

Gravitational potential drives the concentration dependence of the stellar mass-halo mass relation

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ABSTRACT

We investigate the origin of the scatter in the stellar mass-halo mass (SMHM) relation using the COLIBRE cosmological hydrodynamical simulations. At fixed halo mass, we find a clear positive correlation between stellar mass and halo concentration, particularly in low-mass haloes between 10^{11} and $10^{12} M_{\odot}$, where all halo properties are computed from the corresponding dark-matter-only simulation. Two scenarios have been proposed to explain this trend: the earlier formation of higher-concentration haloes allows more time for star formation, or the deeper gravitational potential wells of higher-concentration haloes enhance baryon retention. To distinguish between them, we examine correlations between halo concentration, stellar mass, stellar age, and stellar metallicity. While, at fixed halo mass, halo concentration correlates with stellar age, stellar age itself shows only a weak correlation with stellar mass, indicating that early formation alone cannot account for the concentration-dependence in the scatter of the SMHM relation. In contrast, both stellar metallicity and halo concentration exhibit correlations with stellar mass. The connection between halo concentration and stellar metallicity persists even when simultaneously controlling for both halo mass and stellar mass. These results support the scenario in which the deeper gravitational potentials in higher-concentration haloes suppress feedback-driven outflows, thereby enhancing both baryon and metal retention.

Key words: galaxies: formation - galaxies: evolution - methods: statistical - galaxies: groups: general - dark matter

1 INTRODUCTION

The connection between galaxies and their host dark matter haloes is a cornerstone of modern galaxy formation theory. In the Λ CDM cosmological framework, galaxies form through the condensation of baryons within gravitational potential wells of dark matter haloes (White & Frenk 1991), whose assembly is governed by hierarchical growth. A central empirical relation that encodes this connection is the stellar mass-halo mass (SMHM) relation, which captures the efficiency with which haloes convert baryons into stars (e.g. Eke

et al. 2004; Yang et al. 2003; Moster et al. 2010; Wechsler & Tinker 2018; Behroozi et al. 2019). While the SMHM relation is relatively tight on average, significant intrinsic scatter persists at fixed halo mass in both observations (e.g. Yang et al. 2009; More et al. 2009; Zu & Mandelbaum 2015; Kravtsov et al. 2018; Scholz-Díaz et al. 2022, 2024) and simulations (e.g. Gu et al. 2016; Tojeiro et al. 2017; Matthee et al. 2017; Artale et al. 2018; Correa & Schaye 2020; Lyu et al. 2023; Pei et al. 2024). Understanding the origin of this scatter is crucial for understanding the connection between halo assembly and galaxy formation (e.g. Mandelbaum et al. 2006; Wang et al. 2013; Cui et al. 2021; Zu et al. 2021; Wang et al. 2023a; Zhao et al. 2025; Wang & Peng 2025), and for developing empirical models

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for galaxy formation (e.g. Reddick et al. 2013; Zehavi et al. 2018, 2019).

Previous simulation-based studies have highlighted the role of secondary halo properties—such as concentration, formation time, spin, and environment—in shaping the galaxy-halo connection beyond halo mass (e.g. Matthee et al. 2017; Wechsler & Tinker 2018; Bose et al. 2019; Kulier et al. 2019; Zehavi et al. 2019; Martizzi et al. 2020; Montero-Dorta et al. 2020; Bradshaw et al. 2020). Among these, halo concentration, or equivalently the maximum circular velocity, which can be computed directly from halo mass and concentration, has emerged as a relatively strong correlate: at fixed halo mass, haloes with higher concentration tend to host central galaxies with larger stellar masses. This trend is found in empirical models (e.g. Bradshaw et al. 2020), semi-analytical models (e.g. Tojeiro et al. 2017; Zehavi et al. 2019; Pei et al. 2024), and hydrodynamical simulations (e.g. Matthee et al. 2017; Artale et al. 2018; Bose et al. 2019; Martizzi et al. 2020; Pei et al. 2024). However, the physical origin of this dependence remains uncertain.

Matthee et al. (2017) proposed two broad, physically motivated hypotheses to explain the concentration dependence of stellar mass at fixed halo mass. The first links concentration to halo formation history: haloes that assemble earlier tend to be more concentrated (e.g. Navarro et al. 1997; Wechsler et al. 2002; Zhao et al. 2009; Ludlow et al. 2014; Correa et al. 2015), and may thus have had more time to form stars. In this scenario, stellar mass differences at fixed halo mass arise primarily from temporal effects, with earlier-forming haloes building more massive galaxies due to their more extended star formation histories. The second hypothesis emphasizes gravitational potential depth: at fixed halo mass, more concentrated haloes have deeper gravitational potentials, which can reduce the ability of stellar feedback to eject gas. This mechanism leads to more efficient gas retention and thus higher stellar mass.

Disentangling these two explanations is challenging. While halo formation time and gravitational potential depth are correlated, they are not completely degenerate, and their respective imprints on galaxy properties may differ in subtle but measurable ways. For example, one might expect the formation-time-driven scenario to imprint differences in stellar ages, whereas feedback-regulated growth may more directly influence stellar metallicity through differences in gas retention and recycling (e.g. Peng & Maiolino 2014; Wang 2025).

Here we investigate the origin of the concentration dependence of the SMHM relation using the new COLIBRE cosmological hydrodynamical simulation suite (Schaye et al. 2025; Chaikin et al. 2025a). We focus on central galaxies at $z = 0$, and examine how galaxy stellar mass correlates with halo concentration, stellar age, and stellar metallicity at fixed halo mass. By comparing the relative strength of these correlations, we aim to assess whether the primary driver of the SMHM scatter is the difference in the halo formation time or the gravitational potential depth. This approach provides a physically motivated diagnostic of the two proposed mechanisms and offers a clear picture of how galaxy formation efficiency is modulated by halo structure.

