.016	0.024
6021.771	Mn	0.014	0.018	0.007	0.006	0.062	0.120	-0.050	-0.100	-0.094	-0.043	0.108	0.058	0.037	0.070	-0.016	-0.037
6154.226	Na	0.007	0.015	-0.009	-0.016	0.050	0.101	-0.046	-0.088	-0.057	-0.029	0.053	0.026	0.007	0.016	-0.007	-0.011
6160.747	Na	0.008	0.013	-0.006	-0.010	0.057	0.113	-0.053	-0.103	-0.054	-0.027	0.052	0.025	0.020	0.040	-0.017	-0.032
6166.439	Ca	-0.002	0.031	0.008	0.016	0.048	0.107	-0.033	-0.094	-0.035	-0.031	0.067	0.033	0.031	0.064	-0.026	-0.022
6319.237	Mg	1.064	1.072	0.745	0.740	0.801	0.851	0.695	0.662	0.696	0.719	0.799	0.779	0.752	0.763	0.741	0.737
6371.371	Si	-0.323	-0.173	-0.001	0.041	-0.159	0.064	-0.104	-0.358	0.016	-0.014	-0.223	-0.253	-0.132	0.006	-0.208	-0.177
6455.598	Ca	-0.006	-0.013	0.016	0.027	0.073	0.137	-0.054	-0.133	-0.054	-0.024	0.072	0.037	0.016	0.029	-0.013	0.003
6471.662	Ca	0.014	0.003	0.028	0.045	0.089	0.148	-0.048	-0.083	-0.066	-0.033	0.123	0.074	0.071	0.123	-0.022	-0.057
6499.650	Ca	0.032	0.046	0.017	0.022	0.071	0.124	-0.031	-0.081	-0.037	-0.016	0.089	0.059	0.046	0.084	-0.000	-0.018
6696.023	Al	0.001	-0.001	0.000	-0.005	0.073	0.138	-0.066	-0.126	-0.076	-0.039	0.082	0.041	0.009	0.017	-0.005	-0.011
6698.673	Al	-0.003	-0.004	-0.001	-0.007	0.054	0.114	-0.063	-0.110	-0.073	-0.037	0.066	0.029	0.004	0.008	-0.009	-0.013
6743.540	S	-0.205	-0.315	0.111	0.201	-0.118	0.035	0.147	-0.176	-0.120	0.064	-0.076	0.084	0.085	-0.097	-0.238	-0.031
6757.150	S	-0.200	-0.173	-0.101	-0.081	-0.094	-0.049	-0.063	-0.300	0.016	0.008	-0.166	-0.133	-0.083	-0.014	-0.086	-0.094

Table A.5: This table lists the values of $\Delta[X/H]$ for each spectral line, computed as the mean abundance variation over 10 spectra, in response to changes in the atmospheric parameters (top row of the table). These values quantify the sensitivity of each line to the individual parameter variations.