The remainder of this paper is organized as follows. § 2 describes the COLIBRE simulations, the sample selection, and the physical quantities used in our analysis. In § 3, we examine the correlation between stellar mass and halo concentration and its connection to stellar age and stellar metallicity, from which we can interpret these trends in the context of the proposed scenarios. § 4 presents the summary of our conclusions. We assume the same Λ CDM cosmology as COLIBRE, which is based on the DES

Y3 results (Abbott et al. 2022), where $h = 0.681$, $\Omega_m = 0.306$, $\Omega_b = 0.0486$, $\Omega_\Lambda = 0.693922$.

2 SIMULATIONS

This study is based on the COLIBRE simulation suite, which consists of a set of hydrodynamical simulations and their corresponding dark-matter-only (DMO) counterparts with different box sizes and resolutions (Schaye et al. 2025; Chaikin et al. 2025a). The COLIBRE simulations were run with the SWIFT code (Schaller et al. 2024) with significant improvements to all subgrid prescriptions. The most important new features include allowing radiative cooling below 10^4 K (Ploekinger et al. 2025), star formation with a gravitational instability criterion (Nobels et al. 2024), improved chemical enrichment (Correa et al. 2025), simulating the growth and composition of dust grains with coupling to self-shielding and molecules (Trayford et al. 2025), modelling pre-supernova stellar feedback (Benítez-Llambay et al. 2025), and improving the sampling of supernova and AGN feedback (Chaikin et al. 2023; Schaye et al. 2025). Most of the simulations have two flavours of AGN feedback: one thermally-driven AGN feedback (Booth & Schaye 2009), and one hybrid thermally-driven/kinetic jet-driven AGN feedback (Huško et al. 2025). Spurious transfer of energy from dark matter to baryons is suppressed by using four times more dark matter particles than baryonic particles (Ludlow et al. 2023).

All COLIBRE simulations are calibrated to match the observed $z \approx 0$ galaxy stellar mass function, stellar mass-size relation, and stellar mass-black hole mass relation (Chaikin et al. 2025a). COLIBRE reproduces other properties and scaling relations that are not used in calibration, including the stellar and gas-phase metallicity relations, cosmic star formation history, galaxy quiescent fraction, atomic and molecular gas mass, X-ray luminosity from the circumgalactic medium (Schaye et al. 2025), and the evolution of the galaxy stellar mass function from $z > 10$ to $z = 0$ (Chaikin et al. 2025b).

Here we focus on the L200m6 simulation with thermally-driven AGN feedback, which has a cubic length of 200 comoving Mpc. The masses of dark matter particles and baryonic particles are $2.42 \times 10^6 M_\odot$ and $1.84 \times 10^6 M_\odot$, respectively. Haloes are identified using the Friends-of-Friends (FoF) algorithm, and subhaloes are identified and tracked using the HBT-HERONS method (Forouhar Moreno et al. 2025). Halo and galaxy properties are calculated using the SOAP pipeline (McGibbon et al. 2025). Halo masses are defined as the total mass enclosed within a sphere whose mean density is 200 times the critical density of the universe. Halo concentration, which characterizes the shape of the NFW density profile (Navarro et al. 1997), is estimated from the first moment of the dark matter density profile (Wang et al. 2024). The stellar properties of galaxies are calculated using all bound stellar particles within 50 physical kpc of the most-bound particle of each subhalo. Both stellar age and stellar metallicity are mass-weighted averages over all stellar particles for each galaxy.

To characterize the dark matter halo properties independently of baryonic effects, such as halo contraction (e.g. Blumenthal et al. 1986; Schaller et al. 2015), we follow Matthee et al. (2017) and use the corresponding DMO simulation with the same initial conditions to compute the halo mass, M_{200c}^{DMO} , and concentration, c^{DMO} . Subhaloes in the two simulations are matched by identifying the 50 most bound dark matter particles in each subhalo. This matching yields a bijective success rate of over 98% for central subhaloes

with masses above $10^{11} M_{\odot}$, ensuring a robust mapping between halo properties in the DMO and hydrodynamical simulations¹.

In this work, we focus exclusively on central galaxies² at $z = 0$ residing in haloes with the corresponding DMO mass $M_{200c}^{\text{DMO}} > 10^{11} M_{\odot}$. This selection yields a sample of $\gtrsim 77,000$ haloes. These haloes and their central galaxies are both well-resolved with $\gtrsim 50,000$ and $\gtrsim 100$ dark matter and stellar particles respectively.

3 RESULTS

3.1 The stellar mass-halo mass relation and its dependence on halo concentration

We begin our analysis by examining the relationship between stellar mass and halo mass for central galaxies, and its dependence on halo concentration. Fig. 1 shows the SMHM relation for central galaxies at $z = 0$. Following Matthee et al. (2017), halo masses and concentrations are taken from the matched haloes in the corresponding DMO simulation to better isolate the causal relationship, as the dark matter distribution in the hydrodynamical runs is influenced by baryonic processes (e.g. Duffy et al. 2010; Schaller et al. 2015). As expected, the median stellar mass increases monotonically across bins of increasing halo mass, though individual haloes exhibit noticeable scatter from the median relation, particularly in the low-mass regime. We compute the median and 16th–84th percentile range of stellar mass in bins of 0.1 dex in halo mass, and fit the median relation with equal weighting in 0.1-dex halo mass bins from 10^{11} to $10^{13.5} M_{\odot}$ using an empirical function (Matthee et al. 2017):

$$\log_{10} \left(\frac{M_{\text{star}}}{M_{\odot}} \right) = 11.70 - \exp \left[-0.88 \log_{10} \left(\frac{M_{200c}^{\text{DMO}}}{M_{\odot}} \right) - 10.85 \right], \quad (1)$$

which provides a good description of the overall trend across the full mass range considered.

Overlaid on the SMHM relation, we use colour to encode the concentration of the matched DMO halo for each galaxy. There is a moderately strong secondary dependence of stellar mass on halo concentration: at fixed halo mass, central galaxies residing in higher-concentration haloes exhibit systematically higher stellar masses than their lower-concentration counterparts. This concentration dependence is most pronounced in low-mass haloes. To quantify this effect, the bottom panel of Fig. 1 shows Spearman’s rank correlation coefficient³, $\mathcal{R}_s$, between stellar mass and DMO halo concentration

as a function of halo mass with a bin size of 0.1 dex. We also present three scatter plots between halo concentration and stellar mass in three selected halo mass bins in the smaller inset panels of Fig. 1 for a visual impression of the correlation strength. We find a moderately strong and statistically significant correlation in low-mass haloes, with coefficients of $\mathcal{R}_s \approx 0.50$ in the $M_{200c}^{\text{DMO}} = 10^{11.0-11.1} M_{\odot}$ bin and $\mathcal{R}_s \approx 0.46$ in the $M_{200c}^{\text{DMO}} = 10^{11.5-11.6} M_{\odot}$ bin. The correlation becomes weaker at higher halo mass, with a value of $\mathcal{R}_s \approx 0.15 - 0.35$ above $10^{12} M_{\odot}$, though it remains positive. These results are consistent with the findings of Matthee et al. (2017) for the EAGLE simulation (Schaye et al. 2015).

3.2 Disentangling the physical origin of the concentration dependence

The correlation between stellar mass and halo concentration at fixed halo mass shown in Fig. 1 invites an investigation into its underlying cause. Matthee et al. (2017) proposed two possible explanations for an analogous relation seen in EAGLE. The first takes halo concentration as a proxy for halo formation time since early-forming haloes are typically more concentrated (Navarro et al. 1997; Wechsler et al. 2002; Zhao et al. 2009; Ludlow et al. 2014), and thus have more time to convert gas into stars. The second is based on the fact that higher-concentration haloes have deeper gravitational potentials, which reduces the effectiveness of stellar feedback and allows for more efficient retention of baryons, compared to lower-concentration haloes.

To distinguish these two scenarios, we analyze how halo concentration correlates with stellar age and stellar metallicity, and how both properties relate to stellar mass. The motivation is that stellar age serves as a proxy for star formation history, so that if earlier halo assembly time results in higher stellar mass, we expect, at fixed halo mass, the correlation between halo concentration and stellar age with a strength similar to the correlation between concentration and stellar mass. In contrast, stellar metallicity reflects the integrated effect of metal production and retention, so that in haloes with deeper gravitational potentials, i.e. higher concentrations, stellar feedback is less efficient at ejecting metal-enriched gas, leading to higher metallicities and higher stellar mass (e.g. Peng & Maiolino 2014; Ma et al. 2024; Jia et al. 2025). Nevertheless, at fixed halo mass, a positive correlation between stellar mass and stellar metallicity could alternatively be attributed to more metal production during the star formation process. However, we will shortly show that stellar metallicity correlates moderately with halo concentration even when we control for both halo mass and stellar mass.

We are mainly interested in the halo mass range from 10^{11} to $10^{12} M_{\odot}$. The lower limit is due to the mass resolution of the simulation, as we require at least ≈ 100 stellar particles to minimally resolve the stellar population. The upper limit is chosen to be $10^{12} M_{\odot}$ for two reasons. Firstly, the correlation between stellar mass and halo concentration above this halo mass becomes very weak (see Fig. 1 and also Matthee et al. 2017; Pei et al. 2024). Secondly, AGN feedback starts to play a major role in the evolution of central galaxies above this halo mass (e.g. Bower et al. 2017). Previous studies found that the growth of super-massive black holes correlates with halo concentration, which is also attributed to the gravitational potential depth (Booth & Schaye 2010, 2011). Meanwhile, AGN feedback can suppress star formation activity (e.g. Bower et al. 2006; Di Matteo et al. 2008) and expel the metal-rich gas from the galaxy center and reduce the global metallicity (e.g. De Rossi et al. 2017; Wang et al. 2023b). The additional physical

¹ The L200m6 DMO run experienced a disk failure at $z \approx 0.22$ and was restarted from the most recent snapshot. This restart introduced minor discontinuities in the time integration of some particle trajectories, causing a few massive haloes in the hydrodynamical run to be spuriously matched to low-mass haloes in the DMO run. To remove these artefacts, we exclude all matched haloes with mass differences exceeding 0.3 dex. After this filtering, the bijective matching success rate remains $\gtrsim 98\%$ for haloes with $10^{11-12} M_{\odot}$, and we verified that this has no impact on the results presented in this paper.

² As detailed in Forouhar Moreno et al. (2025), all HBT-HERONS subhaloes are formed initially as centrals. During the subsequent hierarchical merging, the central subhalo is selected to be the one whose bound mass is at least 80% of the most massive progenitor central subhalo, and, if multiple subhaloes survive this threshold, the one with the lowest specific orbital kinetic energy in the frame of the host FoF halo is the new central.

³ Spearman’s rank correlation coefficient quantifies the rank correlation between two variables and ranges from 0 for no correlation to +1/−1 for a strictly monotonic relation.

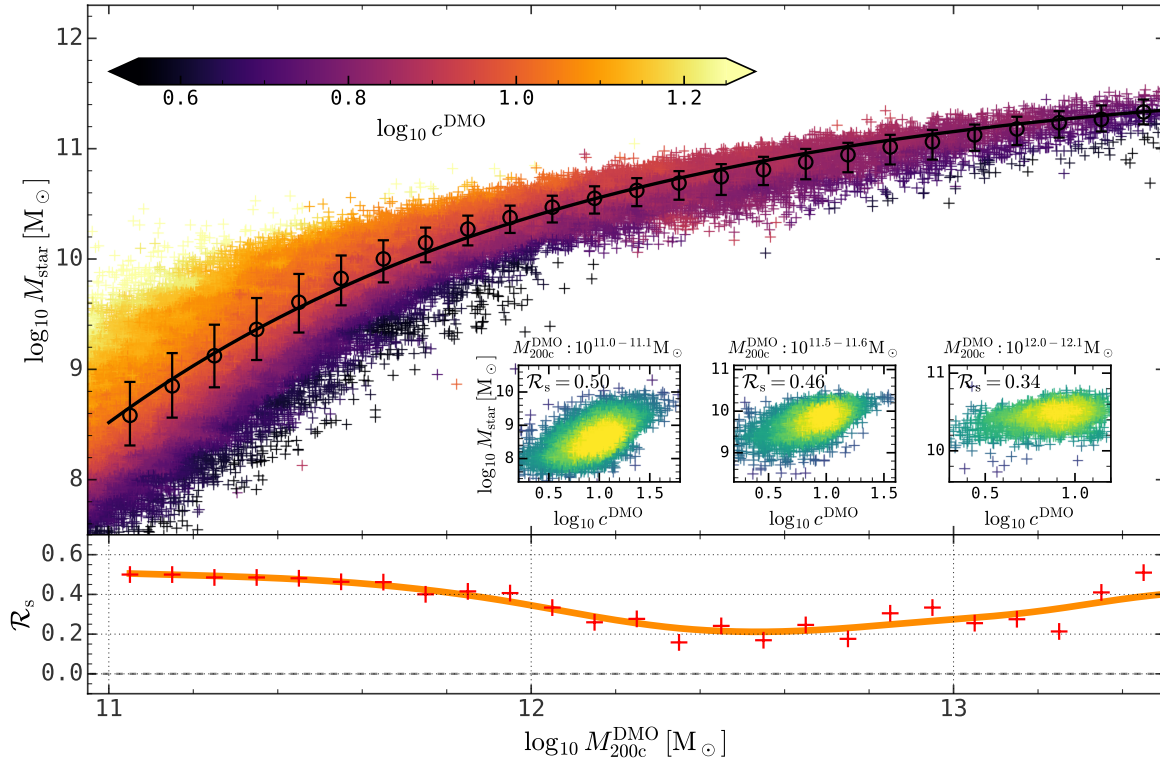


Figure 1. The SMHM relation for central galaxies in the COLIBRE L200m6 simulation, where colour encodes the logarithm of the halo concentration smoothed over 0.1-dex stellar mass and halo mass bins. Here both halo mass and halo concentration come from the corresponding DMO run. The error bars show the median and the 16th – 84th stellar mass quantiles within 0.1-dex-width halo mass bins. The solid line shows the best-fitting result with the functional form in equation (1). Three inset panels show the scatter plot, as well as Spearman’s rank correlation coefficient, between stellar mass and halo concentration in three selected 0.1-dex-width halo mass bins, where the colour encodes the density of the scatter point distribution. The red crosses in the bottom panel show Spearman’s rank correlation coefficients between stellar mass and halo concentration in 0.1-dex-width halo mass bins, where only bins with more than 30 data points are shown. The solid line is a smoothing B-spline line fitted to the data points to show the trend. There is a moderately strong correlation between stellar mass and concentration at fixed halo mass, particularly below $10^{12} M_{\odot}$, and the strength declines above this mass.

processes involved in massive haloes complicate the interpretation, so we will leave this regime to future work.

The correlations between halo concentration, stellar mass, stellar metallicity, and stellar age are summarized in Fig. 2, which shows pairwise Spearman’s rank correlation coefficients between relevant properties measured in 0.1-dex halo mass bins across $10^{11} - 10^{13} M_{\odot}$. For the low-mass regime from 10^{11} to $10^{12} M_{\odot}$, the top two panels reveal a moderately strong positive correlation between halo concentration and stellar age that increases with halo mass from $\mathcal{R}_s \approx 0.40$ to ≈ 0.60 , but only a weak correlation between stellar age and stellar mass, ranging from $\mathcal{R}_s \approx -0.2$ to ≈ 0.2 , which is even negative below $10^{11.2} M_{\odot}$ (see Pei et al. 2024, for consistent results in other simulations and semi-analytical models). This suggests that while galaxies in higher-concentration haloes formed earlier, this does not by itself lead to systematically higher stellar masses, falsifying the early-formation hypothesis.

The bottom two panels of Fig. 2 show moderately strong and coherent trends involving stellar metallicity for halo masses between 10^{11} and $10^{12} M_{\odot}$. The rank correlation coefficient between halo concentration and stellar metallicity ranges from $\mathcal{R}_s \approx 0.60$ to ≈ 0.30 , and the stellar metallicity-stellar mass correlation remains consistently between $\mathcal{R}_s \approx 0.70$ and ≈ 0.60 across the halo mass range $10^{11} - 10^{12} M_{\odot}$. These aligned trends suggest that the deeper gravitational potentials of higher-concentration haloes

enhance metal retention and reduce the impact of feedback, enabling more efficient star formation and, consequently, higher stellar masses. The behavior changes in haloes above $10^{12} M_{\odot}$ due to the involvement of AGN feedback; we leave this for future work.

Although the results in Fig. 2 are consistent with the potential-driven scenario, it remains possible that the observed correlation between halo concentration and stellar metallicity arises indirectly through their mutual dependence on stellar mass. To address this possibility, Fig. 3 shows Spearman’s rank correlation coefficient for the relation between stellar metallicity and halo concentration while simultaneously controlling for both stellar mass and halo mass by binning central galaxies in narrow intervals of stellar mass (0.1 dex) and halo mass (0.3 dex). As a consistency check, we also measure the correlation between stellar mass and concentration within the same bins to verify that stellar mass is effectively held fixed.

Fig. 3 shows that the correlation between stellar mass and concentration disappears under these controls, confirming the successful mitigation of stellar mass effects. In contrast, the concentration-stellar metallicity correlation remains moderately strong with a value of $\mathcal{R}_s \approx 0.34$ for haloes below $10^{12} M_{\odot}$. This confirms that the stellar metallicity-concentration relation is not a byproduct of the correlation between stellar metallicity and stellar mass, and between halo concentration and stellar mass. Instead, this indicates a direct connection between halo potential depth and metal reten-

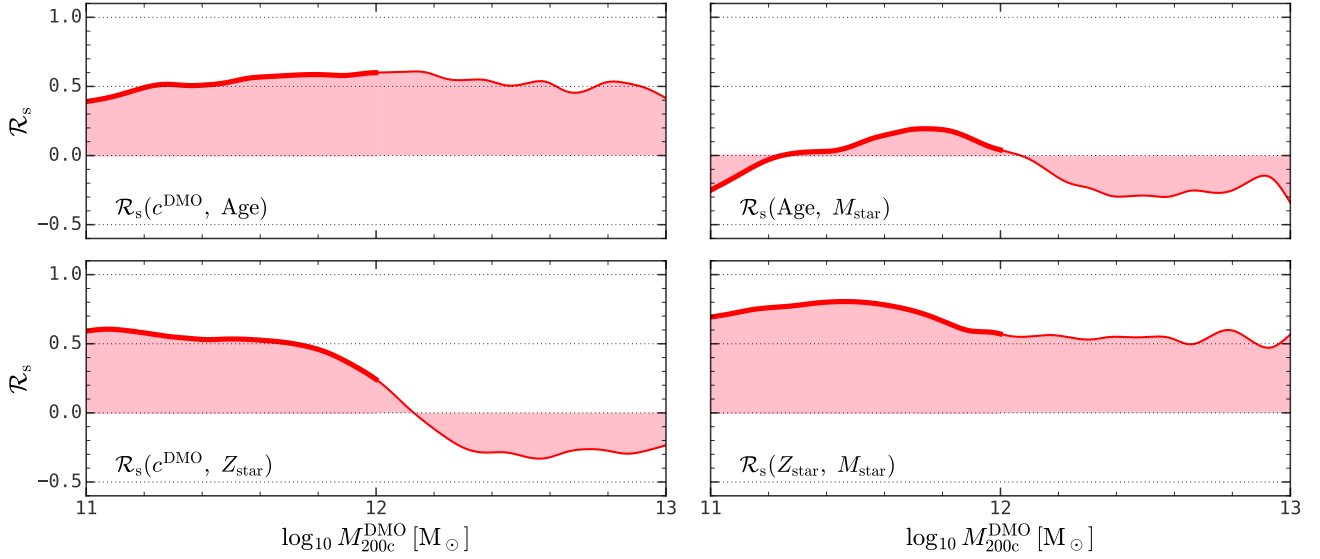


Figure 2. Pairwise Spearman’s rank correlation coefficients for relations between halo and central galaxy properties within 0.1-dex-width halo mass bins from 10^{11} to $10^{13} M_{\odot}$. The four panels correspond to correlations between halo concentration and stellar age, stellar age and stellar mass, halo concentration and stellar metallicity, and stellar metallicity and stellar mass, respectively. Here both stellar age and stellar metallicity are weighted by the mass of stellar particles, and all halo related properties are from the matched haloes in the DMO run. In the halo mass range of $10^{11-12} M_{\odot}$, there is at most a weak correlation between stellar age and stellar mass while all other correlations are of moderate strength, so the correlation between halo concentration and stellar mass is mediated by stellar metallicity rather than by stellar age.

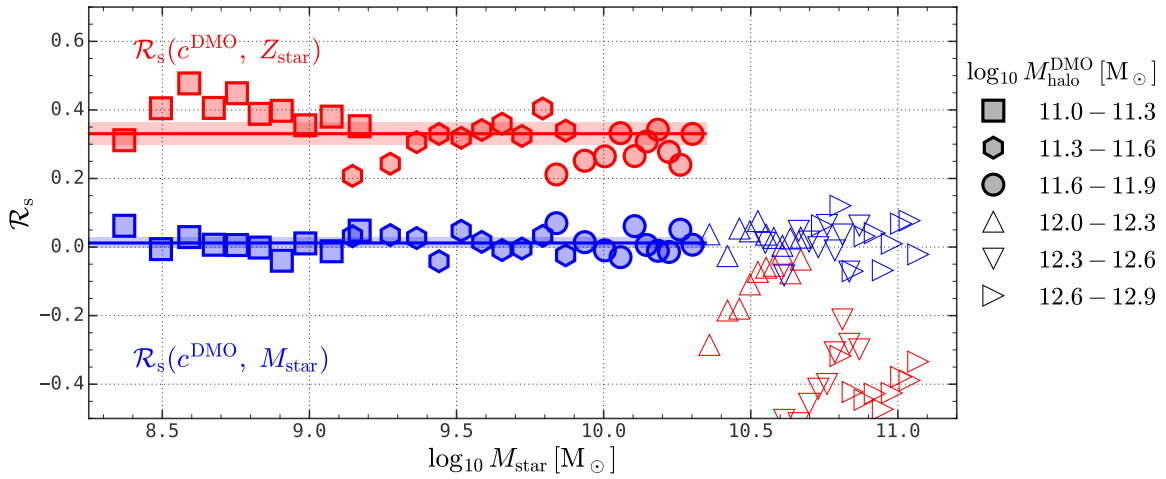


Figure 3. Spearman’s rank correlation coefficients between halo concentration and stellar metallicity (red symbols), and between halo concentration and stellar mass (blue symbols) at fixed stellar and halo mass. Results for haloes within the mass ranges $10^{11-12} M_{\odot}$ and $10^{12-13} M_{\odot}$ are shown with filled and open symbols, respectively. Stellar mass and halo mass are constrained within 0.1-dex and 0.3-dex bins, respectively. The two horizontal lines and the shaded regions are the mean values and standard deviations of the correlation coefficients for haloes of mass $10^{11-12} M_{\odot}$. The correlation coefficient is only evaluated when more than 30 galaxies are present in the bin. All halo related quantities come from the matched haloes in the DMO run. As expected, there is no correlation between halo concentration and stellar mass when both halo mass and stellar mass are controlled. However, there is still a moderately strong correlation between concentration and stellar metallicity. The fact that this correlation is present at fixed stellar mass implies that it is due to the potential depth rather than the amount of metals produced.

tion efficiency. These results thus reinforce the interpretation that deeper gravitational potentials in higher-concentration haloes reduce feedback-driven outflows, enhancing metal retention and leading to higher stellar masses. Again, in haloes with masses above $10^{12} M_{\odot}$, AGN feedback becomes more important than stellar feedback and changes the correlation between halo concentration and stellar metallicity.

4 DISCUSSION AND SUMMARY

In this work, we investigate the origin of the scatter in the SMHM relation using the L200m6 COLIBRE cosmological hydrodynamical simulation. Our analysis focuses on the role of halo concentration as a secondary parameter, and confirms the findings of [Matthee et al. \(2017\)](#) that halo concentration is moderately correlated with stellar mass at fixed halo mass, particularly for haloes in the mass range $10^{11} - 10^{12} M_{\odot}$ (see Fig. 1). We examine whether this correlation is primarily driven by earlier halo formation time or by deeper gravitational potential resulting in less efficient feedback.

We compare the pairwise correlations among halo concentration, stellar age, stellar metallicity, and stellar mass, all with halo mass fixed in 0.1-dex mass bins. We find that, for haloes of mass between 10^{11} and $10^{12} M_{\odot}$, although stellar age correlates moderately with halo concentration, it correlates only weakly with stellar mass. In contrast, stellar metallicity exhibits moderately strong correlations with both halo concentration and stellar mass (see Fig. 2). This suggests that the deeper potential wells of higher-concentration haloes enhance metal retention and star formation efficiency, and give rise to a correlation between halo concentration and galaxy stellar mass, rather than early formation time being responsible.

Since a higher stellar mass could directly cause a higher metallicity through the larger amount of metals produced, we further test our interpretation by computing the stellar metallicity-concentration correlation while simultaneously controlling for both stellar mass and halo mass. The correlation remains significant with a Spearman's rank correlation coefficient of ≈ 0.34 (see Fig. 3). This confirms that the concentration-stellar metallicity link is not mediated through stellar mass and supports a direct and important role of halo potential depth in shaping galaxy properties at fixed halo mass.

We verify the robustness of our results using the COLIBRE runs with hybrid AGN feedback (see [Huško et al. 2025](#)), as well as simulations at different resolutions (see [Schaye et al. 2025](#)). The m5 and m7 runs have average baryonic particle masses of $2.30 \times 10^5 M_{\odot}$ and $1.47 \times 10^7 M_{\odot}$, and the corresponding dark matter particle masses are $3.03 \times 10^5 M_{\odot}$ and $1.94 \times 10^7 M_{\odot}$, respectively. The correlation trends obtained from all these runs are consistent with those in the L200m6 simulation, confirming that our conclusions are insensitive to the adopted AGN feedback model and to numerical resolution.

We recognise that stellar age and metallicity are correlated, as the stellar metallicity is shaped by the star formation history. However, it is unlikely that, at fixed halo mass, the stellar mass-metallicity correlation is driven by stellar age, given that stellar age exhibits only a weak correlation with stellar mass (see Fig. 2). Therefore, the age-metallicity relation does not compromise our main conclusion regarding the role of halo concentration and its effects on gravitational potential in shaping stellar metallicity and mass. Our results therefore provide an appealing explanation for the physical origin of the scatter in the fundamental relation between central galaxy stellar masses and the masses of the haloes in which they reside.

We also find that haloes with masses above $10^{12} M_{\odot}$ exhibit different behavior compared to lower-mass ones. In this high-mass regime, AGN feedback plays a more dominant role than stellar feedback in regulating galaxy evolution. A detailed investigation of this mass range is beyond the scope of the current work and will be explored in future studies.

Finally, our results offer promising avenues for observational validation, as stellar metallicity can be derived from spectroscopic observations and halo mass and concentration can be inferred through gravitational lensing and galaxy kinematics. This makes the predicted connection between stellar metallicity and halo potential depth directly testable with current and upcoming observations.

DATA AVAILABILITY

The data supporting the plots in this article are available on reasonable request to the corresponding author. The COLIBRE simulation data will eventually be made publicly available, although we note that the data volume (several petabytes) may prevent us from simply placing the raw data on a server.

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REFERENCES

- Abbott T. M. C., et al., 2022, *Physical Review D*, 105, 023520
- Artale M. C., Zehavi I., Contreras S., Norberg P., 2018, *Monthly Notices of the Royal Astronomical Society*, 480, 3978
- Behroozi P., Wechsler R. H., Hearin A. P., Conroy C., 2019, *Monthly Notices of the Royal Astronomical Society*, 488, 3143
- Benítez-Llambay A., et al., 2025, *arXiv e-prints*, p. [arXiv:2509.25309](https://arxiv.org/abs/2509.25309)

- Blumenthal G. R., Faber S. M., Flores R., Primack J. R., 1986, *The Astrophysical Journal*, 301, 27
- Booth C. M., Schaye J., 2009, *Monthly Notices of the Royal Astronomical Society*, 398, 53
- Booth C. M., Schaye J., 2010, *Monthly Notices of the Royal Astronomical Society*, 405, L1
- Booth C. M., Schaye J., 2011, *Monthly Notices of the Royal Astronomical Society*, 413, 1158
- Borrow J., Borisov A., 2020, *The Journal of Open Source Software*, 5, 2430
- Bose S., Eisenstein D. J., Hernquist L., Pillepich A., Nelson D., Marinacci F., Springel V., Vogelsberger M., 2019, *Monthly Notices of the Royal Astronomical Society*, 490, 5693
- Bower R. G., Benson A. J., Malbon R., Helly J. C., Frenk C. S., Baugh C. M., Cole S., Lacey C. G., 2006, *Monthly Notices of the Royal Astronomical Society*, 370, 645
- Bower R. G., Schaye J., Frenk C. S., Theuns T., Schaller M., Crain R. A., McAlpine S., 2017, *Monthly Notices of the Royal Astronomical Society*, 465, 32
- Bradshaw C., Leauthaud A., Hearin A., Huang S., Behroozi P., 2020, *Monthly Notices of the Royal Astronomical Society*, 493, 337
- Chaikin E., Schaye J., Schaller M., Benítez-Llambay A., Nobels F. S. J., Ploekinger S., 2023, *Monthly Notices of the Royal Astronomical Society*, 523, 3709
- Chaikin E., et al., 2025a, *arXiv e-prints*, p. [arXiv:2509.04067](#)
- Chaikin E., et al., 2025b, *arXiv e-prints*, p. [arXiv:2509.07960](#)
- Chen Y., Wang K., 2023, HIPP: High-Performance Package for scientific computation, Astrophysics Source Code Library, record ascl:2301.030 (ascl:2301.030)
- Correa C. A., Schaye J., 2020, *Monthly Notices of the Royal Astronomical Society*, 499, 3578
- Correa C. A., Wyithe J. S. B., Schaye J., Duffy A. R., 2015, *Monthly Notices of the Royal Astronomical Society*, 452, 1217
- Correa C., Schaye J., Chaikin E. e. a., 2025, *Monthly Notices of the Royal Astronomical Society*
- Cui W., Davé R., Peacock J. A., Anglés-Alcázar D., Yang X., 2021, *Nature Astronomy*, 5, 1069
- De Rossi M. E., Bower R. G., Font A. S., Schaye J., Theuns T., 2017, *Monthly Notices of the Royal Astronomical Society*, 472, 3354
- Di Matteo T., Colberg J., Springel V., Hernquist L., Sijacki D., 2008, *The Astrophysical Journal*, 676, 33
- Duffy A. R., Schaye J., Kay S. T., Dalla Vecchia C., Battye R. A., Booth C. M., 2010, *Monthly Notices of the Royal Astronomical Society*, 405, 2161
- Eke V. R., et al., 2004, *Monthly Notices of the Royal Astronomical Society*, 355, 769
- Forouhar Moreno V. J., Helly J., McGibbon R., Schaye J., Schaller M., Han J., Kugel R., Bahé Y. M., 2025, *MNRAS*, 543, 1339
- Gu M., Conroy C., Behroozi P., 2016, *The Astrophysical Journal*, 833, 2
- Huško F., et al., 2025, *arXiv e-prints*, p. [arXiv:2509.05179](#)
- Jia C., et al., 2025, *The Astrophysical Journal*, 986, L24
- Kravtsov A. V., Vikhlinin A. A., Meshcheryakov A. V., 2018, *Astronomy Letters*, 44, 8
- Kulier A., Padilla N., Schaye J., Crain R. A., Schaller M., Bower R. G., Theuns T., Paillas E., 2019, *Monthly Notices of the Royal Astronomical Society*, 482, 3261
- Ludlow A. D., Navarro J. F., Angulo R. E., Boylan-Kolchin M., Springel V., Frenk C., White S. D. M., 2014, *Monthly Notices of the Royal Astronomical Society*, 441, 378
- Ludlow A. D., Fall S. M., Wilkinson M. J., Schaye J., Obreschkow D., 2023, *Monthly Notices of the Royal Astronomical Society*, 525, 5614
- Lyu C., et al., 2023, *The Astrophysical Journal*, 959, 5
- Ma C., et al., 2024, *The Astrophysical Journal*, 971, L14
- Mandelbaum R., Seljak U., Kauffmann G., Hirata C. M., Brinkmann J., 2006, *Monthly Notices of the Royal Astronomical Society*, 368, 715
- Martizzi D., Vogelsberger M., Torrey P., Pillepich A., Hansen S. H., Marinacci F., Hernquist L., 2020, *Monthly Notices of the Royal Astronomical Society*, 491, 5747
- Matthee J., Schaye J., Crain R. A., Schaller M., Bower R., Theuns T., 2017, *Monthly Notices of the Royal Astronomical Society*, 465, 2381
- McGibbon R., Helly J., Schaye J., Schaller M., Vandenbroucke B., 2025, *The Journal of Open Source Software*, 10, 8252
- Montero-Dorta A. D., et al., 2020, *Monthly Notices of the Royal Astronomical Society*, 496, 1182
- More S., van den Bosch F. C., Cacciato M., 2009, *Monthly Notices of the Royal Astronomical Society*, 392, 917
- Moster B. P., Somerville R. S., Maulbetsch C., van den Bosch F. C., Macciò A. V., Naab T., Oser L., 2010, *The Astrophysical Journal*, 710, 903
- Navarro J. F., Frenk C. S., White S. D. M., 1997, *The Astrophysical Journal*, 490, 493
- Nobels F. S. J., Schaye J., Schaller M., Ploekinger S., Chaikin E., Richings A. J., 2024, *Monthly Notices of the Royal Astronomical Society*, 532, 3299
- Oman K., 2025, *The Journal of Open Source Software*, 10, 9278
- Pei W., Guo Q., Shao S., He Y., Gu Q., 2024, *Monthly Notices of the Royal Astronomical Society*
- Peng Y.-j., Maiolino R., 2014, *Monthly Notices of the Royal Astronomical Society*, 443, 3643
- Ploekinger S., Richings A. J., Schaye J., Trayford J. W., Schaller M., Chaikin E., 2025, *Monthly Notices of the Royal Astronomical Society*
- Reddick R. M., Wechsler R. H., Tinker J. L., Behroozi P. S., 2013, *The Astrophysical Journal*, 771, 30
- Schaller M., et al., 2015, *Monthly Notices of the Royal Astronomical Society*, 451, 1247
- Schaller M., et al., 2024, *Monthly Notices of the Royal Astronomical Society*, 530, 2378
- Schaye J., et al., 2015, *Monthly Notices of the Royal Astronomical Society*, 446, 521
- Schaye J., et al., 2025, *arXiv e-prints*, p. [arXiv:2508.21126](#)
- Scholz-Díaz L., Martín-Navarro I., Falcón-Barroso J., 2022, *Monthly Notices of the Royal Astronomical Society*, 511, 4900
- Scholz-Díaz L., Martín-Navarro I., Falcón-Barroso J., Lyubenova M., van de Ven G., 2024, *Nature Astronomy*, 8, 648
- Tojeiro R., et al., 2017, *Monthly Notices of the Royal Astronomical Society*, 470, 3720
- Trayford J. W., et al., 2025, *arXiv e-prints*, p. [arXiv:2505.13056](#)
- Wang K., 2025, *arXiv e-prints*, p. [arXiv:2510.02573](#)
- Wang K., Peng Y., 2025, *The Astrophysical Journal*, 980, 233
- Wang L., De Lucia G., Weinmann S. M., 2013, *Monthly Notices of the Royal Astronomical Society*, 431, 600
- Wang K., Chen Y., Li Q., Yang X., 2023a, *Monthly Notices of the Royal Astronomical Society*, 522, 3188
- Wang K., Wang X., Chen Y., 2023b, *The Astrophysical Journal*, 951, 66
- Wang K., Mo H. J., Chen Y., Schaye J., 2024, *Monthly Notices of the Royal Astronomical Society*, 527, 10760
- Wechsler R. H., Tinker J. L., 2018, *Annual Review of Astronomy and Astrophysics*, 56, 435
- Wechsler R. H., Bullock J. S., Primack J. R., Kravtsov A. V., Dekel A., 2002, *The Astrophysical Journal*, 568, 52
- White S. D. M., Frenk C. S., 1991, *The Astrophysical Journal*, 379, 52
- Yang X., Mo H. J., van den Bosch F. C., 2003, *Monthly Notices of the Royal Astronomical Society*, 339, 1057
- Yang X., Mo H. J., van den Bosch F. C., 2009, *The Astrophysical Journal*, 695, 900
- Zehavi I., Contreras S., Padilla N., Smith N. J., Baugh C. M., Norberg P., 2018, *The Astrophysical Journal*, 853, 84
- Zehavi I., Kerby S. E., Contreras S., Jiménez E., Padilla N., Baugh C. M., 2019, *The Astrophysical Journal*, 887, 17
- Zhao D. H., Jing Y. P., Mo H. J., Börner G., 2009, *The Astrophysical Journal*, 707, 354
- Zhao D., et al., 2025, *The Astrophysical Journal*, 979, 42
- Zu Y., Mandelbaum R., 2015, *Monthly Notices of the Royal Astronomical Society*, 454, 1161
- Zu Y., et al., 2021, *Monthly Notices of the Royal Astronomical Society*, 505, 5117

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